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# The Dark Energy Survey Supernova Program results: Type Ia Supernova brightness correlates with host galaxy dust

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## ABSTRACT

Cosmological analyses with type Ia supernovae (SNe Ia) often assume a single empirical relation between color and luminosity ( $\beta$ ) and do not account for varying host-galaxy dust properties. However, from studies of dust in large samples of galaxies, it is known that dust attenuation can vary significantly. Here we take advantage of state-of-the-art modeling of galaxy properties to characterize dust parameters (dust attenuation  $A_V$ , and a parameter describing the dust law slope  $R_V$ ) for the Dark Energy Survey (DES) SN Ia host galaxies using the publicly available **BAGPIPES** code. Utilizing optical and infrared data of the hosts alone, we find three key aspects of host dust that impact SN Ia cosmology: 1) there exists a large range ( $\sim 1 - 6$ ) of host  $R_V$  2) high stellar mass hosts have  $R_V$  on average  $\sim 0.7$  lower than that of low-mass hosts 3) there is a significant ( $> 3\sigma$ ) correlation between the Hubble diagram residuals of red SNe Ia that when corrected for reduces scatter by  $\sim 13\%$  and the significance of the “mass step” to  $\sim 1\sigma$ . These represent independent confirmations of recent predictions based on dust that attempted to explain the puzzling “mass step” and intrinsic scatter ( $\sigma_{\text{int}}$ ) in SN Ia analyses. We also find that red-sequence galaxies have both lower and more peaked dust law slope distributions on average in comparison to non red-sequence galaxies. We find that the SN Ia  $\beta$  and  $\sigma_{\text{int}}$  both differ by  $> 3\sigma$  when determined separately for red-sequence galaxy and all other galaxy hosts. The agreement between fitted host- $R_V$  and SN Ia  $\beta$  &  $\sigma_{\text{int}}$  suggests that host dust properties play a major role in SN Ia color-luminosity standardization and supports the claim that SN Ia intrinsic scatter is driven by  $R_V$  variation.

**Key words:** cosmology: observations — surveys — galaxies: general — supernovae: general

## 1 INTRODUCTION

Type Ia Supernovae (SNe Ia) have been used for decades as standard candles (Riess et al. 1998; Perlmutter et al. 1999; DES Collaboration 2019; Brout et al. 2022), and they remain a powerful cosmological tool to probe the expansion of the Universe, along with many other measurements, including those from the Cosmic Microwave Background (Planck Col-

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laboration 2018), Baryonic Acoustic Oscillations (Eisenstein et al. 2005), and more recently, gravitational wave standard sirens (Abbott et al. 2017; Soares-Santos, Palmese et al. 2019; Palmese et al. 2021). A deep understanding of the factors that cause variation in a standard candle’s luminosity and colour is paramount to increase the level of accuracy and precision of inferred distances and cosmological parameters.

While photometric calibration has historically been thought of as the dominant source of systematic uncertainty, recent analyses (e.g., Brout et al. 2019b) have found that empirical modelling of the residual intrinsic brightness variations after standardization (‘intrinsic scatter’) presents an equally important source of systematic uncertainty. Note that although this scatter is usually referred to as ‘intrinsic’, it may be caused by effects that are external to the SN itself. It is estimated that because of the current lack of understanding of the physical underpinnings of intrinsic scatter, there exists a significant redshift-dependant systematic uncertainty in the estimated SN Ia distances (Brout & Scolnic 2021) therefore a more complete physical description of SN Ia intrinsic brightness variations is critical for the next generation of SN Ia cosmological analyses.

In this work, we focus on the impact of host galaxy dust and its potential to explain the color–luminosity relation, intrinsic scatter, and the puzzling mass step in SN Ia cosmology. The way dust can affect the observed SN colours and cause residual scatter in the Hubble diagram is as follows. Dust obscures light in the rest frame UV-optical wavelengths and thus affects the brightness and spectral energy distribution of objects observed by optical surveys in two ways. The first is extinction, whereby light from observed objects is obfuscated and reddened by intervening dust particles. Secondly, attenuation includes the effect of extinction, but also accounts for the presence of photons scattering off dust and towards the observer. Attenuation and extinction are characterized by laws that describe the difference in luminosity at a given wavelength between the object observed without attenuation/extinction and the observation we see, i.e. with attenuation/extinction. Inadequately accounting for the differential luminosity due to dust and potential variance in dust properties causes incorrect inferred luminosities in the SN Ia Hubble diagram.

However, it is known that there can exist a large range of both dust attenuation ( $A_V$ ) and dust laws ( $R_V$ ) across galaxies. For example, the Milky Way, Large Magellanic Cloud, and Small Magellanic Cloud encompass a large range of dust laws ( $R_V \sim 2 - 5$ ) (Schlafly et al. 2016; Gao et al. 2013; Yanchulova Merica-Jones et al. 2017); similar variations are seen in 230,000 galaxies in the local universe (Salim et al. 2018). Yet, the currently adopted SN Ia methodology does not allow for this freedom in its use of a single colour–luminosity relation. This presents two problems in SN Ia cosmology.

First, because the SN Ia empirical models and relations are mostly agnostic to the underlying physical properties, one needs an additional empirical model to describe the residual scatter in standardized brightness beyond measurement uncertainties. The ‘G10’ intrinsic scatter model (Guy et al. 2010) prescribes 70% of the residual scatter to coherent variation and 30% to chromatic (wavelength dependent) variation, while the ‘C11’ scatter model (Chotard 2011) prescribes only 25% to coherent variation and 75% to chromatic variation.

The more recent ‘SUGAR’ model (Léget et al. 2020) prescribes intrinsic variations to spectral features.

Second, if dust properties vary in a correlated and systematic manner for different galaxy types, it can manifest as residual correlations in the Hubble diagram. This search for correlations between SN Ia standardized brightness with host galaxy properties has been of large focus in the last decade. Both global and local environmental properties of SN Ia-host galaxies have been shown to correlate with the distance residuals in the Hubble diagram (e.g., Hicken et al. 2009; Sullivan et al. 2010; Lampeitl et al. 2010; Gupta et al. 2011; Smith et al. 2012; Childress et al. 2013; Pan et al. 2013; Johansson et al. 2013; Rigault et al. 2013; Rigault et al. 2015; Rose et al. 2019; Rigault et al. 2020; Kelsey et al. 2020a). Remarkably, the strongest correlation has often been seen to be with a global galaxy property that is hardly directly related to the SN itself, the total stellar mass (although see Rigault et al. 2020 and Kelsey et al. 2020a for discussions of the strength of specific star formation rate (sSFR) and rest frame colour respectively). Host galaxy stellar mass has thus been most commonly included as an ad-hoc correction to the standardized brightness in SN Ia cosmology analyses in the form of a step function, hence the name “mass step.”

The physical origins of these correlations have largely remained elusive and have involved novel and sometimes complex interpretations. For the ‘mass step,’ typically defined around stellar masses of  $M_* \sim 10^{10} M_\odot$ , several possibilities have been explored in the literature. For example, the age of the progenitor could be related to the age of the overall stellar population of the galaxy, which is often related to its stellar mass (Kelly et al. 2010; Childress et al. 2013; Rigault et al. 2013; Graur et al. 2015). It is also possible that SNe Ia come in different subclasses, for example those that originate from sub-Chandrasekhar mass and Chandrasekhar mass white dwarfs (Polin et al. 2019, 2021), which show intrinsically different properties. In addition, recent studies from Boone et al. (2021a,b) have used manifold learning techniques in an attempt to model intrinsic spectral diversity of SNe Ia by using pairs of similar (dubbed “twin”) supernovae, in a process called Twins Embedding. These studies indicate that models accurately predicting spectra of Type Ia SNe are highly nonlinear in their parameters. Using these models can halve the size of the mass step ( $0.040 \pm 0.020$  mag compared to  $0.092 \pm 0.024$  mag) but not render it insignificant, hinting that an additional effect potentially unaccounted for in these models, is also driving the presence of the mass step.

Recently, Brout & Scolnic (2021) (Hereafter BS21) suggest a physically motivated dust-based model can explain the colour dependent intrinsic scatter as resulting from a wide range of dust laws and the mass step as an artifact of different dust distributions in different mass galaxies. Neither of these concepts are particularly novel. For example, dust parameters have been found to correlate even more strongly with other physical parameters of interest, like the stellar mass. Salim et al. (2018) find a positive correlation between the stellar mass and the dust column density of galaxies (as given by the parameter  $A_V$ ). In addition, they find lower mass galaxies tend to have steeper attenuation curves (measured by the parameter  $R_V$ , roughly consistent with the independent prediction by BS21).

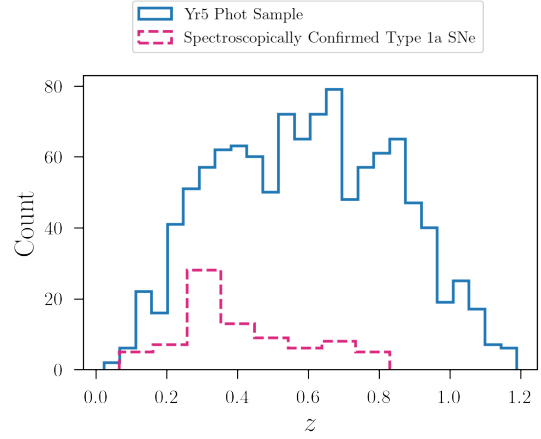
Popovic et al. (2021a) find similar results using SNe Ia with higher certainty using new forward-modeling fitting method

and Popovic et al. (2021b) find that when properly accounting for these host-galaxy correlations with SN properties in simulations, there is no longer a need for ad-hoc corrections. While both BS21 and Popovic et al. (2021a) were able to fit the global dust distributions affecting the SNe Ia, they were unable to distinguish whether the dust properties that they have recovered are intrinsic to the SN (i.e. due to the dust produced by the SN explosion itself), the immediate circumstellar environment, or if they are representative of global galactic properties. In addition, Wiseman et al. (*submitted*) show that the  $R_V$  varying with the age of the host galaxy is also able to account for the observed SN brightness differences when considering SNe split by both their host stellar mass as well as their host colour.

A dust-based SN Ia model could also explain three other key aspects of SN Ia cosmology. First, it could explain why host color and sSFR seem to correlate even stronger with SN Ia Hubble residuals than mass (e.g., Rigault et al. 2020; Kelsey et al. 2020a), as these parameters trace the dust distributions. Second, it could explain why bluer, un-extinguished, type Ia SNe tend to make better standard candles (Foley & Kasen 2011; Gonzalez-Gaitan et al. 2020; Kelsey et al. 2020b; Siebert et al. 2020). Lastly it could explain why observations in the NIR are able to achieve less Hubble diagram scatter and smaller ‘mass steps’ (e.g., Burns et al. 2018; Johansson et al. 2021), as these observations are less sensitive to dust.

