

## THE *CHANDRA* X-RAY SURVEY OF PLANETARY NEBULAE (CHANPLANS): PROBING BINARITY, MAGNETIC FIELDS, AND WIND COLLISIONS

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### ABSTRACT

We present an overview of the initial results from the *Chandra* Planetary Nebula Survey (CHANPLANS), the first systematic (volume-limited) *Chandra X-Ray Observatory* survey of planetary nebulae (PNe) in the solar neighborhood. The first phase of CHANPLANS targeted 21 mostly high-excitation PNe within  $\sim 1.5$  kpc of Earth, yielding four detections of diffuse X-ray emission and nine detections of X-ray-luminous point sources at the central stars (CSPNe) of these objects. Combining these results with those obtained from *Chandra* archival data for all (14) other PNe within  $\sim 1.5$  kpc that have been observed to date, we find an overall X-ray detection rate of  $\sim 70\%$  for the 35 sample objects. Roughly 50% of the PNe observed by *Chandra* harbor X-ray-luminous CSPNe, while soft, diffuse X-ray emission tracing shocks—in most cases, “hot bubbles”—formed by energetic wind collisions is detected in  $\sim 30\%$ ; five objects display both diffuse and point-like emission components. The presence (or absence) of X-ray sources appears correlated with PN density structure, in that molecule-poor, elliptical nebulae are more likely to display X-ray emission (either point-like or diffuse) than molecule-rich, bipolar, or Ring-like nebulae. All but one of the point-like CSPNe X-ray sources display X-ray spectra that are harder than expected from hot ( $\sim 100$  kK) central stars emitting as simple blackbodies; the lone apparent exception is the central star of the Dumbbell nebula, NGC 6853. These hard X-ray excesses may suggest a high frequency of binary companions to CSPNe. Other potential explanations include self-shocking winds or PN mass fallback. Most PNe detected as diffuse X-ray sources are elliptical nebulae that display a nested shell/halo structure and bright ansae; the diffuse X-ray emission regions are confined within inner, sharp-rimmed shells. All sample PNe that display diffuse X-ray emission have inner shell dynamical ages  $\lesssim 5 \times 10^3$  yr, placing firm constraints on the timescale for strong shocks due to wind interactions in PNe. The high-energy emission arising in such wind shocks may contribute to the high excitation states of certain archetypical “hot bubble” nebulae (e.g., NGC 2392, 3242, 6826, and 7009).

**Key words:** binaries: general – planetary nebulae: general – stars: AGB and post-AGB

*Online-only material:* color figures

### 1. INTRODUCTION

Planetary nebulae (PNe), the near endpoints of stellar evolution for intermediate-mass ( $\sim 1\text{--}8 M_{\odot}$ ) stars, have served

as astrophysical laboratories for more than a century (Aller 1956). With their relatively large numbers in close proximity ( $\sim 200$  PNe lie within  $\sim 2$  kpc; Cahn et al. 1992; Frew 2008; Stanghellini et al. 2008), PNe serve as primary examples of plasma and shock processes and provide essential tests of theories of stellar evolution and the origin and enrichment of the

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heavy elements in the universe (Kwok 2000). PNe are best known as  $\sim 10^4$  K optical emission line sources, yet many objects harbor cold ( $< 100$  K), dense ( $\sim 10^6$ – $10^7$  cm $^{-3}$ ) gas, and dust, and some of these same PNe also display emission from rarefied, hot ( $T > 10^6$  K), X-ray-emitting plasma. In principle, each of these temperature and density regimes informs us about the properties of the progenitor star system and its evolution.

Long thought to signify the transition of single stars from asymptotic giant branch (AGB) to white dwarf (WD) evolutionary stages, PNe exhibit a dazzling variety of optical and near-infrared morphologies: round; elliptical; bipolar; highly point symmetric; chaotic; and clumpy (e.g., Sahai et al. 2011 and references therein). The physical mechanisms responsible for this PN morphological menagerie—and, in particular, for the evident transformation from a quasi-isotropic wind during the progenitor star AGB phase to nonspherical or even highly collimated outflow during the PN phase—have been the subject of intense interest and hot debate among PN researchers over the past two decades (e.g., Balick & Frank 2002; Zijlstra et al. 2011; Kastner 2011). At the heart of this debate lies the question: Do some, most, or even all PNe actually represent the ejection and photoionization of the envelopes of AGB stars in *binary systems*? While stellar magnetic fields generated within the AGB progenitor may play a role in the early shaping of PNe (Blackman et al. 2001a and references therein), the large fraction of PNe that display nonspherical geometry (Soker 1997; Sahai & Trauger 1998)—and the fact that at least  $\sim 20\%$  of PNe are known to have binary companions to their central stars (Bond 2000; Miszalski et al. 2009a)—indeed suggests that a significant fraction of PNe represent the products of interacting binary star systems, within which preferred symmetry axes are found (e.g., De Marco 2009; Miszalski et al. 2009b; Jones et al. 2012 and references therein).

A binary companion to the PN central star (hereafter CSPN) can influence the PN shape in various ways, e.g., via formation of an accretion disk around the secondary (e.g., Morris 1987; Mastrodemos & Morris 1998; Soker & Rappaport 2000) or angular momentum injection (during, e.g., a common envelope phase) and the consequent generation of a disk and/or a strong magnetic dynamo at the primary (e.g., Reyes-Ruiz & López 1999; Nordhaus & Blackman 2006; Nordhaus et al. 2007). These models have recently received strong observational support, in the form of examples of close binary CSPNe that drive jets (e.g., Corradi et al. 2011; Miszalski et al. 2011). The potential role of disks as agents of PN outflow collimation, via jet formation (e.g., Soker & Livio 1994; Blackman et al. 2001b), invites analogies to the disk/jet connection in, e.g., young stellar objects (Kastner 2009). There also appears to be a close link between symbiotic binary systems and both “butterfly” (bipolar) PNe (Corradi & Schwarz 1995) and high-[O III]-luminosity PNe (Frankowski & Soker 2009); the latter class of PN serves as an extragalactic standard candle (Ciardullo et al. 2005).

Standard models for the formation of PNe (e.g., Kwok et al. 1978; Schmidt-Voigt & Koepfen 1987; Marten & Schönberner 1991; Villaver et al. 2002; Perinotto et al. 2004) predict that the fast ( $v_w \simeq 500$ – $1500$  km s $^{-1}$ ) wind emanating from the pre-WD at the core of the PN rams into the previously expelled AGB envelope (which was ejected at  $\sim 10$  km s $^{-1}$ ), thereby sweeping the AGB ejecta into a thin shell and shocking the fast wind to temperatures  $\gg 10^6$  K. Such wind interaction models therefore predict that PNe should harbor X-ray-luminous, evacuated bubbles (e.g., Zhekov & Perinotto 1996; Akashi et al. 2006; Steffen et al. 2008; Lou & Zhai 2010). Over the past decade,

X-ray imaging by *Chandra* and *XMM-Newton* has provided compelling observational evidence for such CSPN-wind-blown “hot bubbles” (Kastner et al. 2008, and references therein). About a dozen PNe previously targeted by the two contemporary X-ray observatories have been detected as diffuse X-ray sources (Kastner et al. 2000, 2001, 2003; Chu et al. 2001; Guerrero et al. 2002, 2005; Montez et al. 2005; Gruendl et al. 2006).

*Chandra* imaging has also revealed that certain PNe harbor X-ray point sources at their cores, with source X-ray spectral energy distributions (SEDs) that cannot be explained as the Wien tails of CSPNe emitting as simple hot blackbodies (e.g., Guerrero et al. 2001; Kastner et al. 2003; Montez et al. 2010). Notably—when imaged by *Chandra*—a few PNe reveal both soft, diffuse and harder, point-like X-ray emission (e.g., NGC 6543; Chu et al. 2001). These two “flavors” of PN X-ray sources—diffuse and point-like—serve as complementary probes of the mechanisms underlying PN structural evolution.

*Diffuse X-ray sources.* Several trends have emerged from the examples of diffuse PN X-ray sources detected to date by *Chandra* and *XMM* (Kastner 2007; Kastner et al. 2008 and references therein).

1. Those PNe in which X-ray-luminous “hot bubbles” have been detected thus far all harbor CSPNe that drive particularly energetic winds (speeds  $V_w \sim 1000$  km s $^{-1}$  at mass-loss rates  $\dot{M} \gtrsim$  a few  $\times 10^{-8}$   $M_\odot$  yr $^{-1}$ ). A disproportionate fraction of these CSPNe are of the relatively rare Wolf-Rayet (WR) type (with  $\dot{M} \gtrsim 10^{-6}$   $M_\odot$  yr $^{-1}$ ).
2. Hot bubble X-ray luminosity seems to be weakly correlated with present-day CSPN wind luminosity  $L_w = (1/2)\dot{M}v_w^2$  and anticorrelated with bubble radius, indicative of the close connection between the evolution of CSPN winds and PN hot bubbles (a connection explored in various theoretical investigations; e.g., Akashi et al. 2006, 2007; Steffen et al. 2008).
3. In cases in which hot bubble X-ray emission is detected, the optical/IR structures that enclose the regions of diffuse X-rays have thin, bright, uninterrupted edges, suggesting that the diffuse X-ray-emitting gas is spatially confined and therefore inhibited from expanding adiabatically.

Among the more surprising results obtained from X-ray imaging spectroscopy of PNe are the low temperatures of the shocked (X-ray-emitting) wind gas ( $T_X$ ) in PN hot bubbles. Diffuse X-ray emission regions within PNe typically display  $T_X$  in the narrow range  $\sim 1$ – $2$  MK, which is one to two orders of magnitude lower than expected, based on simple jump conditions, for central star wind speeds  $V_w \sim 1000$  km s $^{-1}$ ; furthermore, hot bubble  $T_X$  does not appear to depend on CSPN wind velocity (Kastner et al. 2008; Montez 2010). Many temperature regulation mechanisms have been proposed to explain these results (see Soker & Kastner 2002; Stute & Sahai 2006; Kastner et al. 2008; Soker et al. 2010 and references therein). The low observed values of  $T_X$  may indicate that hot bubble physical conditions are established during early phases of the post-AGB/pre-PN evolution of central stars with rapidly evolving winds (Akashi et al. 2006, 2007). Alternatively, an active temperature-moderating mechanism may govern the observed  $T_X$ . Potential mechanisms include heat conduction (Steffen et al. 2008; Li et al. 2012), mixing of nebular and fast wind material (Chu et al. 2001), or a small mass of “pickup ions” that wander into the hot bubble from cold, neutral nebular clumps (analogous to a mechanism proposed to cool the solar wind; see Soker et al. 2010). The heat conduction

models appear to hold particular promise (Steffen et al. 2008; Montez 2010; C. Sandin et al. 2012, in preparation). On the other hand, measurements of the elemental abundances within the X-ray-emitting plasma of BD +30°3639—the most luminous diffuse X-ray PN and, hence, the only object for which such precise hot bubble abundance and temperature determinations are presently available (via *Chandra* X-ray gratings spectroscopy; Yu et al. 2009)—closely match that of its [WC]-type central star. This strong resemblance led Yu et al. (2009) to conclude that the superheated plasma within BD +30°3639 consists of “pure” fast wind material, such that neither heat conduction nor mixing likely plays an important role in determining its (low) characteristic  $T_X$  of  $\sim 2 \times 10^6$  K.

*X-ray point sources at CSPNe.* *Einstein* and *ROSAT* established that certain high-excitation PNe harbor soft X-ray sources indicative of emission from hot ( $\gtrsim 100$  kK) CSPN photospheres (e.g., Motch et al. 1993; Guerrero et al. 2000). However, *Chandra* imaging has revealed intriguing examples of X-ray sources at CSPNe that are too hard to be modeled in terms of simple blackbody emission from a pre-WD stellar photosphere (Guerrero et al. 2001; Hoogerwerf et al. 2007; Montez et al. 2010). The Helix Nebula (NGC 7293) central star is perhaps the best-characterized example of such a “hard X-ray excess” source (Guerrero et al. 2001). Among the thousands of WDs that have been observed (mostly serendipitously) by *XMM* or *ROSAT*, only a handful of isolated, supposedly single WDs—including the post-PN object PG 1159—display similarly hard spectra (Bilíková et al. 2010). Via analogy with cataclysmic variables and symbiotic binaries, this new class of relatively hard PN X-ray point source may be hinting at the presence of binary companions and/or accretion processes associated with CSPNe (Kastner 2007). The hard X-rays may arise from accretion onto a compact, hot companion (Kastner et al. 2003), or from the corona of a late-type companion that has been “rejuvenated” via accretion of pre-PN (AGB) wind material (Jeffries & Stevens 1996; Soker & Kastner 2002), as appears to be the case for the PNe DS 1, HFG 1, LoTr 5 (all of which are known binaries; Montez et al. 2010), and K 1–6 (a nearby, binary CSPN that was detected by *ROSAT*; Frew et al. 2011). Other (non-binary) models also could explain the presence of “hard excess” point sources, however, such as internal wind shocks analogous to those observed in O stars (Guerrero et al. 2001) or mass infall, e.g., from a residual, Kuiper-Belt-like debris disk orbiting the CSPN (Su et al. 2007; Bilíková et al. 2010).

The X-ray emission characteristics of PNe just described have been assembled from piecemeal and uncoordinated *Chandra* programs, each of which targeted perhaps one or two objects. The resulting small number statistics, combined with the hazardous nature of the sample of PNe observed thus far in X-rays at *Chandra*’s subarcsecond spatial resolution—which is required to distinguish between point-like and diffuse emission—leaves many fundamental questions unanswered: under what circumstances do wind–wind shocks lead to hot bubbles within PNe, and how do these hot bubbles evolve with time? How are the kinematics of PNe and the wind properties of their central stars related to the luminosity and morphology of PN X-ray emission? What heating and cooling mechanisms govern the temperatures of the X-ray-emitting plasmas within PNe? How might detection and characterization of X-ray point sources at CSPNe improve our knowledge of the frequency and characteristics of binary systems within PNe, and the relationships of such binaries to potentially related systems such as symbiotic stars and SN Ia progenitor binaries?

To address these questions, we are undertaking the *Chandra* Planetary Nebula Survey (CHANPLANS)—a *Chandra* X-Ray Observatory survey of the  $\sim 120$  known PNe within  $\sim 1.5$  kpc of Earth (as drawn from the comprehensive catalogs compiled by Acker et al. 1992; Cahn et al. 1992; Frew 2008). CHANPLANS constitutes one element of a planned comprehensive observational and theoretical campaign to understand the shaping of PNe, as described in the “Rochester White Paper” (De Marco et al. 2011). The CHANPLANS survey began with a 570 ks *Chandra* Cycle 12 Large Program targeting 21 (mostly high-excitation) PNe from among this sample, thereby roughly doubling the number of PNe within  $\sim 1.5$  kpc that have been observed by *Chandra*. In this paper, we describe initial results obtained for this, the first statistically significant, volume-limited sample of PNe to be imaged in X-rays at high spatial resolution.

## 2. SAMPLE SELECTION, OBSERVATIONS, AND DATA REDUCTION

### 2.1. Planetary Nebulae within $\sim 1.5$ kpc Observed by *Chandra*

The sample of 21 PN targeted during *Chandra* Cycle 12 was assembled from the comprehensive lists of well-studied PNe in Gurzadian (1988), Acker et al. (1992), and Frew (2008). Most of the targeted PNe are “high-excitation” objects, characterized by bright lines of highly ionized species of, e.g., He, C, N, O, and Ne that are generally indicative of central stars with high effective temperatures (i.e.,  $T_{\text{eff}} \gtrsim 10^5$  K). We initially selected those PNe for which Gurzadian (1988) lists  $I(\lambda 4686)/I(\text{H}\beta) \gtrsim 0.15$ , corresponding (in principle) to  $T_{\text{eff}} \gtrsim 10^5$  K. To limit our targets to the subset of high-excitation PNe that are closest to Earth, we restricted the Cycle 12 target list to objects with (1) mean distances  $D \leq 1.5$  kpc based on data compiled in Acker et al. (1992), and (2) distances  $D \leq 1.5$  kpc according to either Cahn et al. (1992) or Frew (2008). To this list of objects from Gurzadian (1988), we added the PN Longmore 16 (hereafter Lo 16; Longmore 1977; D. Frew et al. 2012, in preparation).

Based on a search of the *Chandra* archives, we identify 14 additional PNe with  $D \leq 1.5$  kpc previously observed by *Chandra* (13 targeted, and 1 serendipitously observed). In contrast to the volume-limited, excitation-selected Cycle 12 CHANPLANS sample, these 14 PNe were targeted for a variety of (unrelated) reasons—e.g., pre-*Chandra* (e.g., *ROSAT*) X-ray detections (Kastner et al. 2000; Chu et al. 2001; Guerrero et al. 2001), evidence for rapid structural evolution (Kastner et al. 2001), and binary or WR-type central stars (Montez et al. 2005, 2010).

The full (Cycle 12 plus archival data) sample of 35 PNe within  $\sim 1.5$  kpc observed by *Chandra* is listed in Table 1.<sup>23</sup> This table summarizes basic PN and CSPN data for the sample objects; these data are mainly compiled from Frew (2008), with additional morphological classifications following the system described in Sahai et al. (2011, see their Table 2) and results from available molecular ( $\text{H}_2$ ) line observations from Kastner et al. (1996). The last column of Table 1 specifies whether or not the *Chandra* observations (Sections 2.2.1 and 2.3) resulted

<sup>23</sup> Subsets of the sample listed in Table 1 are the subjects of *Herschel Space Observatory* studies of the far-IR emission properties of PNe: “Mass loss of Evolved Stars,” PI: M. Groenewegen, and “The *Herschel* Planetary Nebula Survey,” PI: T. Ueta (early results appear in van Hoof et al. 2011; T. Ueta et al. 2012, in preparation, respectively).

**Table 1**  
Planetary Nebulae Within 1.5 kpc<sup>a</sup> Observed by *Chandra*

Name	PN G	Morph. <sup>b</sup> (F08/SMV11)	<i>D</i> (kpc)	<i>R</i> (pc)	Age (10 <sup>3</sup> yr)	<i>T</i> <sub>*</sub> (kK)	Sp. Type	Comp.	H <sub>2</sub> <sup>c</sup>	X-Rays (Ref.) <sup>d</sup>
PNe observed in Cycle 12										
NGC 650 (M 76)	130.9–10.5	Bas(h)/Bcbpa	1.20	0.40	10	140	PG1159	...	Y	N
NGC 1360	220.3–53.9	Efp/Ecs	0.38	0.31	9	110	O(H)	...	N	P
NGC 1514	165.5–15.2	Ems/Is	0.37	0.12	5	60	O(H):	A0III	...	P
NGC 2346	215.6+03.6	Bs/Bobsp	0.90	0.14	11	≥80	O(H)?	A5V	Y	N
NGC 2371	189.1+19.8	Eps/Bcbspa	1.41	0.13	3	100	[WO1] <sup>e</sup>	...	N	D, P
NGC 2438	231.8+04.1	Emr(h)/Ecs	1.42	0.27	12	124	hgO(H)	...	...	N
NGC 3242	261.0+32.0	Em(h)/Ecspaih	1.00	0.10	4	89	O(H)	...	N	D
NGC 3587 (M 97)	148.4+57.0	Rfm:/Rspi	0.76	0.38	11	105	hgO(H)	...	N	N
NGC 6302	349.5+01.0	Bps/Btp	1.17	0.16	...	220	...	...	Y	N
NGC 6445	008.0+03.9	Bs/Mpi	1.39	0.14	3	170	...	...	Y	P
NGC 6720 (M 57)	063.1+13.9	Ebmr(h)/Ecsh	0.70	0.13	6	148	hgO(H)	...	Y	N
NGC 6772	033.1–06.3	Ep/E	1.20	0.22	19	135	...	...	Y	N
NGC 6781	041.8–02.9	Bam(h:)/Bth	0.95	0.32	26	112	DAO	...	Y	N
NGC 6804	045.7–04.5	Eam/Ms	1.47	0.19	7	85	O(H)	dM? <sup>f</sup>	Y	N
NGC 6853 (M 27)	060.8–03.6	Ebm(h)/Bbpih	0.38	0.37	11	135	DAO	dM?	Y	P
NGC 7008	093.4+05.4	Efp/Bs	0.70	0.15	4	97	O(H)	dG:	...	P
NGC 7009	037.7–34.5	Emps(h)/Lbspa	1.45	0.09	3	87	O(H)	...	N	D, P
NGC 7094	066.7–28.2	Ras/Rs	1.39	0.34	8	110	PG1159	...	N	P
NGC 7662	106.5–17.6	Emp(h)/Esah	1.26	0.09	3	111	O(H) <sup>g</sup>	...	N	D
A 33	238.0+34.8	Ra/R*	1.16	0.78	24	100	O(H)	dK3	...	N
Lo 16	349.3–04.2	Eps/Ispa	0.84	0.17	...	≥82	O(H) <sup>f</sup>	dK?	...	P?
PNe with available archival data										
NGC 40	120.0+09.8	Eas(h)/Bbsh	1.02	0.11	4	48	[WC8]	...	N	D (1)
NGC 246	118.8–74.7	Ea/Es	0.50	0.29	8	140	PG1159	K0V	N	P (2)
NGC 2392	197.8+17.3	Rm/Rsai	1.28	0.14	3	47	Of(H)	dM?	N	D, P (u)
NGC 3132	272.1+12.3	Er/Mtsp	0.81	0.13	6	100	...	A2IV-V	Y	N (u)
NGC 4361	294.1+43.6	Es/Ebs	0.95	0.27	8	126	O(H)	...	N	P (u)
NGC 6543	096.4+29.9	Emps(h)/Mcspa	1.50	0.09	5	48	Of-WR(H) <sup>h</sup>	...	N	D, P (3, 4)
NGC 6826	083.5+12.7	Emp(h)/Ecsah	1.30	0.08	5	50	O3f(H)	...	N	D, P (u)
NGC 7027	084.9–03.4	Bs/Mctspih	0.89	0.03	1.4	175	...	...	Y	D (5)
NGC 7293	036.1–57.1	Bams(h)/Ltspar	0.22	0.46	21	110	DAO	...	Y	P (4)
BD +30°3639	064.7+05.0	Er/Ecsarh	1.30	0.02	1	32	[WC9]	...	Y	D (6)
DS 1	283.9+09.7	Efp/Is	0.73	0.59	19	90	O(H)	M5–6 V <sup>i</sup>	...	P (7)
HFG 1	136.3+05.5	Ea(h:)/R	0.60	0.79	51	100	O(H)	dG:	...	P (7)
IC 418	215.2–24.2	Em(h:)/Ecspih	1.20	0.04	3	38	Of(H)	...	N	D (u)
LoTr 5	339.9+88.4	Eabf/Ss	0.50	0.64	20	100	O(H)	G5III	...	P (7)

**Notes.**

<sup>a</sup> PN and central star data compiled from Frew (2008 and references therein) unless otherwise indicated.

<sup>b</sup> Morphologies as listed in Frew (2008, F08): B: bipolar, E: elliptical, R: round, a: asymmetry present, b: bipolar core present, f: filled (amorphous) center, m: multiple shells present, p: point symmetry present, r: ring structure dominant, s: internal structure noted, (h): distinct outer halo. Morphologies following an abbreviated and very slightly modified version of the classification system described in Sahai et al. (2011) (SMV11; see their Table 2): B: bipolar, M: multipolar, E: elongated, I: irregular, R: round, L: collimated lobe pair, S: spiral arm, c: closed outer lobes, o: open outer lobes; s: CSPN apparent, b: bright (barrel-shaped) central region, t: bright central toroidal structure; p: point symmetry, a: ansae, i: inner bubble, h: halo; r: radial rays.

<sup>c</sup> “Y” = near-IR H<sub>2</sub> detected; “N” = near-IR H<sub>2</sub> not detected; Kastner et al. (1996 and references therein).

<sup>d</sup> X-ray results key: P = point source; D = diffuse source; N = not detected. References to *Chandra* data obtained previous to Cycle 12: (1) Montez et al. 2005; (2) Hoogerwerf et al. 2007; (3) Chu et al. 2001; (4) Guerrero et al. 2001; (5) Kastner et al. 2001; (6) Kastner et al. 2000 (7) Montez et al. 2010; (u) unpublished archival data.

<sup>e</sup> Acker & Neiner (2003).

<sup>f</sup> Frew, unpublished data.

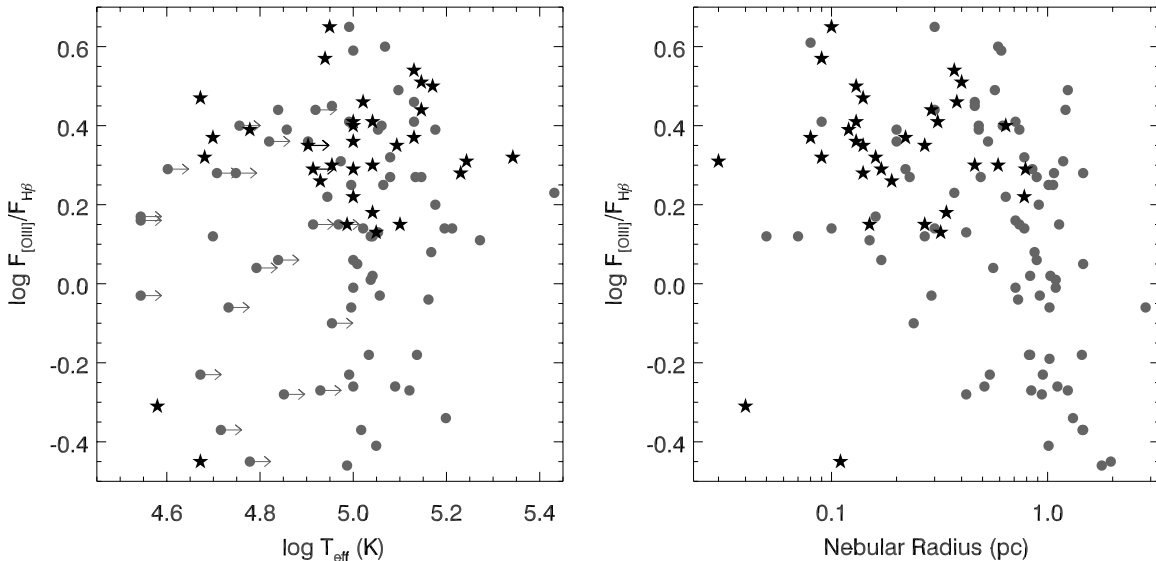
<sup>g</sup> Herald & Bianchi (2011).