Though recent studies have attempted to prove a direct link between dust and the observed mass step (Burns et al. 2018; Ponder et al. 2020; Uddin et al. 2020; Thorp et al. 2021; Johansson et al. 2021), they have not reached definitive consensus as of yet. For example, Ponder et al. (2020); Uddin et al. (2020); Thorp et al. (2021) have suggested evidence for the existence of a mass step in the NIR bands which would invalidate dust-based explanations (e.g. Brout & Scolnic 2021) for the step. However, Burns et al. (2018) and Johansson et al. (2021) find reduced or zero evidence for a mass step in the NIR or when accounting properly for dust in the optical. Furthermore, Johansson et al. (2021) were unable to reproduce the findings of the contradicting papers (Ponder et al. 2020; Thorp et al. 2021; Uddin et al. 2020) and note that these papers may be affected by the biased collection of NIR samples largely at high host stellar mass.

In this work, we aim to test the BS21 model for the effects of dust on SN Ia observations. Furthermore, we seek to test whether the observational contributions of dust are intrinsic/circumstellar or related to the global galaxy dust as whole. We derive galaxy parameters for a subsample of the hosts of the Dark Energy Survey (DES; DES Collaboration 2005, 2016) 5 Year SN sample where state-of-the-art multi-band photometry is available from deep *ugrizY* stacks of Dark Energy Camera (DECam; Flaugher et al. 2015), and from deep observations from the Visible and Infrared Survey Telescope for Astronomy (VISTA; Emerson et al. 2004) in the *JHKs* bands. We note that an analysis of the DES SN hosts from the 3 Year sample has already been carried out in Smith et al. (2020) (here after S20), Wiseman et al. (2020); Kelsey et al. (2020a). Our analysis differs from previous works for a number of reasons. First, we pay particular attention to the dust parameters and their correlation with stellar mass, in an attempt to confirm or rule out the arguments from BS21. Moreover, we use a larger number of filters, in particular optical, including *u*, and NIR data, compared to previous works



**Figure 1.** Redshift distribution of the 1110 DES Year 5 “likely” SN Ia host galaxies (“Phot Sample”, in blue) that overlap with the DES deep fields photometry in Hartley & Choi et al., (2020), as well as the distribution of the subsample of 81 galaxies with spectroscopically confirmed Type Ia SNe in the same fields. Redshifts are from OzDES and various public catalogs.

computing stellar masses with DECam bands (Palmese et al. 2016, 2020; Wiseman et al. 2020; Kelsey et al. 2020a; S20), which helps constrain the dust parameters better compared to optical-only estimates (see e.g., Fig 8 of Childress et al. 2013). However this improved dataset is not available over the whole DES SN fields, and we therefore only analyze a subsample (3 out of 10 fields) of the DES SN hosts, totalling 1110 galaxies.

This paper is organized as follows. In §2 we describe the DES and VISTA data used. In §3 we present the methodology used for the spectral energy distribution (SED) fitting and the priors adopted. §4 and §5 report our results and discussion. We conclude in §6. In Appendix A, we discuss differences between different sets of photometric data used in this work. In Appendices B and C we include alternate versions of figures presented herein, utilizing different input parameters and star formation histories (SFHs).

## 2 DATA

DES (DES Collaboration 2005, 2016)<sup>1</sup> is an optical-near-infrared survey that covered 5000 deg<sup>2</sup> of the South Galactic Cap in *grizY* bands. DES was carried out using the  $\sim 3$  deg<sup>2</sup> DECam on the Blanco 4-m telescope at the Cerro Tololo Inter-American Observatory. The DES Supernova Program (DES-SN) has dedicated shallow and deep drilling fields in *griz* covering 27 deg<sup>2</sup>. SNe are considered as candidates if they satisfy the requirement of two detections at the same location on two separate nights in any of the four bands. A subset of all 5 years of DES-SN candidates were selected for spectroscopic follow-up to obtain a type classification and redshift. A detailed overview of the spectroscopic follow-up program as well as a general overview of the DES-SN program and observing strategy can be found in S20. In total, 415 SNe

<sup>1</sup> [www.darkenergysurvey.org](http://www.darkenergysurvey.org)

have been spectroscopically confirmed as SN Ia. DES-SN executed its first cosmology analysis (Brout et al. 2019b) using scene-modeling photometry from Brout et al. (2019a) for the first 3 years of data (207 spectroscopically confirmed SNe Ia). Here, we utilize difference imaging photometry which is available for all 5 years of SNe Ia and has been shown to be both accurate and achieve similar (within  $\sim 5\%$ ) photometric uncertainties (Brout et al. 2019a) and that the recovered size of the ‘mass step’ is unchanged compared to using scene modeling in the DES-SN-3YR analysis (Brout et al. 2019b). For the difference imaging sample of all 5 years, 290 SNe pass the same cosmology level quality cuts as Brout et al. (2019b). We analyze the subset of 81 spectroscopically classified SNe that are in fields with VISTA infrared bands’ overlap.

A larger sample of 1110 host galaxies is also used in this work in order to study the dust properties of typical SN Ia hosts. To select this sample of hosts we follow the simplest version of cuts defined in Vincenzi et al. (2021) for the DES SN photometric sample. In summary, hosts of ‘likely’ SNe Ia are chosen from the associated sample of unclassified transients from the full 5-year DES-SN sample that, when fit with the SALT2 SN Ia light curve model (Guy et al. 2007), are found to have probability of being a type Ia ( $\text{SNANA FITPROB} > 0.01$  which is determined from the SALT2 data-model  $\chi^2$ ). Additional requirements are that the transients exhibit SN Ia light curve fit parameters  $|c| < 0.5$  ( $c$  being the colour of the SN) and  $|x_1| < 5.0$ , and fitted SN light-curve peak date uncertainty  $< 2$  days. We denote this sample ‘DES 5 year loose photometric sample’, and throughout the text we refer to it as ‘photometric sample’ for simplicity, but note that it differs from the photometrically classified sample in Möller et al. (2022) and that used for DES cosmology as it does not use photometric classification scores. Because for this larger sample of host galaxies we only study the galaxies themselves and not the SNe candidates, this is sufficient for population studies in this work. Most of the results shown here therefore use this larger sample to maximize completeness and statistics, except for those that rely on SN properties (colour) and SN-derived quantities (the Hubble residuals), where we prioritize the purity of the sample and only select spectroscopically-confirmed SNe Ia. Redshifts of all hosts (also in the larger sample) are obtained from the OzDES program on the Anglo-Australian Telescope (AAT; Yuan et al. 2015; Childress et al. 2017; Lidman et al. 2020) or other external catalogs, as described in more detail in Vincenzi et al. (2021). The host matching method used is described in Gupta et al. (2016).

We use the DES Deep Fields photometric catalogue (Hartley et al. 2020) prepared for the Year 3 (Y3) weak lensing and clustering cosmology analyses. Briefly, we complement the DES SN-field data in *griz* with community time DECam data in *u* band and VISTA data from the Deep Extragalactic Observations (VIDEO; Jarvis et al. 2012) and UltraVISTA (McCracken et al. 2012). The Y3 Deep Fields catalogue consists of 2.8 million sources over 4 fields, covering  $\sim 6$  sq deg after masking. One of the fields is in the region of the Cosmic Evolution Survey (COSMOS), which is only used in this analysis for validation tests, while the three of the 10 SN fields covered are C3, E2 and X3 over 4.64 sq. deg. after masking. The depth of the Deep Fields was designed so that the galaxies would have a signal-to-noise  $S/N > \sqrt{10}$  times their equivalent in the main survey. Source detection is performed

in *riz* coadded images, having a  $10\sigma$  depth of 25.7, 25, and 24.3 mag in *r*, *i*, *z* respectively. The VISTA bands have a  $10\sigma$  depth of 23.0–23.8 (*J*), 22.7–23.3 (*H*) and 22.8–22.9 (*Ks*) in  $2''$  apertures, where the ranges take into account the different depths reached in the different fields. Photometry is derived from a bulge plus disk model, fit to the combined *g*, *r*, *i* and *z*-band data using an adapted Multi-Object Fitting (MOF) pipeline, built on the *ngmix* code.<sup>2</sup> Measurements in *u*-band and the VISTA bands, *JHKs*, are made by applying this model in a forced-photometry pipeline with the *fitvd* software.<sup>3</sup> For a full description of the DES Deep Fields data used here, see Hartley & Choi et al., (2020). We compare our results that use the Deep Fields photometry from Hartley & Choi et al., (2020) with those that use the SN coadds in *griz* from Wiseman et al. (2020) and the *uJHK* bands from the Deep Fields. The difference is that the Hartley & Choi et al., (2020) photometry is drawn from coadd images of the best exposures in each region, and in some cases they are contaminated by the SN light if the best exposures happen to be close in time to the SN, leading to very poor SED fits. The photometry from the Wiseman et al. (2020) coadds includes only the seasons that do not contain the SN. We find that, despite the difference in the magnitude estimation between the two methods (Source Extractor – SExtractor, Bertin & Arnouts 1996 – versus MOF), the recovered galaxy properties for the cases in which the Hartley & Choi et al., (2020) photometry is not contaminated are consistent within  $2\sigma$  with each other, as we show in Appendix A. We therefore decide to show results that use the *griz* (Wiseman et al. 2020) and the *uJHK* (Hartley & Choi et al., 2020) photometry, in order to exclude the cases of SN contamination. The *uJHK* photometry comes from observations outside of the nominal DES 5 year (Y5) seasons and therefore they are unlikely to be contaminated by SN light.

The sample considered in this analysis contains 1110 host galaxies. Most of these are relatively bright galaxies, with  $i \lesssim 23$ , and are therefore bright enough to also be detected in the VISTA bands. About 300 of the galaxies do not have a detection in *JHK* as they lie in masked regions of the images. We assume an upper limit given by the VISTA depths for these cases, with a large uncertainty of 10 mag. This sample of 1110 galaxies contains a collection of 81 galaxies that hosted spectroscopically confirmed Type Ia SNe. In Figure 1 we compare the distribution of redshifts in the overall sample versus the spectroscopically confirmed subsample.

Derivation of galaxy properties is described in detail in Section 3, but here we highlight the advantage of combining the VISTA near-infrared photometric data with the DES optical bands. As interstellar dust tends to redden emitted light, near-infrared data is invaluable to estimate dust parameters. The use of VISTA *J*, *H*, and *Ks* strongly improves the precision of our dust parameter measurements as seen in the top panel of Figure 2, where uncertainties on  $A_V$  are reduced for most galaxies. We note that the inclusion of these bands does not cause significant (at more than  $3\sigma$ ) changes in the mean value of the recovered  $A_V$ , as is clear from the bottom panel of Figure 2.