<sup>h</sup> Mendez et al. (1990).

<sup>i</sup> Ribeiro & Baptista (2011).

in the detection of point and/or diffuse X-ray sources for a given PN. These results are described in detail in Section 3. Results obtained from archival *Chandra* observations of four PNe (NGC 2392, 3132, 6826, and IC 418) are presented here for the first time (detailed analyses of the observations of NGC 2392, 6826, and IC 418 will appear in N. Ruiz et al. 2012, in preparation; M. Guerrero et al. 2012, in preparation).

Some characteristics of the present sample of PNe, relative to the general population of PNe in the solar neighborhood, are illustrated in Figure 1. The generally high-excitation states of the sample PNe (a result of the selection criteria applied to target objects in Cycle 12) are readily evident in their overall large ratios of [O III] to H $\beta$  flux (with the notable exceptions of the low-excitation PNe IC 418, BD +30°3639, and NGC 40).



**Figure 1.** Characteristics of the sample of 35 PNe within 1.5 kpc observed in X-rays by *Chandra* (black stars) relative to the other  $\sim 85$  known PNe within 1.5 kpc (gray circles; distances from Frew 2008). Left: ratio of [O III] to H $\beta$  fluxes vs. PN central star effective temperature. Right: [O III] to H $\beta$  flux ratio vs. nebular radius. The very low excitation, compact objects BD +303639 (a *Chandra* diffuse X-ray source; Kastner et al. 2000) and M 1–26 (not yet observed in X-rays) lie well below the range of [O III] to H $\beta$  flux ratio represented in these plots.

In addition, as expected, the sample includes many PNe with high  $T_{\text{eff}}$  central stars (Figure 1, left panel), although about a third (including the three aforementioned low-excitation PNe) have estimated  $T_{\text{eff}}$  in the range 30–100 kK (most of the CSPN  $T_{\text{eff}}$  values listed in Table 1 under column heading  $T_{\text{star}}$  were obtained via the Zanstra method; Frew 2008).

## 2.2. Observations

### 2.2.1. *Chandra* X-Ray Observatory

All sample PNe were observed with *Chandra*'s Advanced CCD Imaging Spectrometer (ACIS) using its primary back-illuminated (BI) CCD (S3). With the exception of the PN LoTr 5 (which was observed serendipitously; Montez et al. 2010), each target PN was positioned at the nominal aim point of S3. *Chandra*/ACIS-S3 has energy sensitivity of  $\sim 0.3$ –8 keV, with a field of view of  $\sim 8' \times 8'$  and pixel size  $0''.492$ . In most observations, additional ACIS CCDs were active, extending the effective field of view; these data are not relevant to the analysis described here, however. Use of the BI CCD S3 provides soft X-ray sensitivity that is superior to that of the front-illuminated CCDs on ACIS, effectively extending energy sensitivity to  $\sim 0.2$  keV for the softest PN sources (Section 3.2.1). In addition—due to the large fraction of photon events in which X-ray photon charge is split among adjacent pixels, in BI devices—use of S3 facilitates subpixel event repositioning (SER) in downstream processing (Section 2.3.1), such that post-SER ACIS-S3 images better sample the ( $\sim 0''.5$  FWHM) core of the point-spread function produced by *Chandra*'s High Resolution Mirror Assembly (Li et al. 2004). Observation IDs, dates, and exposure times are listed in Table 2 (all Cycle 12 Large Program observation identifiers begin with “12,” apart from one short, followup exposure targeting NGC 6302).

### 2.2.2. Ground-based Optical Imaging

Images of a subset of Table 1 PNe (included in some panels of Figure 2; see Section 2.3) were obtained with the

Wisconsin–Indiana–Yale–NOAO (WIYN) 0.9 m telescope.<sup>24</sup> The WIYN 0.9 m images were obtained in 2010 November with the S2KB CCD camera ( $0''.6$  pixels;  $20'.48 \times 20'.48$  field of view) and H $\alpha$  filter. Exposure times ranged from 100 s to 500 s, and images were subject to standard processing (dark subtraction, flat-fielding) and astrometric calibration. The H $\alpha$  image of Lo 16 (exposure time 180 s) used in Figure 2 was obtained with GMOS (Hook et al. 2004) on the Gemini 8 m telescope (as part of program GS-2009A-Q-35; see Miszalski et al. 2009b). The image of DS 1 is taken from the SuperCOSMOS H $\alpha$  Survey (Parker et al. 2005).

## 2.3. Data Reduction: The CHANPLANS Pipeline

### 2.3.1. Reprocessing

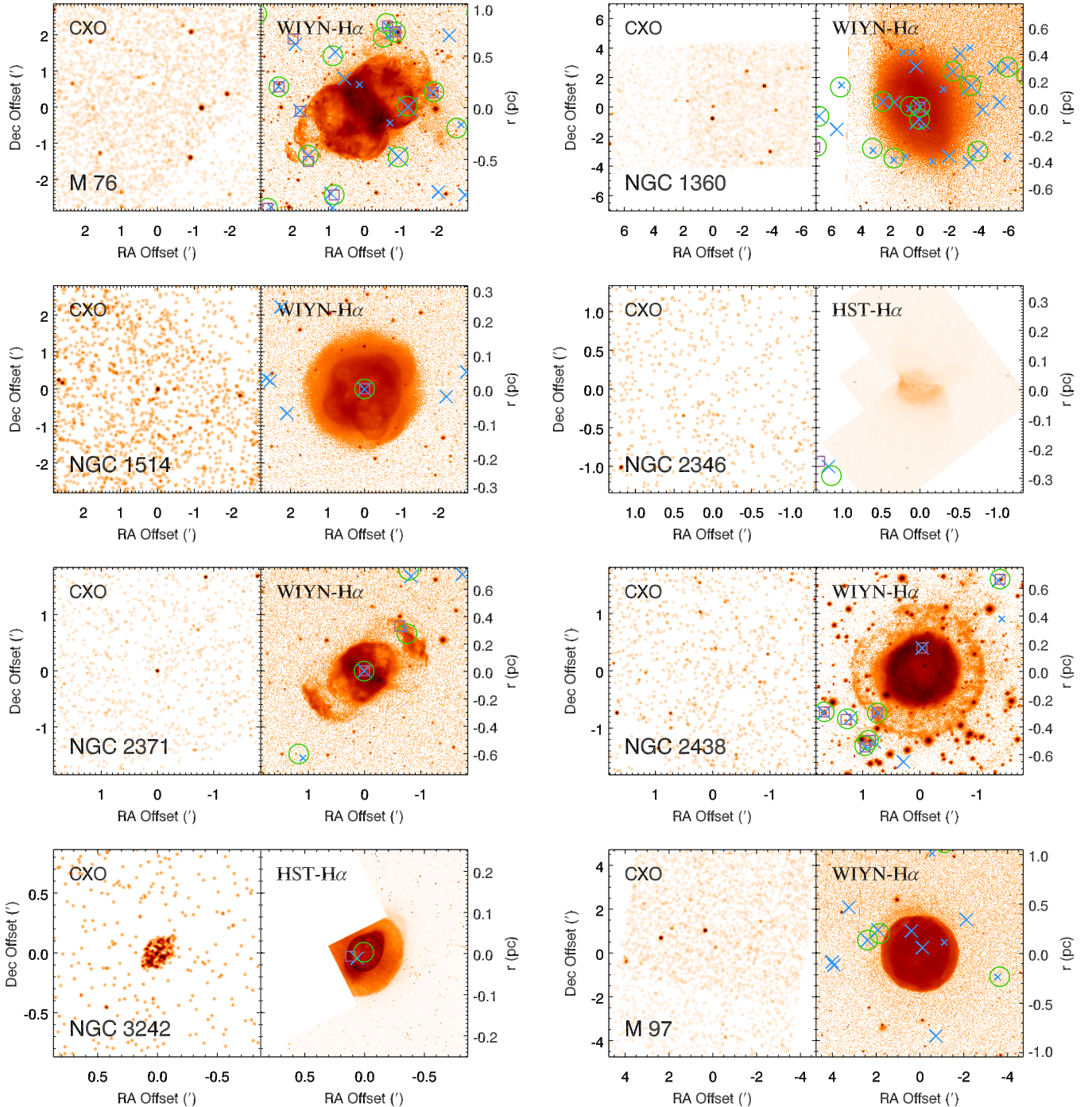
To generate a uniform set of high-level X-ray data products (i.e., images, source lists, spectra, and light curves) from the CHANPLANS observations, we have constructed a processing pipeline consisting of scripts that utilize both CIAO<sup>25</sup> (version 4.3) tools and custom code. The first step is to reprocess the primary and secondary (event and ancillary data) data files provided by the *Chandra* X-ray Center. Reprocessing is performed with the CIAO `chandra_repro` script, which applies the latest calibrations available (CALDB version 4.4.6, in the case of the data presented here) and generates new observation data files (events, bad pixels, aspect solution, etc.). Reprocessing also includes application of SER, so as to optimize the spatial resolution of *Chandra*/ACIS-S3 (Li et al. 2004).

### 2.3.2. Source Detection

For each observation, we search for sources in seven energy filter bands using the CIAO wavelet-based source detection task

<sup>24</sup> The 0.9 m telescope is operated by WIYN Inc. on behalf of a consortium of partner Universities and Organizations that includes RIT. WIYN is a joint partnership of the University of Wisconsin at Madison, Indiana University, Yale University, and the National Optical Astronomical Observatory.

<sup>25</sup> <http://cxc.harvard.edu/ciao/>



**Figure 2.** CHANPLANS pipeline output for the Table 1 PNe. Two panels are presented for each PN. The left panel of each pair shows the *Chandra*/ACIS soft-band (0.3–2.0) keV image, smoothed with a Gaussian function with 3'' FWHM (or 6'' FWHM for images larger than 5' on a side), centered on the SIMBAD coordinates of the PN (which lies on back-illuminated CCD S3), and the right panel shows an optical image (obtained from *HST*, WIYN 0.9 m, or the DSS, as indicated) overlaid with the positions of detected broadband (0.3–8.0 keV) X-ray sources (crosses), USNO catalog stars (circles), and 2MASS Point Source Catalog IR sources (squares). The size of the cross is proportional to the number of X-ray photons detected. (A color version of this figure is available in the online journal.)

wavdetect (Freeman et al. 2002) with wavelength size scales of 1, 2, 4, 8, and 16 pixels. For sources within 4' of the target PN, we use single pixel binning (0.492'' pixels), while for sources at off-axis angles greater than 4', we rebin the images 4 × 4 (~2'' pixels). The wavdetect source detection threshold is set such that the faintest sources identified and recorded are detected at the  $\sim 3\sigma$  significance level. The resulting lists of X-ray sources are then cross-correlated with the USNO-B1 (Monet et al. 2003)

and 2MASS<sup>26</sup> (optical and near-infrared) Point Source Catalogs (PSCs), and the nearest optical and near-infrared sources within 10'' of each detected X-ray source are identified and recorded. Thus far, identification of sources of extended (diffuse) X-ray

<sup>26</sup> This publication makes use of data products from the Two Micron All Sky Survey, which is a joint project of the University of Massachusetts and the Infrared Processing and Analysis Center/California Institute of Technology, funded by the National Aeronautics and Space Administration and the National Science Foundation.