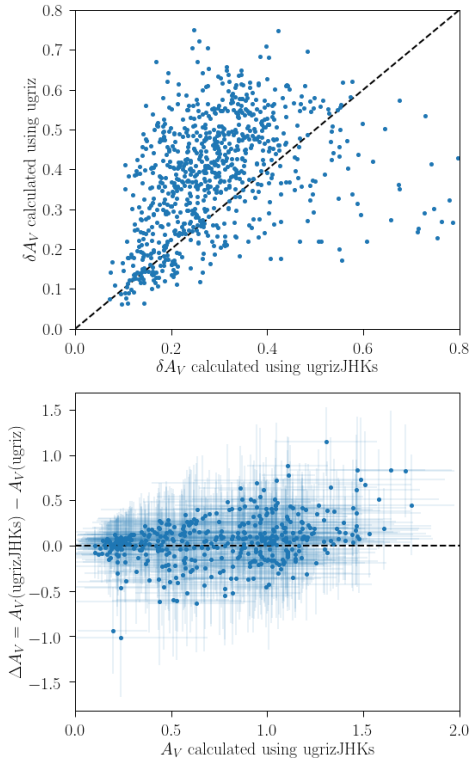
<sup>2</sup> <https://github.com/esheldon/ngmix>

<sup>3</sup> <https://github.com/esheldon/fitvd>



Parameter	Symbol / Unit	Range	Prior	Prior parameters
Redshift	$z$	(0, 1.5)	Gaussian	$(\mu, \sigma) = (z_{\text{spec}}, 0.001)$
Mass Formed	$M_{\text{form}} / M_{\odot}$	(1, $10^{13}$ )	Logarithmic	
Metallicity	$Z / Z_{\odot}$	(0.1, 2.5)	Logarithmic	
V-band dust attenuation	$A_V / \text{mag}$	(0, 2)	Uniform	
Deviation from Calzetti slope	$\delta$	(-1.6, 0.4)	Uniform	
Strength of bump at 2175 Å	$B$	0	Delta function	
log-normal SFH max time	$t_{\text{max}} / \text{Gyr}$	(0, 15)	Uniform	
log-normal SFH FWHM	$\sigma_{\text{SFR}} / \text{Gyr}$	(0, 20)	Uniform	
Exponential SFH age	$t_{\text{age}} / \text{Gyr}$	(0, 15)	Uniform	
Exponential SFH e-folding time	$\tau / \text{Gyr}$	(0, 10)	Uniform	
Burst Age	$t_{\text{age}} / \text{Gyr}$	(0.1, 15)	Uniform	

**Table 1.** Host galaxy parameters estimated using **BAGPIPES**, to which the DECam+VISTA photometry is input. For each parameter, we provide the prior shape and ranges assumed. The top part of the table shows the parameters we use for all cases. The bottom part shows the parameters needed for the SFHs used in this work.



**Figure 2.** *Top:* Uncertainty in the  $A_V$  parameter computed with and without the use of infrared bands. Using the infrared bands consistently leads to lower uncertainties. *Bottom:* residuals between the recovered  $A_V$  mean values with and without the use of infrared bands, as a function of the  $A_V$  computed using all bands.

### 3 METHOD

#### 3.1 Galaxy modelling

In order to derive galaxy properties, we use the Bayesian Analysis of Galaxies for Physical Inference and Parameter ESTimation (**BAGPIPES**; Carnall et al. 2018), a spectral-fitting code used to infer galaxy properties from photometric and

spectroscopic data. **BAGPIPES** employs a fully Bayesian framework and allows us to use sophisticated dust attenuation curves as well as several parametric SFHs. **BAGPIPES** utilizes the stellar population models derived in Bruzual & Charlot (2003) and relies on the **MultiNest** nested sampling algorithm to obtain the posterior distribution based on model, prior distributions and observational data provided by the user.

We experiment with a variety of SFHs to determine their effect on the final estimates of the galaxy properties. In particular, we focus on a log-normal SFH. The log-normal SFH is governed by a peak in star formation rate (henceforth SFR) at some time  $t_{\text{max}}$ , as well as a full width at half maximum parameter,  $\sigma_{\text{SFR}}$ . Both of these parameters are given a uniform prior, over (0, 15) and (0, 20) Gyr, respectively. These broad priors, which follow other works such as Carnall et al. (2019b), let us reproduce a wide range of star formation histories, including those that are flat. We extensively test the effect of these priors on the results. The log-normal SFH is given by the following relation, first presented in Gladders et al. (2013) and revised in Simha et al. (2014):

$$\text{SFR}(t) \propto \frac{1}{t} \exp \left[ -\frac{(\ln(t) - T_{0,t})^2}{2\tau_t^2} \right], \quad (1)$$

where  $t$  is the age of the universe, SFR is the star formation rate, and  $\tau_t$  and  $T_{0,t}$  are free parameters. We follow Diemer et al. (2017) and Carnall et al. (2019b) in redefining these parameters in terms of  $t_{\text{max}}$  and  $\sigma_{\text{SFR}}$  through the following relations:

$$t_{\text{max}} = e^{T_{0,t} - \tau_t^2}, \quad (2)$$

and

$$\sigma_{\text{SFR}} = 2t_{\text{max}} \sinh(\sqrt{2\ln(2)}\tau_t). \quad (3)$$

In addition, we experiment with modifications to the log-normal SFH by adding a delta function “burst” in SFR at some time  $t_{\text{age}}$ , defined as the time since the burst. This SFH is referred to as log-normal + Burst.

The exponential SFH has been widely adopted in the literature (e.g., Tinsley 1972; Bruzual A. 1983; S20). It is parameterized by a start time,  $T_{0,e}$ , where it increases from a star

forming rate of 0 to its maximum value. It then exponentially decays with an  $e$ -folding time of  $\tau$ , following:

$$\text{SFR}(t) \propto \begin{cases} \exp\left(-\frac{t-T_0}{\tau}\right) & t > T_0 \\ 0 & t < T_0. \end{cases} \quad (4)$$

We give  $\tau$  a uniform prior of (0, 10) Gyr in general, but we will occasionally use a (0, 2) Gyr uniform prior to compare to other analyses. We give  $T_0$  a broad uniform prior over (0, 14) Gyr.

For all SFH models, we assume uniform priors in logarithm space ranging over  $(1, 10^{13}) M_\odot$  (i.e., an uninformative prior) and (0.1, 2.5)  $Z_\odot$  for the total stellar mass formed and for the metallicity  $Z$ , respectively. The redshift prior is a Gaussian centered at the spectroscopic redshift  $z_{\text{spec}}$  for each SN, and has a standard deviation of 0.001.

Dust attenuation curves can be parameterized in several ways, and they are often described by the parameters  $A_V$  and  $R_V$ .  $A_V$  is the dust attenuation in V band (at  $\lambda \simeq 5500\text{\AA}$ ). It is often used as an overall normalization constant for the attenuation curve, and is proportional to dust column density. In other words, the more dust between the observer and the observed point, the more dust affects the attenuation curve at all wavelengths.  $R_V$  determines the slope of the attenuation curve in the optical wavelengths.  $R_V$  is defined as  $\frac{A_V}{E(B-V)}$ , the ratio between  $A_V$  and the reddening  $E(B-V)$ , which is the difference in extinction in the B and V bands. For our analysis we use the dust attenuation law from Salim et al. (2018), as also done in Carnall et al. (2020), governed by the parameters  $A_V$ ,  $\delta$  and  $B$ . The  $\delta$  parameter represents the deviation from the Calzetti attenuation curve model Calzetti et al. (1994) and is directly related to  $R_V$  via:

$$R_V = \frac{R_{V,\text{Cal}}}{(R_{V,\text{Cal}} + 1) (4400/5500)^\delta - R_{V,\text{Cal}}}, \quad (5)$$

where  $R_{V,\text{Cal}} = 4.05$ . It should be noted that this relationship is only valid for  $E(B-V) = 1$ , and thus should be considered only an approximation for other reddening values. The  $B$  parameter corresponds to the strength of the ‘‘bump’’ to the total attenuation at  $2175\text{\AA}$ .  $B$  is fixed at 0 in our case, so it does not have any effect on the dust law and we do not take it into account in the equations above, nor in the following. This choice is driven from the fact that we are not particularly interested in this parameter, and adding another degree of freedom to the extinction law could be problematic considering the limited number of filters available. We perform a test run leaving  $B$  free, and no results are altered enough to be noteworthy.  $A_V$  is given a uniform prior over (0,2) while  $\delta$  has a uniform prior over (-1.6, 0.4) (following Salim & Narayanan 2020, who revisit the bounds used in Salim et al. 2018) which corresponds to  $R_V$  values in a range (1.2, 7.1). Parameters and priors used in this work are summarized in Table 1.

The ability of BAGPIPES to recover galaxy properties, including dust, has been extensively studied in Carnall et al. (2018) and Carnall et al. (2019b) using simulations. Carnall et al. (2019a) finds trends of galaxy properties with dust attenuation parameters that are in agreement with independent studies. We also perform a large number of validation tests using COSMOS galaxies, and comparisons with the catalogs of Laigle et al. (2016); Weaver et al. (2022), and with Salim et al. (2018).

In this work, we take  $\chi^2 < 12.5$  as our standard SED fit

cutoff, as this value corresponds to a p-value of  $< 0.1$  for 7 degrees of freedom. Out of the 1110 galaxies in the SN loose photometric sample, 870 pass this cut.

### 3.2 Supernova Distances

Our analysis of the spectroscopically-confirmed DES SN Ia light curves (i.e. a subsample of the 1110 SNe that is used in the analyses that rely on Hubble residuals and SN properties) follows that of Brout et al. (2019b). In summary, we fit light curves with the SALT2 model as presented in Guy et al. (2010) and updated in B14 to obtain light curve fit parameters ( $m_B$ ,  $c$ ,  $x_1$ ). Distance moduli ( $\mu$ ) are then inferred with the Tripp estimator (Tripp 1998) for each SN Ia by:

$$\mu = m_B + \alpha_{\text{SALT2}}x_1 - \beta_{\text{SALT2}}c - M + \delta_{\text{bias}} \quad (6)$$

where  $m_B$  is peak-brightness based off of the light-curve amplitude ( $\log_{10}(x_0)$ ) and where  $M$  is the absolute magnitude of a SN Ia with  $x_1 = c = 0$  and noting that  $M$  is degenerate with the determination of  $H_0$ .  $\alpha_{\text{SALT2}}$  and  $\beta_{\text{SALT2}}$  are the correlation coefficients that standardize the SNe Ia and are determined following Marriner et al. (2011), in a similar process to what is done in B14. Marriner et al. (2011) minimize a  $\chi^2$  expression that depends on the Hubble residuals after applying the Tripp estimator (see Eq. 6) and normalize residuals by the quadrature sum of the measurement uncertainties and intrinsic scatter. The method separates the sample into redshift bins in order to remove the cosmological dependence of the fitting procedure. The procedure iterates to determine the intrinsic scatter  $\sigma_{\text{int}}$ ,  $\alpha$ ,  $\beta$  and the resultant distance modulus values.

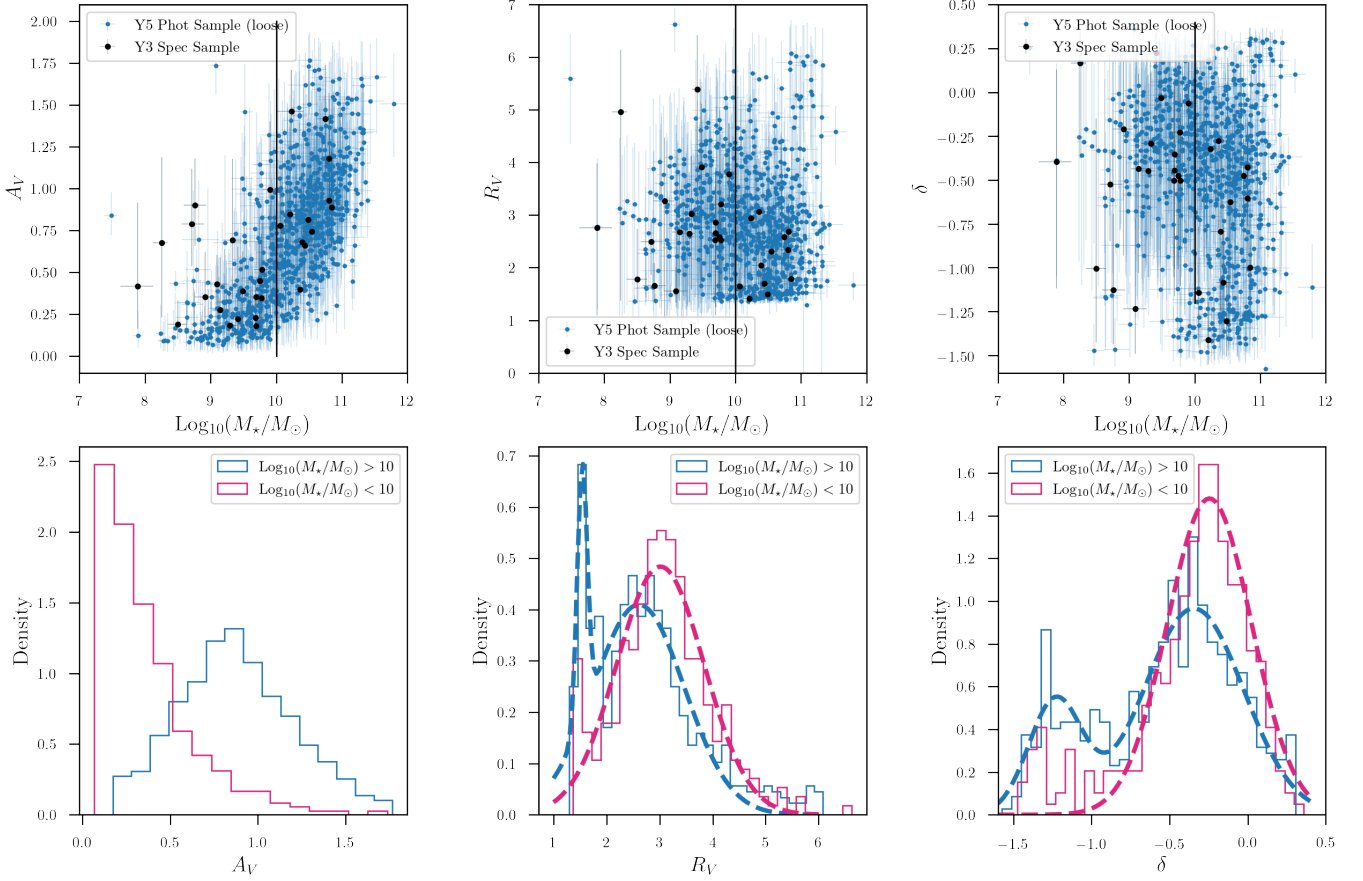
The observed SN Ia distances of Eq. 6 are bias corrected by the expectation ( $\delta_{\text{bias}}$ ) from simulations of the survey. We employ one crucial difference with the method used in Brout et al. (2019b) in accordance with the result found by S20 that the 5-dimensional bias correction ( $c, x_1, m_B, \alpha, \beta$ ) of Kessler & Scolnic (2017) artificially suppresses the size of the observed correlations between Hubble diagram residuals and host galaxy properties. For this reason we employ a solely redshift-dependent Malmquist bias correction (1D) similar to that of Betoule et al. (2014). While Kessler & Scolnic (2017) found that 1D  $\delta_{\text{bias}}$  can result in subtle cosmological parameter inaccuracies, we utilize the 1D method because it preserves the observed correlations between light-curve fit parameters ( $c$ ,  $x_1$ ) and host galaxy properties, which is the focus of this paper.

## 4 RESULTS I: HOST PROPERTIES

In this Section we show the fitted parameters and resulting properties of the DES host galaxies in the loose photometric sample with VISTA overlap. We examine the relations between dust parameters and mass properties themselves as this has the potential to explain the ‘mass step’. Finally we explore the relation between host color and the inferred dust parameters.

### 4.1 The diversity of host galaxy dust parameters

In Figure 3, we show the recovered distributions of  $A_V$ ,  $R_V$ , and  $\delta$  values. We find a large range of  $A_V$ ,  $\delta$ , and  $R_V$  values



**Figure 3.** *Top:* Dust parameters  $A_V$ ,  $R_V$  (as defined in Eq. 5), and  $\delta$  estimates from this work as a function of stellar mass, derived using a log-normal SFH. Blue points represent galaxies whose photometry data are found in the DES Year 5 sample (Y5 Phot Sample loose), black points are those galaxies found in the DES Year 3 spectroscopic Sample. In all cases we use 3 of the SN fields. In all six plots, galaxies that did not meet a  $\chi^2$  cut are not shown and are not used in calculation of fits. The vertical line denotes the location of the  $10^{10} M_\odot$  split. *Bottom:* Distributions of  $A_V$ ,  $R_V$ , and  $\delta$  values above and below the  $10^{10} M_\odot$  split, the location of the observed mass step. For the  $R_V$  and  $\delta$  distributions, we include Gaussian curves to highlight the distributions' dependence on mass.

Gaussian fit to:	$\mu$	$\sigma$	$N$
$R_V$ High $M_*$ 1	$2.61 \pm 0.07$	$0.86 \pm 0.06$	$0.88 \pm 0.02$
$R_V$ High $M_*$ 2	$1.54 \pm 0.02$	$0.10 \pm 0.01$	$0.12 \pm 0.02$
$R_V$ Low $M_*$	$3.02 \pm 0.05$	$0.82 \pm 0.04$	1
$\delta$ High $M_*$ 1	$-0.35 \pm 0.03$	$0.32 \pm 0.03$	$0.77 \pm 0.04$
$\delta$ High $M_*$ 2	$-1.23 \pm 0.04$	$0.17 \pm 0.04$	$0.23 \pm 0.04$
$\delta$ Low $M_*$	$-0.25 \pm 0.02$	$0.27 \pm 0.01$	1

**Table 2.** Gaussian approximation to the  $R_V$  and  $\delta$  distributions shown in Figure 3 for low and high mass galaxies split at  $10^{10} M_\odot$ . The high mass galaxies are described by two Gaussian distributions. The parameters  $\mu$  and  $\sigma$  are the mean and standard deviation respectively, while  $N$  is a normalization constant ensuring that the total integral of the double Gaussian equal to one.

for each host, specifically for  $\delta$  mostly between values of  $-1.2$  and  $0$ , or  $1$  and  $4$  for  $R_V$ . This is a wide range in comparison to the distribution found in the Milky Way,  $R_V$  between  $2.9$

	Lower bound	Median	Upper bound
High Mass $A_V$	0.41	0.88	1.35
Low Mass $A_V$	0.00	0.33	0.54
High Mass $\delta$	-1.03	-0.50	0.17
Low Mass $\delta$	-0.72	-0.34	0.40

**Table 3.** Lower and upper 68% HDI values and medians for the dust parameters  $A_V$  and  $\delta$  and for both high and low mass galaxies.

and  $3.3$  (Schlafly et al. 2016). Our large range found in  $R_V$  agrees with determination of a large  $R_V$  spread ( $1$ - $5$ ) found in Brout & Scolnic (2021).

We note that typical values of  $R_V=3$  are often assumed for SN analyses or have found to be tightly peaked in  $R_V$  as seen in Mandel et al. (2020) and Thorp et al. (2021). In particular, Thorp et al. (2021) find that global  $R_V$  values lie within the range  $R_V = 2.61 \pm 0.21$ , a distribution about ten times narrower than ours. It is however to be noted that the  $R_V$  estimation of Thorp et al. (2021) comes from the SN light



curve, while our determination is a global property of the host galaxy. A large spread is also found for a more generic galaxy population such as that of [Salim et al. \(2018\)](#), finding  $R_V$  ranging from 1 to 10. Values of  $R_V$  that differ from 3.1 have also been measured from the SN data (as opposed to their hosts), e.g., [Foley & Kasen \(2011\)](#).

Novel galaxy evolution studies have shown there is a large spread in  $R_V$  for different galaxies, and we show, for the first time, that this spread exists in a population of SN hosts in this work. First, deviations from  $R_V = 3.1$  or 4.05 have been found over the past decades using various galaxy samples (e.g., [Conroy et al. 2010](#); [Salmon et al. 2016](#); [Salim et al. 2016, 2018](#)). Moreover, [Salim & Narayanan \(2020\)](#) find a range of  $\delta$  values which is similar to what found here. They show that there exists a correlation between stellar mass and slope of the dust law, as we find. The fact that we reach similar conclusions to what is well known in the galaxy evolution literature motivates the use of these findings to improve SN cosmology analyses, as we show in the following.

## 4.2 Dust Parameters and correlation with host mass

### 4.2.1 Correlation between $R_V$ and host mass

The correlation between  $R_V$ , and equivalently  $\delta$ , and stellar masses is shown in the top middle and right hand side panels of Figure 3. Splitting the data at  $\log_{10}(\frac{M_*}{M_\odot}) = 10$ , yields two slightly different distributions of  $R_V$ , where higher mass galaxies have an additional peak around  $R_V \approx 1.7$  compared to the lower mass galaxies, see bottom middle panel of Figure 3. The possibility that lower mass galaxies may have, on average, higher  $R_V$  values and lower  $A_V$  values compared to higher mass galaxies, has been noticed in [Salim et al. \(2018\)](#) for a generic sample of galaxies. Here, we have confirmed that this is relevant also for a specific selection of SN host galaxies.