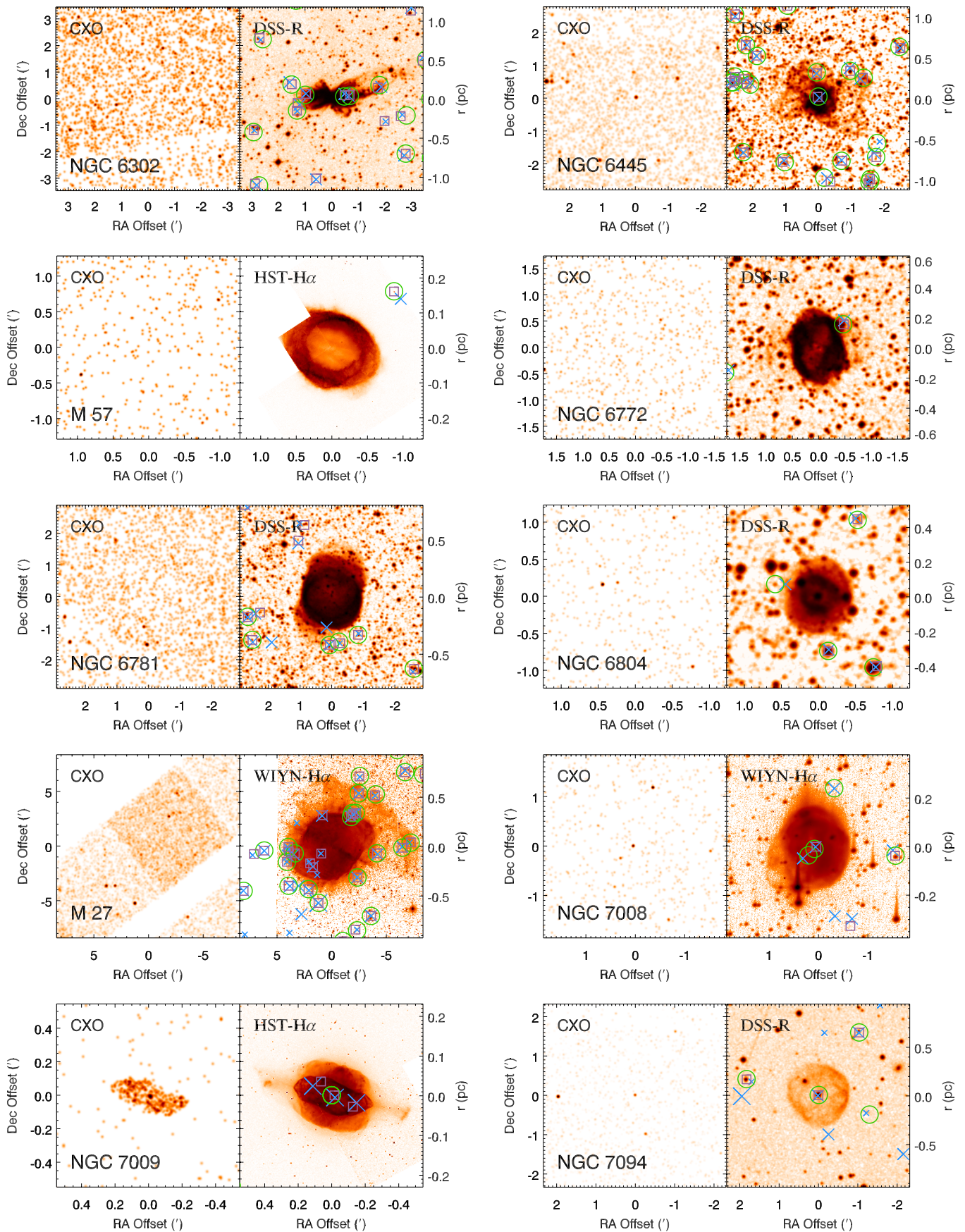


Figure 2. (Continued)

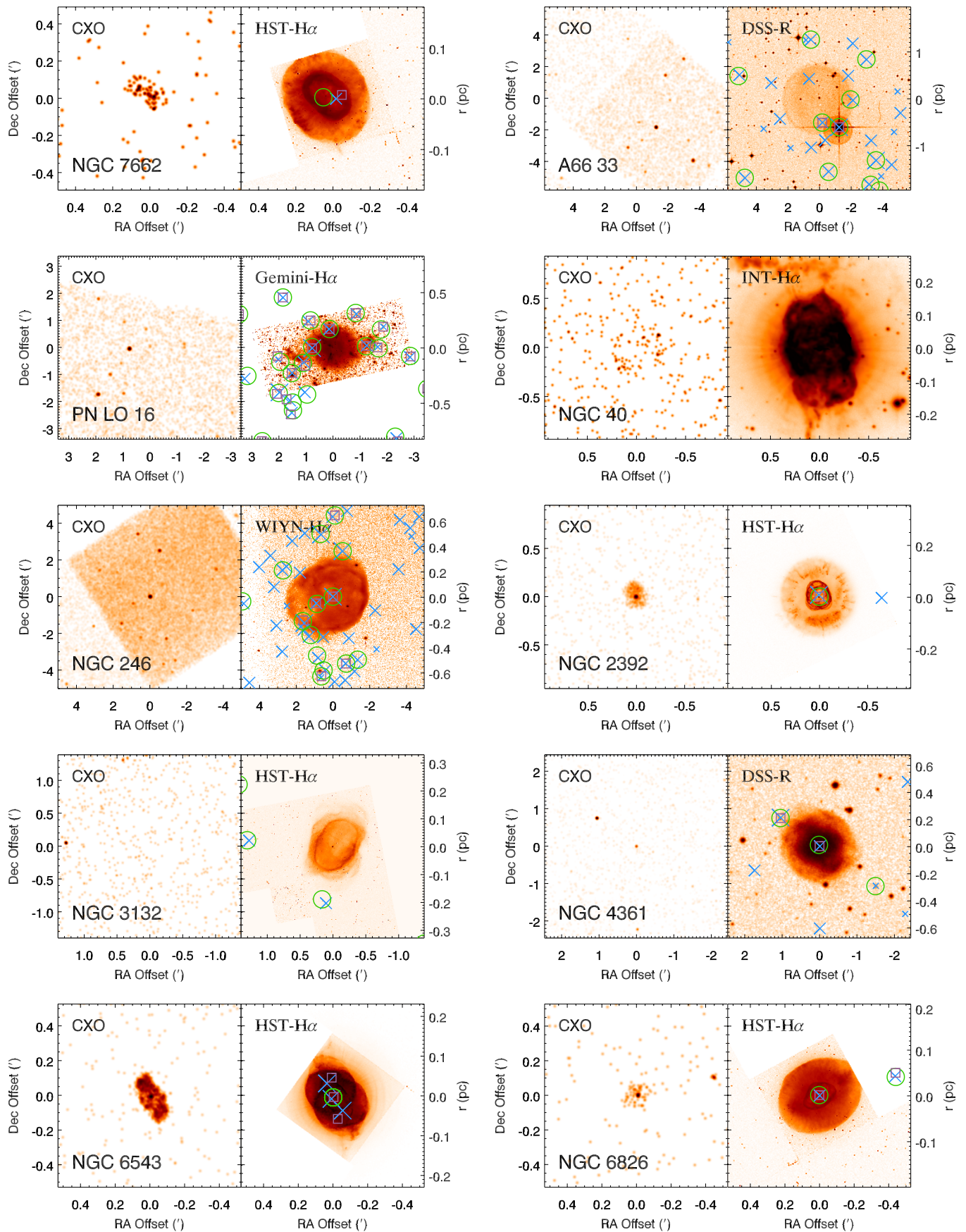


Figure 2. (Continued)



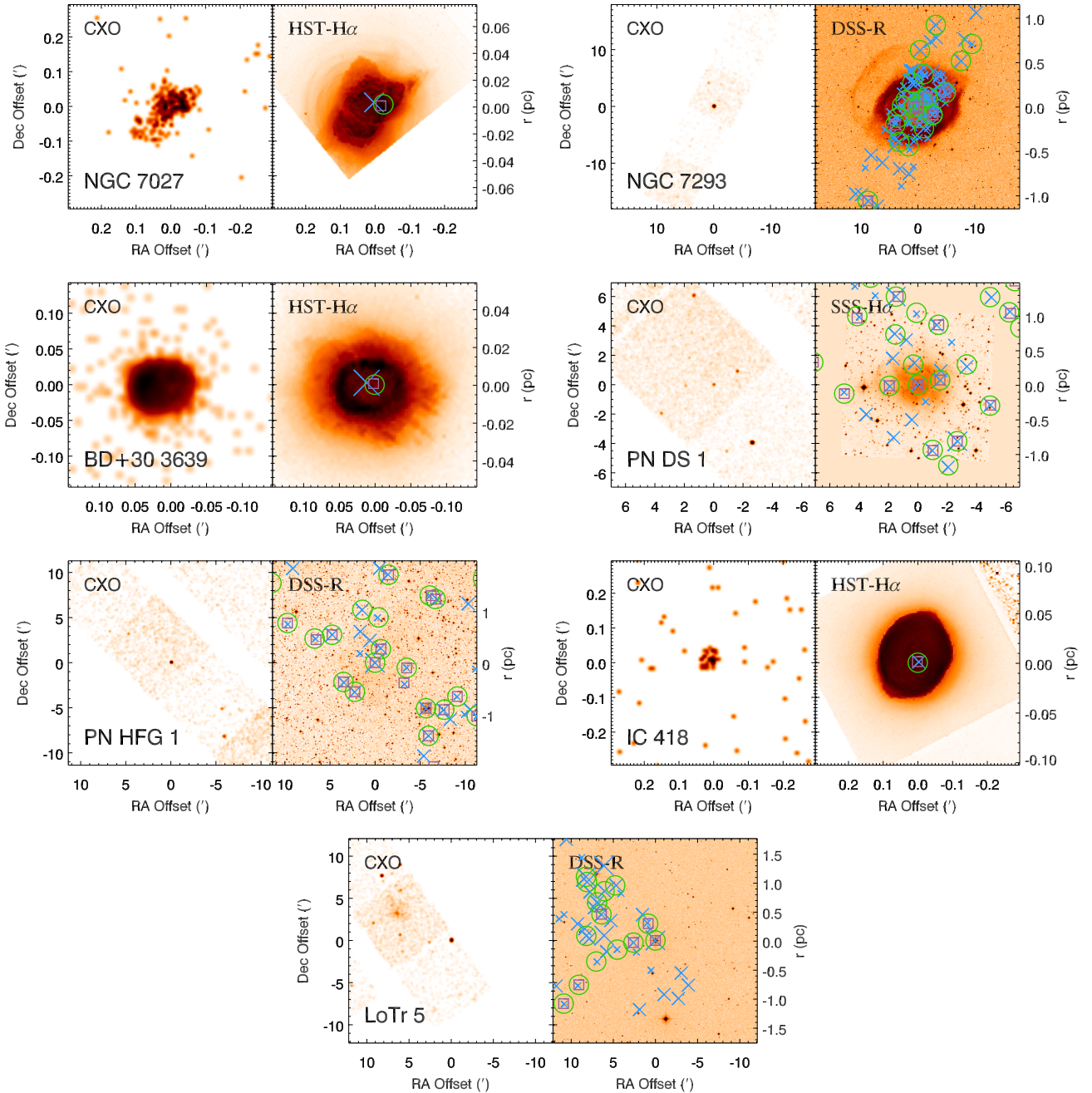


Figure 2. (Continued)

emission (Section 3.2) has been restricted to visual inspection of the soft-band (0.3–2.0 keV) images, assisted in some cases by image rebinning and/or smoothing.

### 2.3.3. Source Event Statistics and Spectral Extraction

For those PNe whose central stars are detected as X-ray sources (Sections 3.1 and 3.2), we calculate statistics for the events in a  $3''.5$  radius region centered on the CSPN, so as to determine the total number of source photons and the mean, median, and first and second quartile photon energies. The median energy is an observed quantity that is dependent on instrumental energy response; for an ensemble of sources observed with a particular instrument, however, median photon energy is indicative of the source plasma temperature and intervening absorbing column (Getman et al. 2010). We also perform spectral extractions within regions of interest encompassing the

CSPN, nebula (both including the central star and excluding the central star), and source-free background regions. The sizes and morphologies of the nebular extraction regions are determined from the optical morphologies of the nebula. This extraction—which results in generation of source and background region X-ray spectra and all associated *Chandra*/ACIS response (source-specific calibration) files necessary for spectral model fitting—is performed for all objects, whether detected or not detected. Analysis of the resulting X-ray spectra will be presented in forthcoming CHANPLANS papers (e.g., R. Montez et al. 2012, in preparation).