We then quantify the validity of our conclusion that higher and lower mass hosts defined at the mass step correspond to lower and higher  $\delta$  (or  $R_V$ ) galaxies, respectively. Utilizing a Kolmogorov–Smirnov (KS) test, we assess the validity of the hypothesis that the  $\delta$  (or  $R_V$ ) distribution for lower versus higher mass SN hosts come from the same distribution. We repeat this test for different values of stellar mass, changing the location where we divide the hosts sample into two. While the absolute minimum of the p-value occurs at  $\log_{10}(\frac{M_*}{M_\odot}) \sim 9.7$ , the distribution for the p-statistic is almost completely flat from  $\log_{10}(\frac{M_*}{M_\odot}) \sim 9.6$  to 10.3. Thus we conclude that the location of the mass step around  $\log_{10}(\frac{M_*}{M_\odot}) \sim 10$  corresponds to two different distributions of  $R_V$  values. This difference is in line with the expectations from BS21 that higher mass galaxies have lower  $R_V$  values, which could explain the mass step.

### 4.2.2 Correlation between $A_V$ and host mass

In the left panels of Figure 3 we show the relation between  $A_V$  and the stellar mass of SN host galaxies, for which we find a strong correlation.  $A_V$  values for host galaxies with  $\log_{10}(\frac{M_*}{M_\odot}) \geq 10$  are higher on average, and cover a larger range of values, than those below  $\log_{10}(\frac{M_*}{M_\odot}) < 10$ . This strong correlation is not surprising, and the fact that our findings are similar to that found in previous galaxy evolution studies

is a confirmation of the robustness of our galaxy property estimates, and that similar correlations between dust extinction and stellar mass also hold for SN hosts. For example, [Garn & Best \(2010\)](#); [Zahid et al. \(2013\)](#) find that the main correlation between galaxy properties and dust attenuation in galaxies is that with the stellar mass, as galaxies building up their stellar mass over time, also build up their dust content. The same trends have been recovered by several more recent studies, including [Zahid et al. \(2017\)](#).

We note that the  $A_V$  presented here is measured from the host galaxy, not from the SN. Since this is measured as a global parameter using photometry from the entire host galaxy, its value can be different and usually larger than that measured from the SN itself. Our  $A_V$  estimate in fact accounts for the dust along the line of sight crossing the entire galaxy, averaged over its light profile, while a generic SN in our sample is more likely to be positioned at some intermediate position along the line of sight within the galaxy, and not close to the nucleus. We find a mean value of  $A_V$  of 0.7, with the low mass galaxies being close to 0.4 and the high mass galaxies averaging at 0.9. On the other hand, the typical  $A_V$  spread from SNe Ia can be approximated by the term  $\beta c$  in the Tripp relation in Eq. (6). Considering a typical color distribution with standard deviation 0.1 and  $\beta \sim 3$  ([Scolnic et al. 2018](#)), the  $\beta c$  term has a standard deviation of 0.3, which is quite different from the two  $A_V$  distributions we find in this work. A comparison with the mean value of  $A_V$  from the SNe is not as straightforward as a spread comparison, as the mean is taken out in SALT2 using an empirical relation.

We estimate the location of the peak of the core part of the distributions in  $R_V$  and  $\delta$  for bins of stellar mass split at  $10^{10} M_\odot$ , with results reported in Table 2. The distributions are binned with Poisson errors. For the low mass bin we use a single Gaussian, while for the high mass bin we assume a double Gaussian, where each Gaussian is normalized with its own normalization factor  $N$  summing up to 1 for the overall distribution. Let us first compare the higher  $R_V$  and  $\delta$  population. We find that the higher mass galaxies present a peak at slightly lower  $R_V$  and  $\delta$  values (2.6 and  $-0.3$ , respectively) than their lower mass counterparts (which peak at 3 and  $-0.2$ ). The standard deviation is slightly smaller for the lower mass galaxies (0.8) compared to the higher mass bin (0.9). The lower  $R_V$  or  $\delta$  population at  $R_V \lesssim 2$  or  $\delta \lesssim -0.9$  only resembles a Gaussian consistently for any mass split from  $\log_{10}(\frac{M_*}{M_\odot}) = 9.6$  to 10.3 for the higher mass galaxies, which is why we only include this secondary Gaussian for this mass bin. The peak is found to be at  $R_V = 1.5$  or  $\delta = -1.2$ , with a standard deviation of 0.1 and 0.2 respectively. The lower  $R_V$  and  $\delta$  population of galaxies comprises a significant 10-20% of the high mass galaxies, as shown by the normalization value  $N$  in Table 2. These results are stable for variation of the mass split location between  $\log_{10}(\frac{M_*}{M_\odot}) = 9.6$  and 10.3. We remind the reader that the  $R_V$  distribution findings should be used with care as  $R_V$  is estimated through an approximation. The  $\delta$  results should be considered instead.

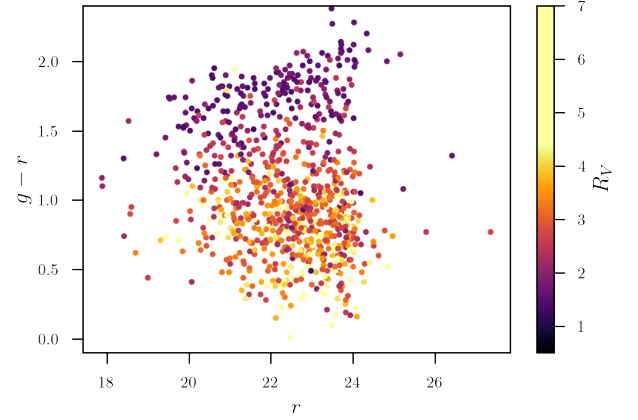
The choice of  $\delta$  prior we make in the SED fitting is similar to that in [Salim & Narayanan \(2020\)](#), and results in a small pile up of galaxies at the low-end of the prior (see rightmost column in Figure 3). This effect appears to be somewhat more prominent in the  $R_V$  case (see middle column in Figure 3) compared to  $\delta$ , but this is expected because  $\delta$  is an expo-

nent in Eq. 5. Moreover, since our  $R_V$  estimation is only an approximation for  $E(B - V) = 1$ , these results are better understood with  $\delta$ , which we directly estimate. As a test, we extend the prior range on  $\delta$  to lower values to approximate  $R_V \sim 0$ . We find that the resulting  $\delta$  distribution ends up extending smoothly to lower values, but our conclusions on the  $A_V$  and  $\delta$  distributions being different for high versus low mass galaxies, and on lower mass galaxies having smaller  $A_V$  and larger  $\delta$  than higher mass galaxies, are largely unchanged.

It is important to quantify the impact of uncertainties in the dust parameters estimates in order to understand whether the difference between high and low mass  $A_V$  and  $\delta$  (or equivalently,  $R_V$ ) distributions is significant. The average of  $\sigma_{A_V}$  is 0.3, meaning that galaxies at the peak of the low mass  $A_V$  distribution are  $\sim 3\sigma$  away from the peak of the high mass distribution. The average of  $\sigma_\delta$  is 0.4, and  $\sim 1$  for  $\sigma_{R_V}$ . This implies that galaxies with a  $\delta$  (similarly for  $R_V$ ) estimate around the peak of the low mass galaxies distribution at  $\delta = -0.2$  are statistically different at  $\sim 2.5\sigma$  from those in the secondary Gaussian of the high mass sample around  $\delta = -1.2$ . We therefore expect differences in the distributions to be significant despite the large measurement uncertainties. A test to verify this consists in using the full posterior distributions on the dust parameters. In Table 3 we show the 68% Highest Density Interval (HDI) for the distributions of  $A_V$  and  $\delta$  that take into account the measurement uncertainties. We construct these distributions by taking 500 samples from the output posteriors of the dust parameters for each galaxy. We provide distributions for  $A_V$  and  $\delta$  using this method for both high and low mass galaxies in Table 3. Applying a KS test to both of these distributions yields p-values much less than 0.05, indicating it is highly likely these two samples were drawn from two different fundamental distributions. Thorp et al. (2021) find that their fitted  $R_V$  is consistent between low-mass and high-mass, contrary to our findings and the findings of Johansson et al. (2021). We note that when measurement uncertainties are taken into account, the 68% HDI for higher mass galaxies occupies lower values in the  $\delta$  parameter space compared to lower mass galaxies, similarly to the case of  $\delta$  point estimates. This supports the arguments of BS21. When it comes to the spread of the distributions, however, we are not able to argue if one of the two populations has a larger scatter. In other words, with the current precision on dust parameter estimates and host statistics, we are unable to quantify how much of the spread in the distributions is given by intrinsic variations in dust parameters rather than measurement uncertainties. In Table 3, we report the specific boundary values of the highest density intervals, as well as the median of the distributions.

## 5 RESULTS II: CORRELATIONS BETWEEN HOST PROPERTIES AND SN IA PROPERTIES

Here we examine the relation between host properties and the SN Ia properties. First, we explore the relation between fitted  $R_V$  of the host and fitted  $\beta_{SALT2}$  from the SNe Ia on similar populations of galaxies. Then, we show correlations between host properties and the Hubble residuals from the SNe in the DES5YR spec sample. Lastly, we show the impact of dust



**Figure 4.** Color-magnitude plot of the hosts considered in this work, color-coded by their best-fit dust law  $R_V$  value. The red sequence galaxies (such as the redMaGiC galaxies) clearly represent the largest fraction of low  $\delta$  (or similarly,  $R_V$ ) galaxies, meaning that they tend to have a steeper dust law slope.

Host Color Split	Host $\langle R_V \rangle$	Host $\sigma_{R_V}$	SN Ia $\beta_{SALT2}$	SN Ia $\sigma_{int}$
$g - r > 1.25$	1.82	0.47	$2.09 \pm 0.15$	$0.05 \pm 0.015$
$g - r < 1.25$	3.15	0.90	$2.71 \pm 0.11$	$0.10 \pm 0.007$

**Table 4.** Fitted distributions of  $R_V$  when split on galaxy  $g - i$  color. Mean  $R_V$  and  $1\sigma$  width  $\sigma_{R_V}$  are reported.

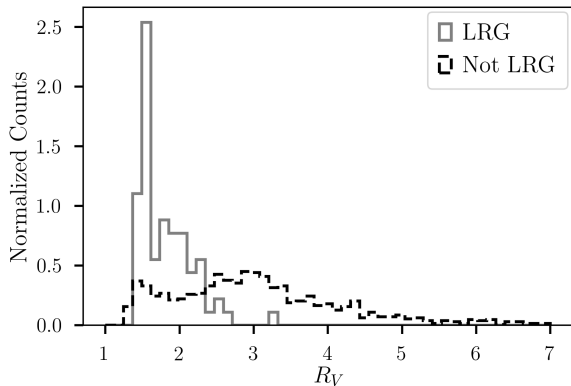
based corrections to the SN Ia standardized luminosity and its effect on the inferred mass step.

### 5.1 Host Dust Drives the SN Ia Color-Luminosity Relation

In order to explore the cause of the low  $R_V$  values seen in Figure 3, the recovered host  $R_V$  values are plotted as a function of host color. In Figure 4, it can be seen that there is a dependence of  $R_V$  with a split occurring roughly at host color of  $g - r = 1.25$ . For host colors redder than this split the  $\langle R_V \rangle = 1.8$  and bluer than the split  $\langle R_V \rangle = 3.2$  (shown in Table 4). This follows the expectation that steeper dust law slope produces redder SEDs, and that red sequence galaxies are those with a lower  $\delta$  or  $R_V$ . On the other hand, the very high  $\delta$  or  $R_V$  galaxies are very blue, star forming galaxies.

For the subset of DES spectroscopically confirmed SNe Ia in the entire DES Y5 spectroscopic sample (i.e. not only those with VISTA bands analyzed in this work, since here we are not using a galaxy SED fit) of red (81 SNe) and blue (209 SNe) hosts, we perform two separate fits of  $\beta_{SALT2}$  when splitting on host  $g - r$  color = 1.25. For the SNe Ia in red hosts we find  $\beta_{SALT2} = 2.09 \pm 0.15$  and for the SNe Ia in other hosts we find  $\beta_{SALT2} = 2.71 \pm 0.11$ . This 3.3 sigma difference in  $\beta_{SALT2}$  is consistent with the difference found in dust  $R_V$  for the same split on host color. We note that  $\beta_{SALT2}$  is defined for  $B$  band magnitudes and that  $\beta_{SALT2} \sim R_B = R_V + 1$ .

We perform the same test using the subsample of the 870 host galaxies that is identified as luminous red galaxies (LRG) by the redMaGiC (Rozo et al. 2016) algorithm applied to the DES wide field survey. redMaGiC identifies red galaxies based on the redMaPPer (Rykoff et al. 2014) red sequence tem-



**Figure 5.** Fitted  $R_V$  for galaxies that are matched to the DES red-MaGiC catalog (solid grey) and those that are not (dashed black). These galaxies are LRGs selected using the redMaGiC algorithm.

plate, given a minimum luminosity cut. Figure 5 shows the  $R_V$  distribution for the 79 LRG galaxies in our sample (solid line) and for all other galaxies (dashed). From Figure 5 it is evident that the LRG sample is found to typically have much lower values of  $R_V$ . For the set of SNe in LRGs, Chen et al. (2022) find a very low value of SALT2  $\beta \sim 2$  instead of the typical  $\beta \sim 3$  that has found in past cosmological analyses of SNe Ia from unselected samples (e.g., Brout et al. 2019b).

These tests mark a concrete advancement over BS21, which was not able to discriminate if the dust that causes a diversity in reddening behavior was due to circumstellar or galactic dust; by connecting the recovered low value of  $\beta$  with a predicted low value of  $R_V$ , this finding points towards galactic dust as a driving force of SN reddening.

## 5.2 Host Dust Drives Residual SN Ia Scatter

From the same test of splitting on host color (at  $g-r = 1.25$ ) described in Section 5.1, this time we consider the spread in  $R_V$  above and below the split and its relation to SN Ia intrinsic scatter in distance. For host colors redder than this split, the  $1-\sigma$  width of the  $R_V$  distribution ( $\sigma_{R_V}$ ) is 0.47 and for hosts bluer than the split  $\sigma_{R_V}$  is 0.90, roughly twice as wide (Shown in Table 4).

The scatter in the SN Ia standardized distances under the same host galaxy conditions is shown in Table 4. We find that for the SN Ia in red hosts the intrinsic scatter ( $\sigma_{\text{int}}$ ) is  $0.05 \pm 0.015$  mag whereas for the blue hosts we find the SN Ia  $\sigma_{\text{int}}$  is  $0.10 \pm 0.007$  mag. These values are distinctly different, at  $3\sigma$  significance. The scatter of the SNe Ia is twice as small for the sample in red hosts in comparison to blue hosts, which is consistent with the scatter found for  $R_V$ , suggesting that dust is indeed driving the scatter in the SN Ia luminosities.

## 5.3 Host Dust Corrections for Current SN Ia Analyses

In this subsection we seek to identify correlations between various host galaxy dust parameters and the SN Ia Hubble Residuals ( $\Delta\mu$ ) to be made as ad-hoc corrections in current analysis methodologies. We note that the residuals here are only computed for the subsample of 81 SN Ia candidates from DES Year 5 that have been spectroscopically classified and

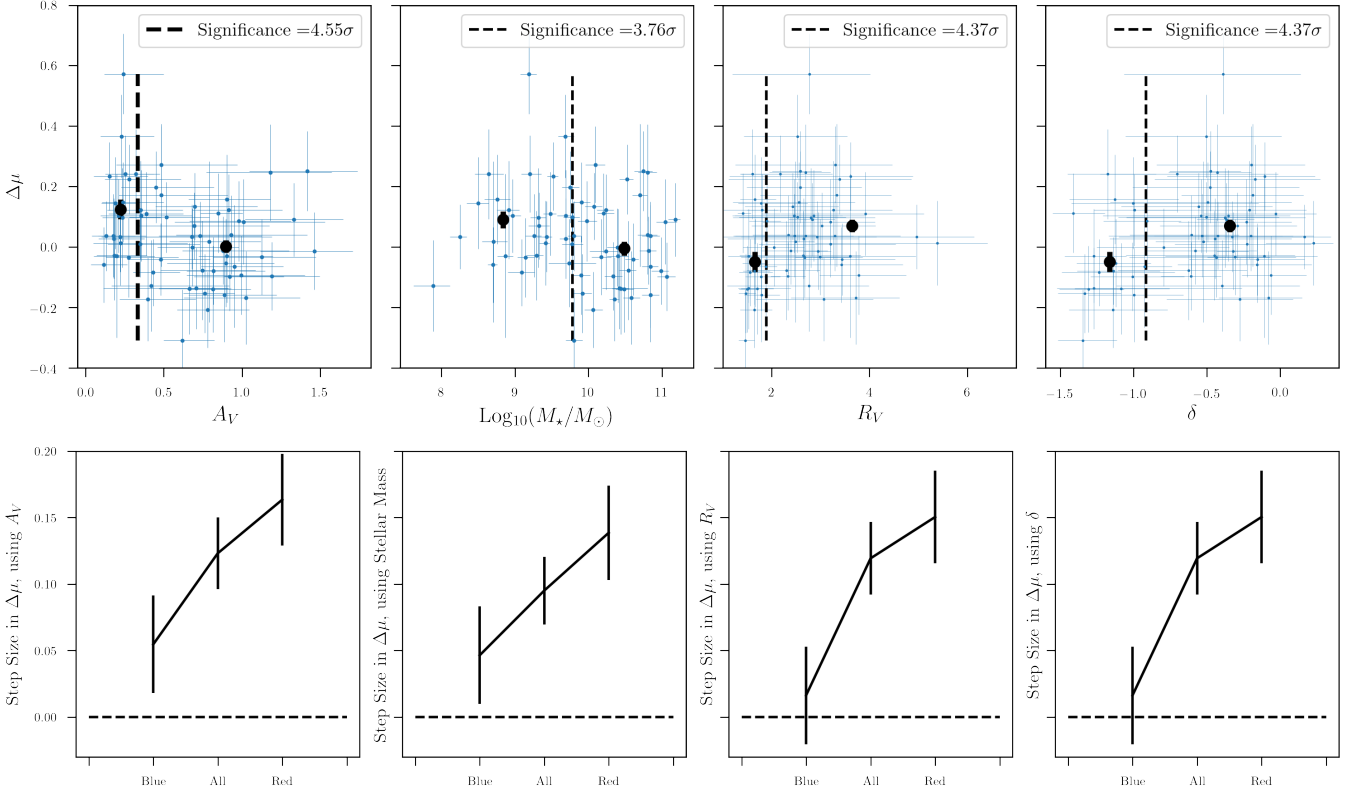
that fall within the DES deep fields with available VISTA data.

In the top panels of Figure 6 we analyze the step-like properties in the Hubble Residuals with respect to several host-galaxy parameters. We scan over the range of each parameter, finding the difference in the mean on either side of a step, and reporting results for the step position that maximizes significance. All steps considered here (i.e. stellar mass,  $A_V$  and  $\delta$  - which is equivalent to  $R_V$ ) are found to be significant and are plotted in black in the top row of Figure 6. We note that the step in dust parameters is as significant as the step in stellar mass, and even more significant in the case of  $\delta$  or  $R_V$ .

In order to further understand the role of dust, we divide the SNe into red and blue SNe, by splitting them in colour at  $c = -0.025$ . Brout & Scolnic (2021) determined that this value separates the SNe that are significantly affected by dust (red sample) from those that are closer to their intrinsic colour (blue sample). We decided to adopt this split rather than trying to model the entire sample as a whole for simplicity: we do not expect the blue SN sample to be significantly sensitive to the dust properties and so we expect it to mostly add noise when measuring relations between Hubble residuals and dust parameters. We report the steps in host properties for blue and red SNe separately in the bottom row of Figure 6. For all galaxy parameters, the step in Hubble residuals for the blue SN sample is not significant ( $< 2\sigma$ ) while it is significant for the red SN sample, further suggesting that it is the SNe extinguished by dust that drive the observed SN Ia relations with host properties. Dixon et al (in prep) use host galaxy spectra instead of photometry to derive galaxy parameters and find similar trends between Hubble residuals and host galaxy dust, which when split by SN Ia color, are only present in the redder SNe Ia.

In addition, we specifically analyze a linear correlation between the dust law slope parameter  $\delta$  (and  $R_V$ ) with Hubble Residuals ( $\Delta\mu$ ). When performing a cut on SNe based on colour at  $c = -0.025$  as done above, a significant linear correlation between  $\delta$  or  $R_V$  and  $\Delta\mu$  becomes apparent (Top of Figure 7). For the  $\delta$  relation with the red SNe Ia ( $c > -0.025$ ), the slope and  $1\sigma$  confidence interval is  $m = 0.219^{+0.076}_{-0.069}$ , and the intercept is  $b = 0.179^{+0.058}_{-0.053}$ , using an MCMC linear fitting algorithm. The maximum likelihood fit is  $\sim 3\sigma$  away from a relation with  $m = 0$ , and we therefore deem the correlation significant. For blue SNe ( $c < -0.025$ ) we find that the correlation is absent in this sample:  $m = 0.033^{+0.052}_{-0.051}$ ,  $b = 0.063^{+0.036}_{-0.035}$ , as can be seen from the bottom panel of Figure 7. We stress that the correlation found for the red SNe is not a result of outliers skewing the overall slope. Repeating the fit with outliers removed (any galaxy that was an outlier in  $\delta$ ,  $\Delta\mu$ , or colour was removed, and one  $3\sigma$  SN Ia  $\Delta\mu$  outlier was removed) achieved consistent results. This result is compelling as it had not been detected in previous works: for example, Pan et al. (2013) find no correlation between SN colour and reddening, and conclude that extinction is not a main cause of the diversity of SN properties and of the mass step.