### 2.3.4. Pipeline Output: Annotated Images

Figure 2 illustrates the results of the processing pipeline just described. The two panels included for each PN—presented in the order listed in Table 1—display the *Chandra* soft-band

**Table 2**  
Log of *Chandra* Observations

Name	OBSID	Date	Exposure (ks)
A 33	12369	2011 Jan 21	29.67
BD +30	587	2000 Mar 21	19.22
"	10821	2009 Jan 22	38.63
"	9932	2009 Jan 27	38.03
DS 1	9953	2009 Jul 19	24.15
HFG 1	9954	2008 Dec 11	11.45
IC 418	7440	2006 Dec 12	49.38
Lo 16	12367	2012 Jan 30	30.00
LoTr 5	3212	2002 Dec 4	27.74
NGC 40	4481	2004 Jun 13	19.89
NGC 246	2565	2002 Oct 22	40.95
NGC 650-51	12371	2010 Nov 15	29.90
NGC 1360	12362	2010 Nov 19	19.65
NGC 1514	12361	2010 Nov 4	20.78
NGC 2346	12379	2010 Dec 19	29.89
NGC 2371-72	12376	2012 Feb 18	29.67
NGC 2392	7421	2007 Sep 13	57.41
NGC 2438	3765	2003 Apr 21	49.75
"	12377	2011 Feb 6	29.67
NGC 3132	4514	2004 Aug 8	23.97
NGC 3242	12380	2011 Feb 28	29.26
NGC 3587	12366	2011 Jul 5	19.24
NGC 4361	3760	2003 Feb 17	29.38
NGC 6302	14364	2012 Apr 25	7.47
"	12370	2011 Oct 25	22.54
NGC 6445	12375	2011 Feb 19	29.68
NGC 6543	630	2000 May 10	46.06
NGC 6543	11999	2009 Sep 26	25.03
NGC 6543	10443	2009 Sep 21	22.03
NGC 6720	12364	2011 Jan 23	19.80
NGC 6772	12372	2011 Jun 30	29.37
NGC 6781	12368	2011 Nov 22	28.25
NGC 6804	12378	2011 Nov 30	29.58
NGC 6826	8559	2007 Jul 24	15.04
"	7439	2007 Jun 11	34.08
NGC 6853	12363	2010 Dec 17	19.80
NGC 7008	12365	2011 Nov 10	18.22
NGC 7009	12381	2011 Jun 25	29.66
NGC 7027	588	2000 Jun 1	18.22
NGC 7094	12374	2011 Apr 21	29.67
NGC 7293	631	1999 Nov 17	37.13
"	1480	1999 Nov 18	11.02
NGC 7662	12373	2012 May 15	27.61

**Table 3**  
Planetary Nebula X-Ray Point Source Characteristics

Name	$N^a$ (photons)	$C^b$ (ks $^{-1}$ )	Median $E^c$ (keV)	$E$ Range $^d$ (keV)
NGC 40 $^e$	...	...	...	...
NGC 246	749	18.5 $\pm$ 0.3	0.33	0.26–0.39
NGC 650	...	<0.16	...	...
NGC 1360	25	1.26 $\pm$ 0.25	0.27	0.17–0.56
NGC 1514	13	0.62 $\pm$ 0.17	0.72	0.59–0.98
NGC 2346	...	<0.13	...	...
NGC 2371-72	29	0.96 $\pm$ 0.18	0.36	0.31–0.47
NGC 2392 $^e$	241:	4.2:	0.95:	0.61–1.46
NGC 2438	...	<0.13	...	...
NGC 3132	...	<0.16	...	...
NGC 3242 $^e$	...	<0.5:	...	...
NGC 3587	...	<0.25	...	...
NGC 4361	43	1.48 $\pm$ 0.22	0.26	0.19–0.40
NGC 6302	...	<0.17	...	...
NGC 6445	33	1.10 $\pm$ 0.19	1.04	0.90–1.20
NGC 6543 $^e$	165:	3.6:	0.55:	0.43–0.73
NGC 6720	...	<0.20	...	...
NGC 6772	...	<0.13	...	...
NGC 6781	...	<0.13	...	...
NGC 6804	...	<0.13	...	...
NGC 6826 $^e$	27:	0.8:	0.71:	0.55–0.90
NGC 6853	173	8.74 $\pm$ 0.66	0.18	0.17–0.20
NGC 7008	23	1.26 $\pm$ 0.26	0.85	0.67–1.29
NGC 7009 $^e$	31:	1.0:	0.74:	0.57–1.01
NGC 7027 $^e$	...	<0.5:	...	...
NGC 7094	28	0.93 $\pm$ 0.18	0.55	0.41–0.95
NGC 7293	396	35.9 $\pm$ 1.8	0.89	0.73–1.06
NGC 7662 $^e$	...	<0.5:	...	...
A 33	...	<0.13	...	...
BD +30 $^{\circ}$ 3639 $^e$	...	...	...	...
DS 1	55	2.29 $\pm$ 0.31	1.01	0.79–1.28
HFG 1	143	12.6 $\pm$ 1.0	1.05	0.83–1.38
IC 418 $^e$	...	...	...	...
Lo 16	8	0.26 $\pm$ 0.09	1.09	0.92–1.84
LoTr 5	285	10.27 $\pm$ 0.61	1.13	0.88–1.55

**Notes.**<sup>a</sup> Number of source photons, after background subtraction.<sup>b</sup> Source photon count rate.<sup>c</sup> Median source photon energy.<sup>d</sup> Source photon energy range (25th through 75th percentiles).<sup>e</sup> Point source counts, count rate (or upper limit), median energy, and energy ranges are uncertain due to presence of diffuse emission component.

(0.3–2.0 keV) X-ray image centered on the position of the PN (left panel) and the positions of *Chandra*-detected broadband (0.3–8.0 keV) X-ray sources, USNO-B1.0 catalog stars, and 2MASS PSC IR sources overlaid on an optical ( $H\alpha$  or  $R$  band) image (right panel). The *Hubble Space Telescope* (*HST*) optical ( $H\alpha$ ) images in the right-hand panels were obtained from the *Hubble* Legacy Archive<sup>27</sup> where available; otherwise, we display the ground-based images described in Section 2.2.2 or  $R$ -band images from the Digital Sky Survey<sup>28</sup> (DSS).

<sup>27</sup> Based on observations made with the NASA/ESA Hubble Space Telescope, and obtained from the *Hubble* Legacy Archive, which is a collaboration between the Space Telescope Science Institute (STScI/NASA), the Space Telescope European Coordinating Facility (ST-ECF/ESA) and the Canadian Astronomy Data Centre (CADM/NRC/CSA) where available; otherwise, <http://archive.stsci.edu/hst>

<sup>28</sup> <http://archive.stsci.edu/dss>

## 3. RESULTS

Results from the *Chandra* observations listed in Table 2 and illustrated in Figure 2 are summarized in Tables 1 and 3. The rightmost column of Table 1 states whether or not the PN was detected and, in the case of a detection, whether the PN displays a point-like X-ray source at the CSPN, diffuse X-ray emission, or a combination of the two. Based on preliminary model fits to the spectra extracted for the X-ray-faintest objects in Table 2—i.e., the diffuse source and soft X-ray point sources within NGC 2371 (each of which displays a count rate  $\sim 1$  ks $^{-1}$ ) and the “hard X-ray” point source within Lo 16 (count rate  $\sim 0.3$  ks $^{-1}$ )—we conservatively estimate that our sensitivity limits for diffuse and hard (soft) point-like X-ray sources are  $\sim 10^{30}$  and  $\sim 10^{29}$  ( $\sim 10^{31}$ ) erg s $^{-1}$ , respectively, at the limiting (1.5 kpc) distance of the survey (where the soft source luminosity limit is strongly dependent on the intervening absorbing column along the line of sight to the CSPN).

**Table 4**  
Planetary Nebulae: *Chandra* X-Ray Detection Statistics

Category <sup>a</sup>	$N^b$	$N_{PX}^c$	$N_{DX}^d$
Entire sample	35	18 <sup>e</sup> (53%)	11 (31%)
Round/elliptical, F08	28	16 <sup>e</sup> (57%)	10 (36%)
Bipolar, F08	7	2 (28%)	1 (14%)
Round/elliptical/irregular, SMV11	20	11 <sup>e</sup> (55%)	6 (30%)
Bipolar/multipolar, SMV11	15	7 (47%)	5 (33%)
Near-IR H <sub>2</sub> not detected	15	9 (60%)	9 (60%)
Near-IR H <sub>2</sub> detected	13	3 (24%)	2 (15%)
Known binary CSPN	13	9 <sup>e</sup> (69%)	1 (8%)

#### Notes.

<sup>a</sup> Morphologies as listed and defined in Column 3 of Table 1 and associated footnotes; CSPN binary detections and H<sub>2</sub> detections as listed, respectively, in Columns 9 and 10 of Table 1.

<sup>b</sup> Total number of sample PNe in each category.

<sup>c</sup> Number of PNe in each category displaying point-like X-ray emission in *Chandra* imaging.

<sup>d</sup> Number of PNe in each category displaying diffuse X-ray emission in *Chandra* imaging.

<sup>e</sup> Includes tentative detection of Lo 16.

Table 3 lists basic characteristics of the point-like X-ray source (net background-subtracted photon counts, count rates, median energy, and energy ranges)—or point source nondetection (upper limit on count rate), as the case may be—that is associated with each PN. Comparisons of X-ray and optical emission morphologies for the Table 1 PNe detected as diffuse X-ray emission sources are presented in Figures 3 and 4. A summary of the PNe observed and detected, broken down into various object categories (i.e., primary PN morphology descriptor, as listed in Column 3 of Table 1 and described in associated footnotes, detection and nondetection of near-IR H<sub>2</sub> emission, and known binary CSPNe), is presented in Table 4 and Figure 5. Discussions of individual diffuse and point-like PN X-ray sources (including the implications of these sources for the origin and evolution of their “host” PNe), as well as presentations of the results of detailed modeling of individual sources, are deferred to subsequent papers (e.g., R. Montez et al. 2012, in preparation). Here, we present summaries of the main results.

### 3.1. PNe Displaying Diffuse X-Ray Emission

*BD+30°3639*, *IC 418*, *NGC 40*, *2371*, *2392*, *3242*, *6543*, *6826*, *7009*, *7027*, and *7662*: The *Chandra* observations reported here establish that these 11 PNe are diffuse X-ray sources. The CHANPLANS and archival images of these PNe (as well as contours of these X-ray images overlaid on optical images) are displayed in Figure 3 and are grouped according to whether or not the PN also displays a point-like emission component at its central star.

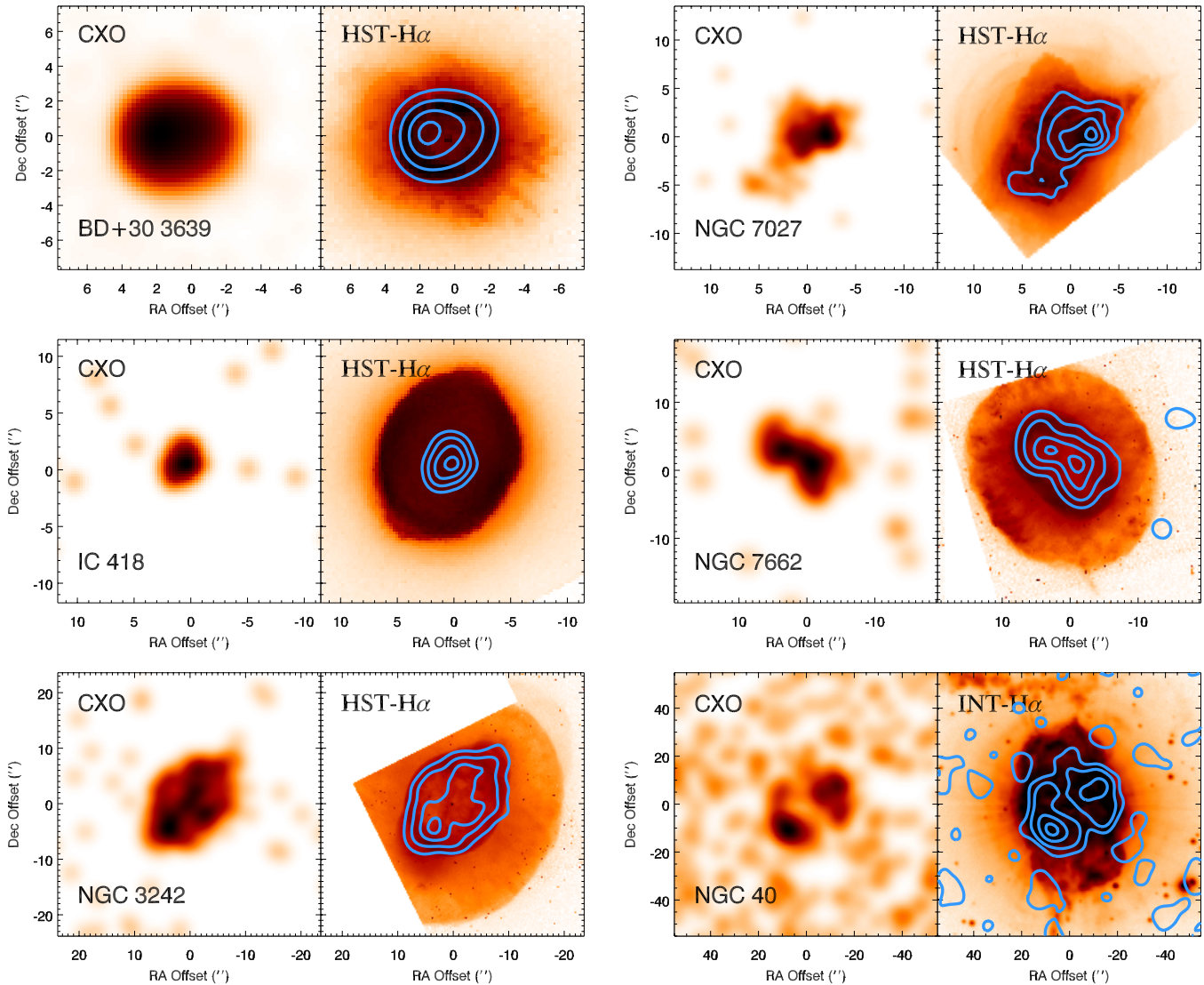
*BD+30°3639* (Campbell’s Star; Kastner et al. 2000; Yu et al. 2009) and *NGC 6543* (Cat’s Eye; Chu et al. 2001; Kastner et al. 2002) were the earliest established—and remain the best documented—examples of PNe displaying emission from wind-shock-generated “hot bubbles” (Figure 3). No X-rays appear to be specifically associated with the CSPN of *BD+30°3639*. However, *NGC 6543* represents a case study of a PN harboring both diffuse (hot bubble) and harder, point-like (CSPN) X-ray emission components (Chu et al. 2001; Guerrero et al. 2001). Similarly, the archival CHANPLANS observations of *NGC 2392* (Eskimo) and *NGC 6826* and the Cycle 12 CHANPLANS observation of *NGC 7009* (Saturn) reveal that

both soft diffuse and harder point-like (CSPN) emission components are also clearly present in these PNe (Figure 3; the diffuse and point-source X-ray components, respectively, of *NGC 2392* and *6826* are the subject of forthcoming papers by N. Ruiz et al. 2012, in preparation; M. Guerrero et al. 2012, in preparation). The *Chandra* detection of relatively hard point source emission in *NGC 2392* likely explains the energy dependence of the X-ray morphology apparent in the earlier *XMM* data; see Guerrero et al. (2005). Likewise, in the case of *NGC 7009*—which was also previously detected by *XMM* (Figure 4, lower left; Guerrero et al. 2002)—*Chandra*’s spatial resolution establishes the presence of an X-ray point source embedded within diffuse X-ray nebulosity (Figure 4, lower right). The remaining example of a PN harboring both diffuse and point-like X-ray emission, *NGC 2371*, displays a softer CSPN X-ray source and very faint diffuse emission that partially fills its central regions (Figure 3).