We remind the reader that the  $R_V$  or  $\delta$  we compute here reflect a global quantity from the galaxy, and will not fully describe the change in SN colour compared to its intrinsic value: there are other factors in play, including the dust at the location of the SN, and Circumstellar Material (CSM) interactions. On the other hand, it is worth stressing that



**Figure 6.** *Top:* Step-like behaviour in Hubble residuals as a function of  $A_V$ , stellar mass,  $R_V$  and  $\delta$ , calculated using a log-normal SFH. We scan over each parameter to search for a significant step in the Hubble Residuals, i.e. a certain point where the mean of the Hubble residuals on either side of said point is significantly different. The most significant step found is displayed in each panel, along with its  $1\sigma$  uncertainty on the mean in black. The location of the splitting that gives the most significant step is displayed as a vertical dashed line. *Bottom:* Step size in the Hubble Residuals as we move from what we expect to be a less dusty to a more dusty sample. The “red” and “blue” samples refer to red and blue SNe, having a colour parameter above or below  $c = -0.025$ , respectively. The “All” sample refers to what was shown in the top row of this figure. Note that in each case the blue (presumably less dusty) sample yields a step which does not appear to be significant, while the red (presumably more dusty) sample always yields a significant step. We define as significant a step size that it at least  $3\sigma$  away from 0. The errorbars in the plots are  $1\sigma$ . A version of this plot with more examples using different SFHs is available in Appendix C.

our dust measurements from the host galaxy do not use any information about the SN, and are still found to correlate with the SN Hubble residuals.

#### 5.4 Impact of Host Dust Corrections on the ‘Mass Step’

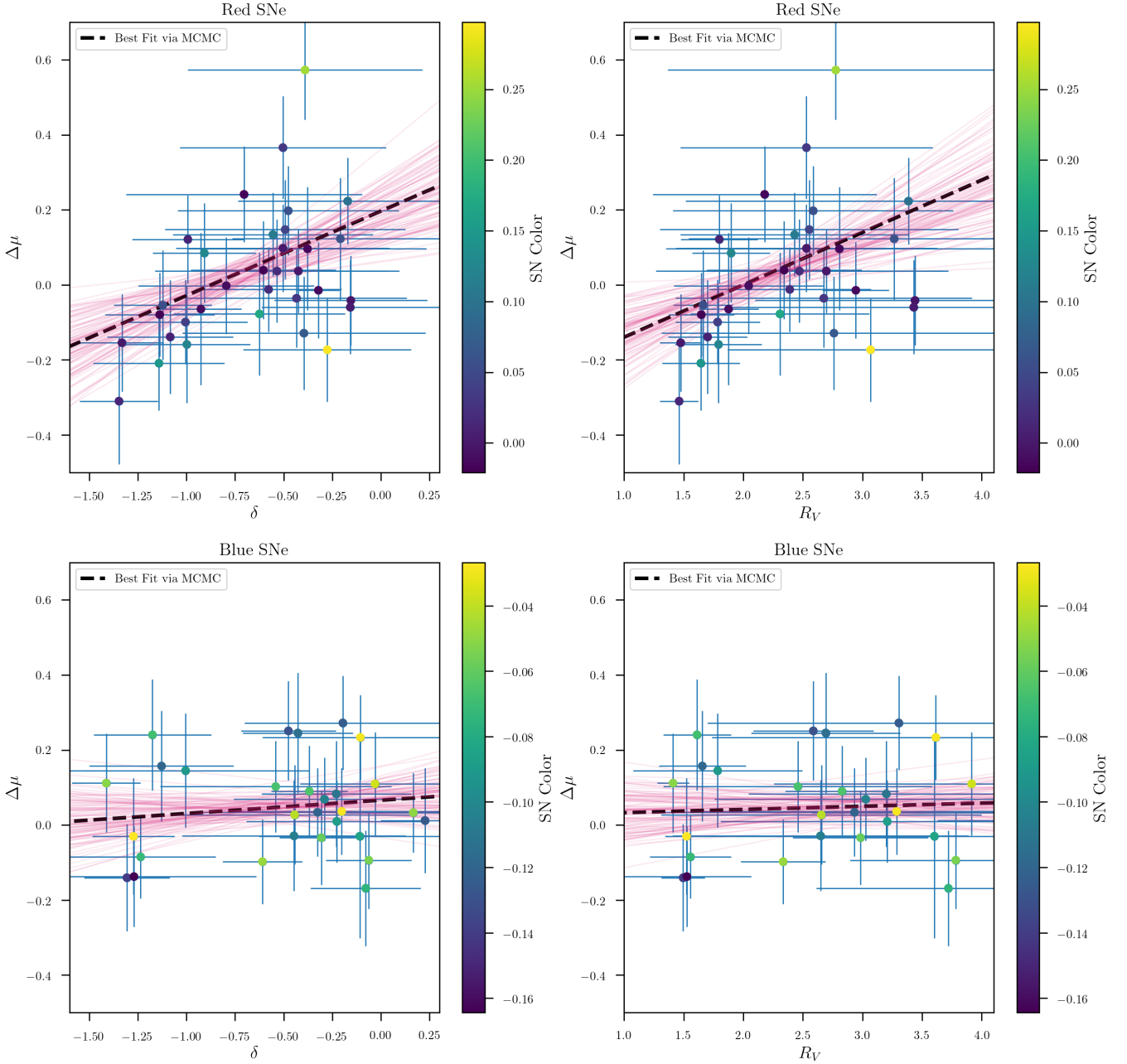
Finally, we apply fits to the observed correlations between the dust parameters and the Hubble residuals as corrections to the SN Ia luminosities, to determine if accounting for host-galaxy dust can explain the mass step. In the top left panel of Figure 8, we show uncorrected Hubble Residuals of 33 red DES SNe Ia (taken from the total 81 DES SNe Ia after their host galaxy fits are subjected to a  $\chi^2$  cut and a cut on color) versus the stellar mass of their host galaxies. By scanning from low to high mass, we determine that the highest significance mass step occurs at  $\log_{10}(M_*/M_\odot) = 9.8$  with a significance of  $3.76\sigma$ . The RMS scatter in the data is 0.17. The darker blue points represent the mean  $\Delta\mu$  value on either side of the mass step split.

After fitting the  $\Delta\mu$ - $\delta$  relation as described in Section , we fit a second order polynomial to the  $\Delta\mu$ - $A_V$  data that are

shown in Fig. 6. We then subtract the recovered relations out of the  $\Delta\mu$ ’s based on each galaxy’s  $A_V$  and refit the mass step. We choose a second order polynomial instead of a linear relation for  $A_V$  because the data shows a parabolic shape, and because the Akaike Information Criterion (AIC) is improved compared to the linear regression. The best fit relation is:  $\Delta\mu = 0.333A_V^2 - 0.657A_V + 0.253$ . Subtracting out the  $\Delta\mu$  value based on the best-fit relation and its  $A_V$  value results in the red points in the top-right plot of Figure 8. Remarkably, this procedure renders the mass step much less significant: the difference between the means on either side of the mass step split after the  $A_V$ - $\Delta\mu$  correction drops to a significance of only  $1.4\sigma$  (black points). Furthermore, the scatter decreases from 0.173 mag to 0.152 mag, meaning that the decrease in the mass step significance is not due to an increase in scatter. The bottom left panel of Figure 8 shows the results of subtracting out a linear  $\delta$  correlation shown in Figure 7. Again, both mass-step size and RMS scatter in  $\Delta\mu$  are reduced to  $2.7\sigma$  and 0.153 mag respectively.

Finally, we fit a 2-D linear correlation in  $A_V$  and  $\delta$  to the Hubble Residuals simultaneously and find  $\Delta\mu = -0.075A_V + 0.197\delta + 0.184$ . This provides the smallest scatter (0.150 mag)