In an *XMM* observation (Figure 4, upper left), *NGC 3242* (the Ghost of Jupiter nebula) is well detected and appears as a marginally extended, asymmetric X-ray source (Ruiz et al. 2011). However, as in the cases of *NGC 2392* (Guerrero et al. 2005) and *7009* (Guerrero et al. 2002), the diameter of the inner nebula of *NGC 3242* is similar to the width of the *XMM*/EPIC (pn and MOS) point spread functions, rendering its *XMM* X-ray morphology difficult to interpret and (in particular) the potential contribution from an X-ray-luminous CSPN impossible to ascertain. In CHANPLANS imaging (Figure 4, upper right), the X-ray emission from *NGC 3242* is clearly established as diffuse, with a smooth surface brightness distribution that traces the inner shell, including the protrusions along the shell’s major axis. Notably, no X-ray point source is evident at the CSPN of *NGC 3242* in the CHANPLANS image. The morphologically similar *NGC 7662* (the Blue Snowball nebula) is a new (Cycle 12 *Chandra*) X-ray detection. As in the case of *NGC 3242*, we detect only diffuse emission within *NGC 7662*.

All three low-excitation nebulae (as measured in terms of [O III] to H $\beta$  line ratio) observed thus far by *Chandra*—*BD+30°3639*, *IC 418*, and *NGC 40* (all of which were observed prior to Cycle 12)—display diffuse X-rays from hot bubbles, and all three lack point source (CSPN) emission (Figure 3). Two of the three, *BD+30°3639* and *NGC 40*, harbor late [WC]-type CSPNe with dense, fast winds, leading to the suggestion that WR-type central stars are efficient at generating X-ray-luminous wind-blown bubbles within PNe (Montez et al. 2005; Kastner et al. 2008). The CHANPLANS detection of diffuse X-rays within *NGC 2371*, which harbors an early [WO]-type CSPN, further reinforces this notion. Indeed, the three objects with WR-type CSPNe in Table 1—*BD+30°3639*, *NGC 40*, and *NGC 2371*—appear to represent a sequence in which both CSPN effective temperature and PN radius increase as diffuse X-ray luminosity decreases; *NGC 2371* is the only one of the three to display X-rays from its CSPN. While the CSPN of *IC 418* has not been classified as a [WC] type, this CSPN is as cool as the late-type [WC] stars within *BD+30°3639* and *NGC 40* and, like such stars, it drives a relatively strong, fast wind (Cerruti-Sola & Perinotto 1989).

In contrast to the PNe just discussed, all of which appear to display diffuse X-ray emission that is confined to “hot bubbles,” *NGC 7027* (which was observed early in the *Chandra* mission; Kastner et al. 2001, 2002) provides a rare, clear example of X-ray emission from collimated flows within a nearby PN (Figure 3; Kastner 2009). Specifically, the X-ray emission closely traces



**Figure 3.** Images of Table 1 PNe in which diffuse X-ray emission has been detected by *Chandra*. The left and right panel pairs for each PN display, respectively, *Chandra* 0.3–2.0 keV images and *Chandra* contours overlaid on optical images. The *Chandra* images of all but two PNe (NGC 40 and NGC 2371) have been smoothed with a 3'' FWHM Gaussian (the *Chandra* images of NGC 40 and NGC 2371 have been smoothed with a FWHM of 8''); contour levels are 10%, 30%, 60%, and 90% of the maximum X-ray surface brightness. The PNe in the first six panels display only diffuse emission; the PNe in the subsequent five panels display both diffuse and point-source X-ray emission components.

(A color version of this figure is available in the online journal.)

portions of the central, elliptical shell that have evidently been punctured by high-velocity bullets or jets (Cox et al. 2002).

### 3.2. PNe Displaying Only Point-like X-Ray Emission at Central Stars

In addition to the five PNe that display both diffuse and point-like X-ray emission components, 13 objects display only point-like X-ray emission from their CSPNe. These X-ray sources appear to belong to two general classes: very soft sources with median photon energies  $< 0.4$  keV, most of which have hard X-ray “tails” (Section 3.2.1); and harder sources, with median photon energies ranging from  $\sim 0.5$  keV to  $\sim 1.0$  keV (Section 3.2.2).

#### 3.2.1. Objects with Strong or Dominant “Hot CSPN Photosphere” X-Ray Spectral Components

NGC 246, 1360, 4361, and 6853. These nebulae appear to represent a distinct group of CSPNe that display a

combination of relatively high  $T_{\text{eff}}$  and low median X-ray photon energy (Figure 6). However, only NGC 6853 (the Dumbbell nebula) displays an X-ray (ACIS-S3) SED that is consistent with “pure” photospheric emission from a hot ( $\gtrsim 100$  kK) CSPN (this CHANPLANS result supports the previous analysis of *ROSAT* X-ray data by Chu et al. 1993). The detection of such emission from the CSPN of NGC 6583 by *Chandra* is facilitated by its proximity (see Table 1 and Section 3.3). As is evident in their detected photon energy ranges (Figure 6), each of the other CSPNe X-ray sources in this group—although significantly softer than the sources discussed in Section 3.2.2—displays a  $\sim 0.4$ – $0.6$  keV “tail” in its X-ray SED, indicative of excess emission above that expected from a hot CSPN photosphere (R. Montez et al. 2012, in preparation). The CSPN X-ray source within NGC 2371 (which shows very faint diffuse emission; Section 3.1.1) is also in this category, i.e., a “soft” source with a “hard tail” (Figure 6). The (archival) ACIS-S3 data for NGC 246 were previously published by Hoogerwerf et al. (2007), who found that the soft portion of the CSPN X-ray spectrum was

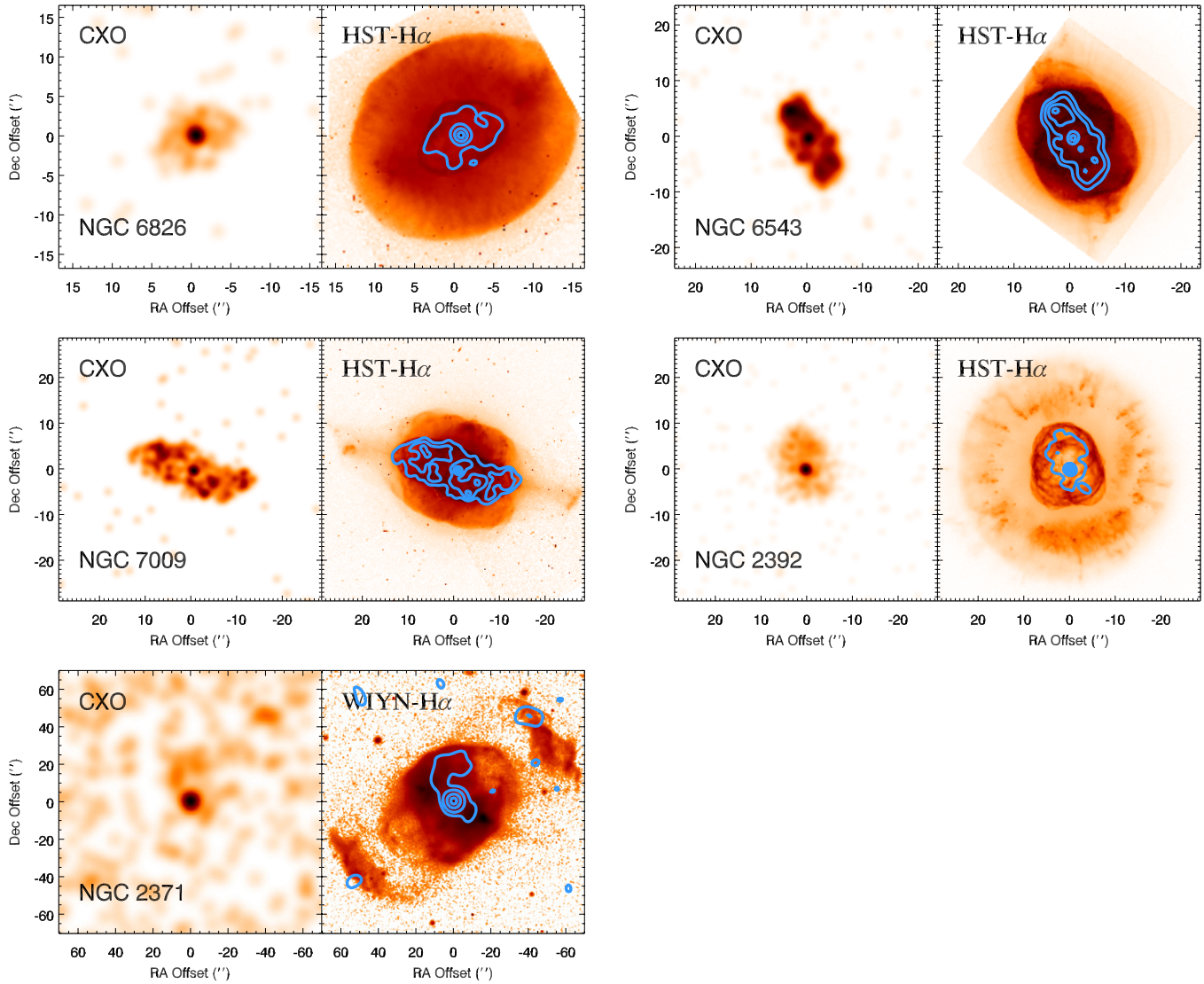


Figure 3. (Continued)

consistent with non-LTE models describing PG1159 star atmospheres, but that an additional component (consisting of emission lines of highly ionized C) was necessary to account for excess flux in the 0.3–0.4 keV energy range.

### 3.2.2. Objects with Strong or Dominant “Hard” X-Ray Spectral Components

*NGC 1514, 6445, 7008, 7094, 7293, DS 1, HFG 1, Lo 16, LoTr 5.* NGC 7293 (Helix) is the prototype of an X-ray-luminous CSPN whose X-ray SED extends to energies far too high to be explained as due to a hot pre-WD photosphere (Guerrero et al. 2000, 2001). Cycle 12 CHANPLANS and archival observations have yielded a dozen more examples of such hard X-ray excesses at CSPNe (Figure 6), including point sources clearly associated with five PNe that display “hot bubble” X-ray emission (Section 3.1). All of these CSPN X-ray sources have median X-ray energies in the range  $\sim 0.5$ – $1.0$  keV, i.e., a factor  $\sim 2$ – $5$  larger than the CSPN X-ray sources described in Section 3.2.1. Remarks on most of these “hard X-ray CSPNe” follow.

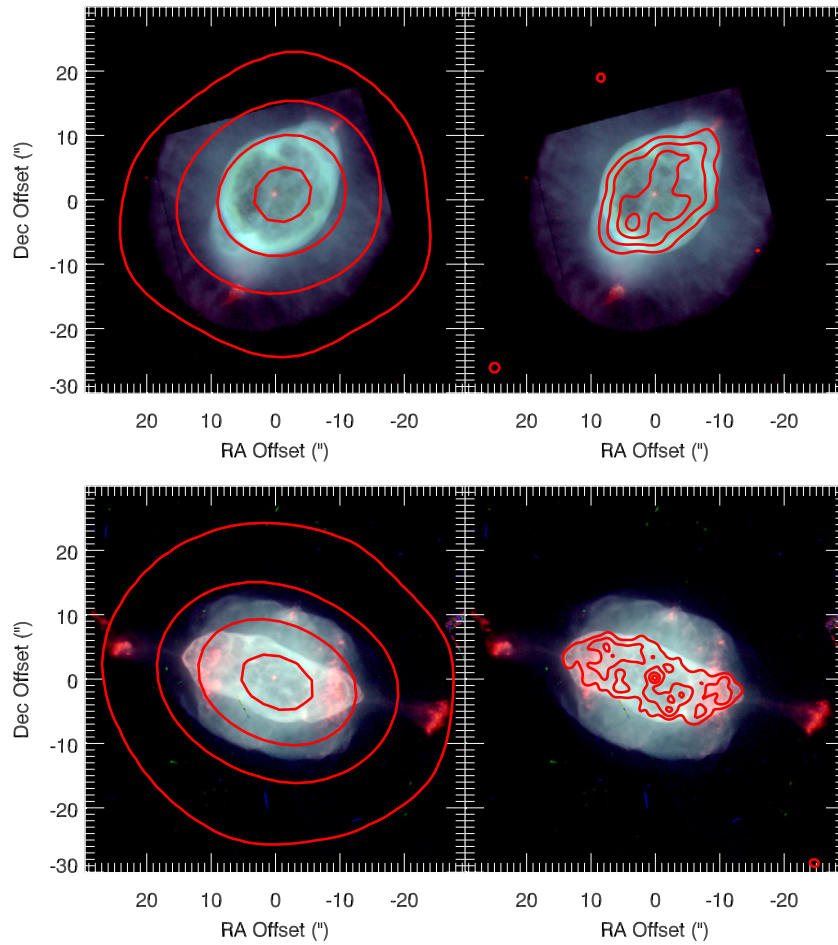
**NGC 1514.** This PN has an unremarkable, amorphous optical morphology, but mid- to far-infrared imaging has revealed

striking bipolar, double-ring dust structures exterior to the ionized nebula (Ressler et al. 2010; Aryal et al. 2010). The central star has a companion of type A0 (Ciardullo et al. 1999)—the earliest spectral type among known binary companions of “hard X-ray CSPNe” (see Section 4.2).

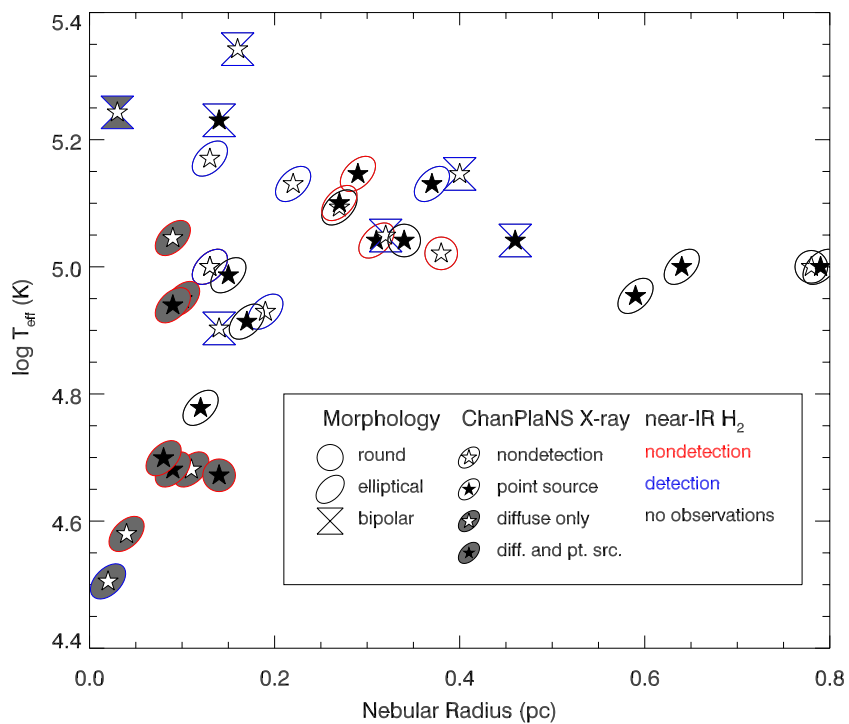
**NGC 6445.** This object is a bipolar PN with faint lobes and a bright central ring or torus. The central, point-like X-ray source detected in CHANPLANS imaging, which is offset by  $\sim 3''$  from the SIMBAD coordinates of the PN, may be the first secure detection of the central star (or central binary, as the case may be) of this PN. This X-ray source is slightly off-center within the central ring.

**NGC 7008.** This PN displays an X-ray point source coincident with the position of its central star as listed in the *HST* Guide Star Catalog and 2MASS PSC. The central star has a probable late-type (G) companion at separation  $\sim 160$  AU (Ciardullo et al. 1999).

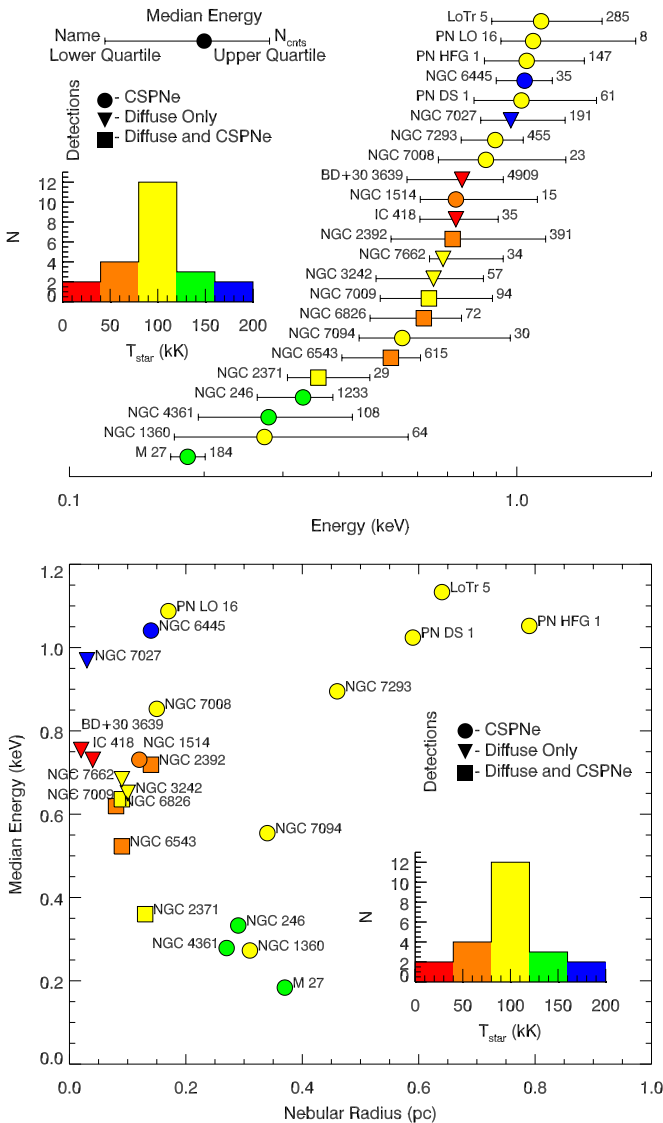
**NGC 7293.** This PN displays a composite X-ray SED consisting of a soft “hot blackbody” component and harder, higher-temperature component (Guerrero et al. 2001). As in the cases of known binary CSPNe (see next), the harder X-ray component may arise in the corona of a late-type



**Figure 4.** Color montages of *HST* archival images of NGC 3242 (top panels) and NGC 7009 (bottom panels) overlaid with contours of X-ray surface brightness as imaged by *XMM* (left panels) and *Chandra* (right panels). Contour levels are 10%, 30%, 60%, and 90% of the peak.



**Figure 5.** Plot of CSPN  $T_{\text{eff}}$  vs. PN radius for the Table 1 objects, with symbols indicating presence or absence of diffuse or point-like X-ray emission, as well as PN morphology and presence or absence of H<sub>2</sub> emission (see Table 1 and associated footnotes, and Table 4).



**Figure 6.** Top: photon energy statistics (net source counts, median energy, and first and third quartile energies) for PN X-ray sources, ordered from lowest to highest median energy (bottom to top). Symbols indicate nature of emission (point-like at CSPN, diffuse, or both CSPN and diffuse) and are color-coded according to CSPN  $T_{\text{eff}}$  (see inset histogram, which also displays the distribution of  $T_{\text{eff}}$  for CSPNe detected as X-ray sources). Bottom: PN X-ray source median energy vs. PN radius, with symbols as in the top panel.

(dM) companion—a possibility bolstered by detection of variable  $H\alpha$  emission from the CSPN (Gruendl et al. 2001). Additional supporting evidence for the presence of such a companion remains elusive, however (e.g., O’Dwyer et al. 2003).

**DS 1, HFG 1, and LoTr 5.** All three PNe feature late-type companions to their CSPNe. The PNe DS 1 and HFG 1 were targeted by *Chandra* because their CSPNe are known close binaries and, hence, candidate post-common envelope objects, while the companion to the CSPN of LoTr 5, which was observed serendipitously, is a Ba-rich giant (Montez et al. 2010 and references therein). In each case, the X-ray characteristics of the central, point-like source are consistent with coronal emission—as expected if the source is a late-type binary companion that has been spun up by accretion of material lost by the PN progenitor (see Section 4.2 and Montez et al. 2010).

**Lo 16.** This chaotic nebula that also harbors a (close, 0.49 day period) binary CSPN (D. Frew et al. 2012, in preparation) and is very tentatively detected in Cycle 12 imaging as a rather hard CSPN X-ray source (median energy  $\sim 1.1$  keV).

### 3.3. X-Ray Nondetections

NGC 650–1, 2346, 2438, 3132, 3587, 6720, 6772, 6781, 6804, A33. The majority of these X-ray-nondetected objects are molecule-rich PNe (Kastner et al. 1996) with morphologies that are either sharply bipolar (NGC 650–1, 2346) or Ring-like (NGC 3132, 6720, 6772, 6781). The Ring-like PNe likely have intrinsically axisymmetric density structure—fundamentally similar to the structures of clearly morphologically bipolar (pinched-waist) PNe such as NGC 650–1 and NGC 2346 (Kastner et al. 1994). NGC 2346 also was undetected in *XMM* imaging (Gruendl et al. 2006), but the Cycle 12 CHANPLANS nondetection places more severe constraints on the X-ray luminosity of its CSPN (R. Montez et al. 2012, in preparation). Tarafdar & Apparao (1988) reported a  $\sim 3\sigma$  *Einstein* X-ray Observatory detection of A33, but CHANPLANS imaging demonstrates this association is spurious, and the *Einstein* source can most likely be attributed to an X-ray-luminous (and optically bright) field star near the southwest edge of the nebula, possibly combined with 1–2 weaker field X-ray sources within the boundaries of the PN. Although the CSPN of NGC 3587 was detected as a very soft X-ray source by *ROSAT* (Chu et al. 1998), its nondetection here is not surprising, given its lower luminosity, lower photospheric temperature, and larger distance relative to the CSPN of NGC 6583 (Section 3.2.1).

## 4. DISCUSSION

The sample of 35 PNe within  $\sim 1.5$  kpc observed to date by *Chandra* affords the first opportunity for relatively unbiased statistical investigations of the spatial and spectral characteristics of PN X-ray emission, so as to inform studies of PN formation and evolution. The initial sample (Table 1) is still rather heterogeneous and prone to selection effects, as it is composed of a mixture of high-excitation PNe and small subsets of objects specifically targeted for various reasons (Section 2.1); furthermore, data analysis is still in its early stages. Hence, it is premature to draw firm conclusions concerning the nature(s) of X-ray sources within PNe, much less the implications of PN X-ray emission (or lack thereof) for the origin and shaping of PNe. Nevertheless, a few preliminary trends are apparent in the initial CHANPLANS results described in Section 3. We highlight and comment on these trends in the following subsections.

### 4.1. Diffuse X-Ray Emission from PNe

Ten of the eleven sample PNe displaying diffuse emission (Section 3.1; Figure 3) are classified by Frew (2008) as elliptical or round nebulae (the lone exception being NGC 7027, which is classified as bipolar by Frew 2008). These diffuse X-ray PNe are also generally molecule-poor (e.g., they lack detections of near-IR  $H_2$  emission; see Table 4 and Kastner et al. 1996), with the notable exceptions of BD +30°3639 and NGC 7027. In optical imaging, the diffuse X-ray PNe display multiple, nested shells with well-defined innermost bubbles (Frew morphology subclass of “m” and/or Sahai et al. secondary morphology characteristic of “i” in Table 1) and, in all but one PN, the diffuse X-ray emission lies within the confines of these elliptical inner bubbles (the lone exception is, once again, NGC 7027;

Kastner et al. 2001, 2002). Most of the “hot bubble X-ray” PNe (Section 3.1) also show ansae (Sahai et al. secondary morphology characteristic of “a” in Table 1) associated with bullet-like mass ejections (FLIERS; Balick et al. 1994), and four objects (NGC 40, 2371, 6543, 7009) display axisymmetric and/or point-symmetric structures that are further indicative of fast, collimated flows, leading to classifications of bipolar, multipolar, or “collimated lobe pair” nebulae under the PN classification system of Sahai et al. (2011).

Figure 5 and Table 1 readily demonstrate that (1) the central bubbles within all of the diffuse X-ray PNe have radii  $\lesssim 0.15$  pc, corresponding to dynamical ages  $\lesssim 5 \times 10^3$  yr; and (2) most diffuse X-ray PNe have inferred CSPN effective temperatures  $T_{\text{eff}} \lesssim 100$  kK (the only exceptions thus far being NGC 7027 and 7662). These two observations suggest, respectively, that (1) the timescale for energetic wind interactions in elliptical PNe is  $\sim 5 \times 10^3$  yr; and (2) the luminous X-ray emission arising in wind shocks may contribute to the high-excitation states of the subclass of multiple-shell elliptical PNe that display well-defined central bubbles (provided such nebulae harbor sufficient masses of high-density gas; Ercolano 2009).

The foregoing results reinforce previous assertions by Gruendl et al. (2006), Kastner et al. (2008), and Kastner (2009) that the necessary conditions for detectable diffuse X-ray emission in PNe are either (1) a combination of energetic central star (pre-WD) winds and enclosed inner PN shells (or lobes) that can effectively confine wind-shock-heated plasma, as is the case for all CHANPLANS survey objects apart from NGC 7027; or (2) high-velocity, collimated post-AGB flows impinging on AGB ejecta, as is the case for NGC 7027 (as well as for the more distant objects Mz 3 and Hen 3-1475; Kastner et al. 2003; Sahai et al. 2003). Moreover, in the former case, it appears that the conditions described in the interacting winds scenarios which predict the production of a classical “hot bubble” (e.g., Zhekov & Perinotto 1996) are met only for a limited class of PNe: specifically, elliptical, nested-shell PNe with “young” inner bubbles and ansae. This result appears to underscore the importance of accounting for the rapid time evolution of pre-PN and CSPN wind properties in modeling the key, early stages in the structural evolution of PNe (see, e.g., Villaver et al. 2002; Akashi et al. 2006; Huarte-Espinosa et al. 2012). Additional analysis of the CHANPLANS and archival data obtained to date, combined with further *Chandra* X-ray observations of PNe, should lead to an improved understanding of the potential implications of PN diffuse X-rays for models of the evolution of CSPN temperatures, masses, and winds, as well as the consequences of such CSPN evolution on the surrounding PN.

#### 4.2. Point-like X-Ray Emission from CSPNe

It is apparent from Figure 5 and Table 4 that, like diffuse X-ray emission, point-like (CSPN) X-ray emission is more often associated with molecule-poor than molecule-rich ( $\text{H}_2$ -detected) PNe; furthermore, the majority of PNe with X-ray-luminous central stars have elliptical or round morphologies, according to the Frew (2008) classifications. Specifically,  $\sim 60\%$ – $65\%$  of elliptical (or round), molecule-poor PNe harbor CSPN X-ray sources, whereas only  $\sim 25\%$  of PNe with bipolar morphologies and/or in which near-IR  $\text{H}_2$  has been detected display such CSPN X-ray point sources. It also appears that the majority of PNe hosting X-ray-luminous central stars have morphologies that are perhaps best characterized as amorphous and internally disorganized—in stark contrast to the (generally) highly structured morphologies of PNe that display diffuse X-ray emission

from hot bubbles (Sections 3.1.1 and 4.1) or that lack detectable X-rays (Section 3.3).

The X-ray SEDs of these CSPN X-ray sources appear to represent two general classes (Figure 6, top): (1) objects that display very soft X-ray SEDs, indicative of strong or dominant hot ( $\sim 100$ – $200$  kK) photospheric components (Section 3.2.1), and (2) CSPNe that display harder X-ray SEDs, dominated by photons in the range  $\sim 0.6$ – $1.0$  keV (Section 3.2.2). The former (soft X-ray) CSPN group, all of which have rather high CSPNe  $T_{\text{eff}}$ , are thus far confined to a rather narrow range in PN radius (i.e., radii  $\sim 0.1$ – $0.4$  pc; Figure 6, bottom), corresponding to a short, well-defined dynamical timespan. This suggests that the epoch of significant X-ray contributions from hot CSPN photospheric radiation is both delayed and short-lived. Specifically—setting aside the young, [WO]-type CSPN in NGC 2371—it appears that the epoch of detectable soft X-ray emission from the CSPN photosphere corresponds to a dynamical PN age of  $\sim 10^4$  yr (Table 1), and that such photospheric X-ray emission then declines significantly after another few  $\times 10^3$  yr.

The second (harder X-ray SED) group represents the majority of CSPN X-ray sources. Unlike the soft CSPN point sources, the harder CSPN X-ray sources span a wide range of PN radii, indicating that either a single long-lived process or a combination of short-lived and longer-timescale processes, intrinsic to the PN stellar component, are responsible. There is a broad range of potential explanations for these CSPN sources, including (see Guerrero et al. 2001; Blackman et al. 2001a; Soker & Kastner 2002; Montez et al. 2010; Bilíková et al. 2010, and references therein): coronal emission from late-type binary companions that have been “spun up” (and hence become highly magnetically active) via accretion of mass lost by the PN progenitor, or whose coronae have been compressed by the CSPN wind; post-AGB magnetic activity at the CSPN itself (possibly instigated by interactions with a past or present binary companion); emission arising from an actively accreting companion (e.g., accretion shocks at a main-sequence companion, or an accretion disk associated with a compact companion); re-accretion (“fallback”) of PN material onto the CSPN; or self-shocking, variable, fast CSPN winds analogous to those of massive OB stars. The data available thus far do not particularly favor any one of these alternative models; indeed, it is likely that different mechanisms may apply to different CSPNe. However, there are several points worth noting.

1. The three X-ray CSPNe for which the most likely cause of the X-rays is coronal emission from spun-up companions, DS 1, HFG 1, and LoTr 5 (Montez et al. 2010), are associated with some of the dynamically oldest (largest) PNe in the CHANPLANS sample (Figure 6, bottom). This is consistent with the notion that the spin-down (hence enhanced magnetic activity) timescale for the companions in these systems should be significantly greater than characteristic PN lifetimes ( $\sim 10^5$  yr; e.g., Frew 2008 and references therein).
2. The median energies of the DS 1, HFG 1, and LoTr 5 CSPN X-ray sources are very similar, and are among the hardest thus far detected ( $\sim 1.0$  keV), reflecting the relatively high temperatures of these sources ( $T_X \sim 10$  MK, consistent with coronal emission from late-type companions; Montez et al. 2010). This suggests that other X-ray sources in this same median energy range (i.e., those within NGC 6445 and 7008; Figure 6) may also be due to the coronae of spun-up,



late-type CSPN companions—an assertion bolstered by the disproportionately large fraction of PNe with known binary CSPN that display point-like X-ray sources (Table 4). On the other hand, in the case of the point source within NGC 7293—whose median X-ray photon energy is similar to the other PNe CSPN X-ray sources just mentioned—the presence of a CSPN companion earlier than late M spectral type is excluded (O’Dwyer et al. 2003).

3. The constraints placed on the intrinsic X-ray luminosities of the CSPNe within NGC 650-1 and 2346 by their *Chandra* nondetections are compromised somewhat by the fact that these CSPNe suffer considerable extinction within the equatorial regions of their host bipolar PNe. Such absorption effects are important for the class of very soft (blackbody-dominated) CSPN sources (Section 3.2.1). On the other hand, our detection of the CSPN of the bipolar PN NGC 6445 indicates absorption-related selection effects are less important for sources with median energies  $\gtrsim 0.6$  keV (S 3.2.2). Furthermore, the X-ray-undetected CSPNe of the similarly molecule-rich, “Ring-like” objects NGC 3132, 6720, 6772, 6781—which are lower-inclination analogs to bipolar PNe (Kastner et al. 1994, 1996)—are subject to relatively small line-of-sight absorbing columns. Hence, the initial CHANPLANS results suggest that strongly axisymmetric, molecule-rich PNe generally have X-ray-faint CSPNe.
4. The presence of a binary companion to the CSPN is widely believed to be responsible for bipolar (axisymmetric) structure in PNe (see, e.g., Balick & Frank 2002 and references therein). The apparent lack of X-ray point sources associated with the central stars of molecule-rich, axisymmetric (bipolar and Ring-like) PNe among the CHANPLANS sample (Table 4) would therefore appear to contradict the hypothesis that spun-up companions are widely responsible for “hard X-ray CSPNe.” It is possible that the lack of X-ray-luminous CSPNe within bipolar and Ring-like PNe reflects the fact that these objects are descended from progenitors whose masses are higher than average for PNe ( $\gtrsim 1.5 M_{\odot}$ ; Kastner et al. 1996, and references therein); hence, on average, the central stars of bipolar and Ring-like PNe may have a higher incidence of X-ray-inactive, intermediate-mass binary companions than elliptical and round PNe. Indeed, the nondetection of X-ray sources at the CSPNe of NGC 2346 and 3132 is not inconsistent with these CSPNe harboring wide-separation, intermediate-mass companions; such (A type) stars are generally less magnetically active than late-type stars, due to their lack of envelope convective zones (although it is then noteworthy that the CSPN of NGC 1514, which has an A-type companion, is an X-ray source). Regardless, the relative rarity of CSPN X-ray emission in the case of other bipolar and Ring-like PNe is difficult to explain, if a close companion was responsible for their axisymmetric structures. Perhaps the companion is a magnetically inactive, nonaccreting WD, or has already merged with the CSPN (during, e.g., a common envelope phase). Alternatively, the lack of point-like X-ray emission from CSPNe of bipolar and Ring-like PNe may somehow reflect the more rapid evolution of such objects.
5. The X-ray point sources associated with “hot bubble X-ray” PNe (Section 3.1) display a narrow range of median photon energy ( $\sim 0.5$ – $0.7$  keV; Figure 6, bottom). Though it remains to assess to what extent these median energies are affected by contamination from the underlying diffuse

X-ray emission, their clustering in median energy may suggest a common CSPN X-ray emission mechanism. This mechanism could be internal (small-scale) wind shocks in the near-CSPN environment, given that these PNe all exhibit the effects of ongoing wind collisions at large scales.

## 5. SUMMARY

We are undertaking CHANPLANS, the first systematic *Chandra* X-Ray Observatory survey of PNe in the solar neighborhood. CHANPLANS began with a 570 ks *Chandra* Cycle 12 Large Program targeting 21 (mostly high-excitation) PNe within  $\sim 1.5$  kpc of Earth. We have combined the results of these observations with those obtained from *Chandra* archival data for the (14) other PNe within  $\sim 1.5$  kpc that have been observed to date. The highlights of the early CHANPLANS results include the following.

1. The overall X-ray detection rate for the 35 PNe within  $\sim 1.5$  kpc observed thus far by *Chandra* is  $\sim 70\%$ .
2. Roughly 50% of the sample PNe harbor X-ray-luminous point sources at their CSPNe. This fraction includes nine new detections of CSPNe X-ray sources among the Cycle 12 sample PNe, and another three CSPN point sources identified via analysis of previously unpublished archival data.
3. All but one of the point-like X-ray sources detected at CSPNe display X-ray spectra that are harder than expected from hot cores emitting as simple blackbodies (the lone apparent exception is the central star of the Dumbbell nebula, NGC 6853). These hard X-ray excesses may suggest a high frequency of binary companions to CSPNe. Other potential explanations include self-shocking winds or PN mass fallback.
4. Soft, diffuse X-ray emission tracing shocks (in most cases, “hot bubbles”) formed by energetic wind collisions is detected in  $\sim 30\%$  of the sample PNe. The PNe detected as diffuse X-ray sources include four nebulae imaged by *Chandra* in Cycle 12 (NGC 2371, 3242, 7009, 7662) and three PNe for which archival X-ray images are presented here for the first time (NGC 2392, 6826; IC 418).
5. Five objects (NGC 2371, 2392, 6543, 6826, and 7009) display both diffuse and point-like emission components in *Chandra* imaging.
6. The presence (or absence) of X-ray sources appears correlated with PN density structure: molecule-poor, elliptical nebulae are more likely to display X-ray emission (either point-like or diffuse) than molecule-rich, bipolar, or Ring-like nebulae.
7. In addition to displaying elliptical morphologies, most PNe detected as diffuse X-ray sources have a nested shell/halo structure and display bright ansae; the diffuse X-ray emission regions are enclosed within the innermost, compact, sharp-rimmed shells. All of these inner shells have dynamical ages  $\lesssim 5 \times 10^3$  yr, placing firm constraints on the timescale for strong shocks due to wind interactions in PNe.
8. The central stars of all but two diffuse X-ray-emitting PN—the exceptions being NGC 7027 and 7662—have effective temperatures  $T_{\text{eff}} \lesssim 100$  kK, further reflecting the youth of these objects and suggesting that high-energy emission arising in wind shocks may contribute to the high-excitation states of archetypical “hot bubble” nebulae such as NGC 2392, 3242, 6826, and 7009.

Further analysis of these and future CHANPLANS X-ray imaging spectroscopy data and results describing both point-like and diffuse X-ray emission from PNe will serve to inform and refine models describing PN shaping mechanisms and, in particular, the role of binarity in determining PN structure and evolution.

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