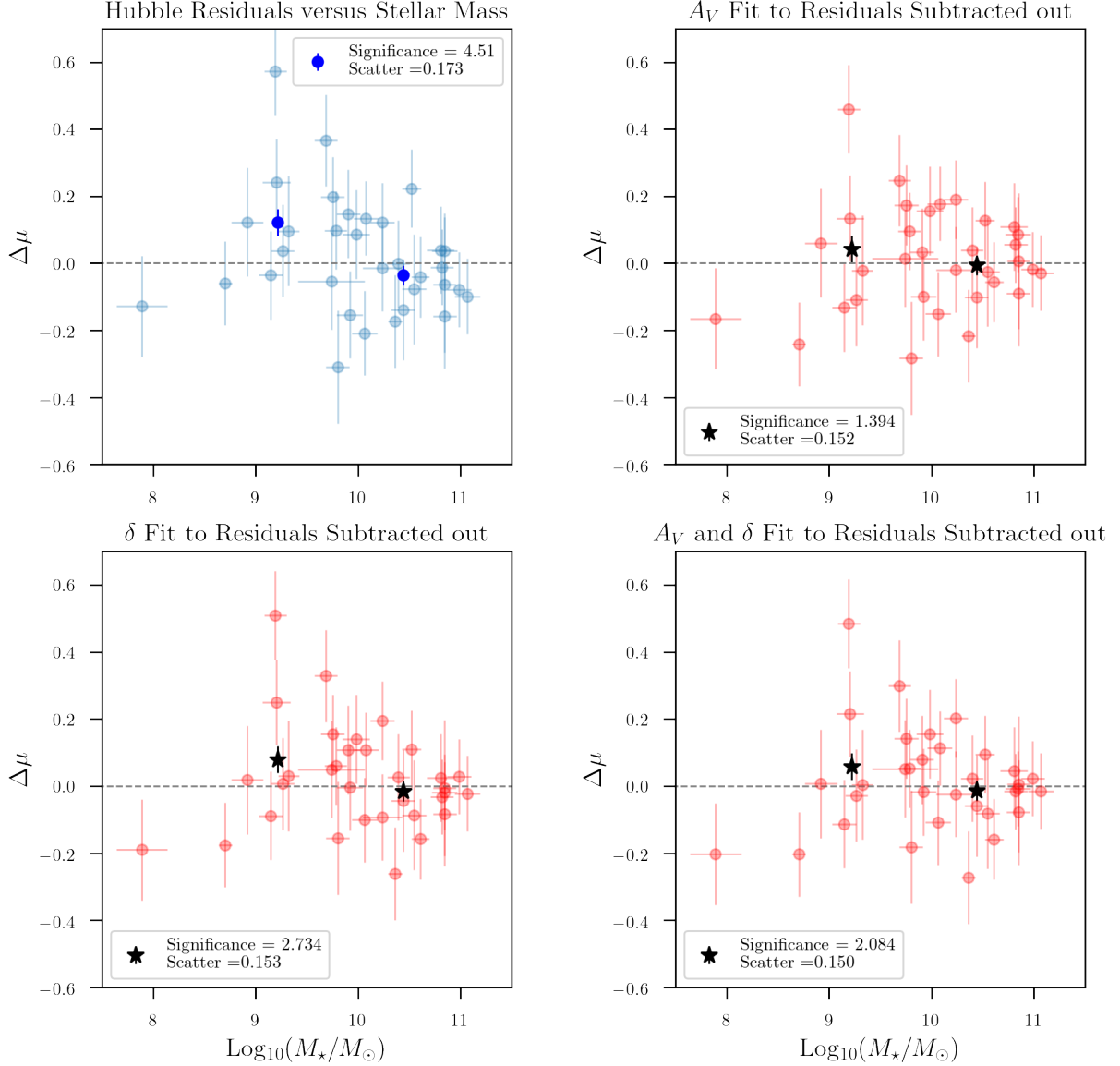


**Figure 7.** *Top:* Correlation between the host galaxy dust parameters  $\delta$  (left) and  $R_V$  (right) and Hubble Residuals ( $\Delta\mu$ ) for red (and presumably reddened by dust) SNe with SALT2  $c > -0.025$ , with the best fit linear relation (dashed line). The thinner lines represent an MCMC sampling of the posteriors for the slope and intercept of the relation. The slope is  $m = 0.219^{+0.076}_{-0.069}$  for the relation with  $\delta$  (which is  $\gtrsim 3\sigma$  away from a flat relation) while the  $R_V$  relation has a slope of  $m = 0.122^{+0.056}_{-0.046}$ . The colour of the data points represents the colour of the SN. *Bottom:* same as the top panels, but for the blue SNe with SALT2  $c < -0.025$ , which are less likely to be affected by dust than their redder counterparts. The trend between  $\delta$  or  $R_V$  and  $\Delta\mu$  vanishes, giving random scatter around zero with  $m = 0.008^{+0.024}_{-0.024}$  for the  $R_V$  case and  $m = 0.033^{+0.052}_{-0.051}$  for the  $\delta$  case.

amongst the three correlations tested. Step size is also dramatically reduced, down to  $\sim 2.1\sigma$ .

Salim & Narayanan (2020) already noted the stellar mass dependence of the attenuation law slope (i.e., similarly for  $\delta$  and  $R_V$ ) is driven by the observed relation between slope and  $A_V$ . In other words, the stellar mass correlation with dust parameters is driven by  $A_V$ . We interpret our findings in light of Salim et al. (2018); Salim & Narayanan (2020): the Hubble

residuals dependence on  $A_V$  is most effective in reducing the mass step because  $A_V$  drives the correlation between dust law and stellar mass. Evidently, there is a connection between the dust parameters of the host galaxy and the mass step. These results offer a compelling physical mechanism for understanding the mass step and also show that host galaxy dust properties have the potential to be used to reduce the scatter in the Hubble diagram.



**Figure 8.** Mass step for red SNe Ia in our sample (color  $c > -0.025$ ), with corrections that use host galaxy dust information. *Top left:* unaltered Hubble Residuals as a function of stellar mass for red SNe. The largest step occurs at  $\log_{10}(M_*/M_\odot) = 9.8$ , with a significance of  $4.5\sigma$ . The darker blue points show the mean of the data on each side of the split. *Top right:* we fit a 2nd-order polynomial to the Hubble Residuals as a function of  $A_V$ , and subtract this relationship out for each individual galaxy based on its measured  $A_V$ . This change reduces the significance of the mass step down to  $1.4\sigma$ , as it can be seen from the black star markers, showing the new mean values on either side of the split, while also decreasing the scatter to 0.15. *Bottom left:* another correction to the mass step, this time subtracting out the linear relation fit between Hubble residuals and  $\delta$ . This correction also reduces step size and scatter, down to  $2.7\sigma$  and 0.15 respectively. *Bottom right:* we fit the Hubble residuals as a function of both  $A_V$  and  $\delta$ , and correct the mass step plot with its best fit. This correction provides the most interesting results, with the step being reduced to  $2\sigma$  and the scatter being reduced to 0.15. We claim that the significant scatter reduction coupled with mass step decrease points to an interesting connection between the mass step and dust parameters, and to promising applications of host galaxy dust estimates to reduce the scatter in the Hubble diagram.

### 5.5 Challenges with measurements of dust parameters from broadband photometry

It is well known that there exist degeneracies between several galaxy parameters, especially as measured from optical/NIR broadband photometry (e.g., Qin et al. 2022). Because of the way the dust parameters explored here are defined, it is obvious that there will be a certain degree of degeneracy between  $\delta$  and  $A_V$ . SFR is another parameter that is degen-

erate with the dust parameter, as a red and quenched galaxy with little dust content may look very similar to a star forming galaxy with a steep dust attenuation law in optical/NIR photometry. A galaxy sample such as the one presented by Salim et al. (2018); Salim & Narayanan (2020) has exquisite photometry throughout a wide range of filters, which makes it possible to break those parameter degeneracies, especially with the inclusion of longer wavelength, IR bands. Their robust analysis shows that the correlations found between  $\delta$

and  $A_V$  are not an artifact. Our results are widely in agreement with their findings, despite the very different datasets and techniques used, which supports the conclusion that this work is also not retrieving some spurious relations. To further test this conclusion, we test our method on a state-of-the-art panchromatic extragalactic catalog, the COSMOS catalog from [Weaver et al. \(2022\)](#).

Our test consists in mapping our SN hosts onto similar galaxies in the COSMOS catalog, and then comparing the derived galaxy properties of the two samples. By doing so, we hope to understand whether similar trends are found at the population level when a similar galaxy sample, with significantly improved photometric data including IR bands, is analyzed. First, for each SN host, we find the top 2 nearest neighbours in the DES COSMOS data based on their magnitude, in redshift bins. We then find the COSMOS data from [Weaver et al. \(2022\)](#) that matches the DES COSMOS galaxies, and run BAGPIPES using all of the available COSMOS data. We assume the COSMOS photometric redshifts as priors in the BAGPIPES run. Because the COSMOS field is only  $2 \text{ deg}^2$ , we do not expect to find appropriate neighbour matches for all galaxies, especially those at the lowest redshifts, so some differences between the SN hosts and the COSMOS populations are expected. We are, however, able to recover with the COSMOS neighbours population: 1) two significantly different distributions for low versus high mass galaxies for the dust parameters shown in Fig. 3; 2) similar trends for  $A_V$  and  $\delta$  versus mass as those in Section 4; 3) Hubble residuals steps and trends similar to those in 3, although at a lower significance. We therefore conclude that, although there may be degeneracies amongst dust and star formation parameters due to the limited photometric bands available, the trends are confirmed when using data from one of the richest photometric galaxy catalogs available, and are likely not spurious. In other words, although the dust parameters may not be well constrained for each single galaxy, as indicated by the large uncertainties in e.g., Fig. 3, the overall trends found for the high versus low mass galaxies and for the step function are not expected to be artifacts.

## 6 DISCUSSION AND CONCLUSIONS

In this work, we have presented a detailed study of DES SN hosts, posing particular attention to the estimation of dust parameters and to correlations of dust parameters with SN Ia Hubble diagram residuals. While some host galaxy properties, such as dust, star formation histories, and stellar mass, have been recently explored with simulations or control samples, this is the first work exploring the impact of dust estimates from host photometry alone on the SN Ia mass step. We detect significant correlations between global host galaxy dust parameters and Hubble residuals for the first time. We show how these correlations are able to account for the mass step to a large extent, and how host galaxy dust parameters can be used to reduce the scatter in the SNe Ia Hubble diagram.

We show in Table 5 the predictions from SNe Ia and compare them with the findings for host dust properties; all of which are consistent with the SN Ia color-luminosity relation, SN Ia intrinsic scatter, and SN Ia ‘mass step’ being driven by host-dust. The following points summarize our findings:

- There is significant ( $\gtrsim 4\sigma$ ) step-like behaviour in the Hubble residuals as a function of dust parameters  $A_V$  and  $\delta$  (and consequently,  $R_V$ ). These step-like relations are driven by the red SNe (with color parameter  $c > -0.025$ ), while they wash out for the bluer SNe. This finding, together with the fact that the step is as strong or stronger than the mass step, could be indicating that these relations are due to the presence of dust, rather than being a simple consequence of the underlying correlation between host galaxy mass and dust.
- There exists a linear correlation, significant at  $3\sigma$ , between the Hubble residual and the dust parameter  $\delta$  when one removes SNe where dust attenuation is negligible, i.e. the bluer SNe with color parameter  $c < -0.025$ .
- Correlations of host galaxy dust parameters with the Hubble residuals can account for the size of the mass step and simultaneously reduce scatter in red SNe, indicating that the origin of the mass step could be in the host galaxy dust. Due to the small sample size, this study is unable to confirm or rule out the presence of other mechanisms that could be partially responsible for the mass step, along with dust. Future analyses with larger data sets should be able to provide a more quantitative statement on the extent to which the mass step can be explained by dust alone.
- Host galaxy dust parameters can be used to reduce the scatter in the Hubble diagram by up to 13%, thus potentially improving SN cosmology results.
- We find evidence for two different populations of  $\delta$  (or equivalently,  $R_V$ ) when separating the sample into low and high mass galaxies, following a typical mass step split at  $10^{10} M_\odot$ . These two populations show a similar behaviour to what [Brout & Scolnic \(2021\)](#) predicted would be observed if dust were to explain the observed mass step: lower mass galaxies tend to have a larger  $\delta$  (or equivalently,  $R_V$ ).
- These results for  $\delta$  are similar to what is found in a generic galaxy sample. Moreover, SN Ia hosts show a strong correlation between the dust attenuation in V band  $A_V$  and stellar mass, which is also consistent with a generic galaxy population.
- We find that the  $A_V$  derived from the integrated galaxy photometry is not strongly correlated with the amount of dust along the line of sight to the SNe (SALT2 color). It is important to note that host-galaxy attenuation and its wavelength dependence ( $R_V$ ) in this work have been derived from the non-local galaxy photometry, yet we compare our results with the dust extinction and properties derived from line of sight photometry to the SNe. Our findings agree with the general expectation that while the amount of dust can vary greatly across a galaxy, the observed SN color is highly dependent on its location and line of sight within the galaxy. This is also consistent with findings of sibling SNe in the same galaxy having large variation in SN colors ([Scolnic et al. 2020](#); [Biswas et al. 2021](#); [Scolnic et al. 2021](#)). However, we do find a significant correlation between the  $R_V$  (or  $\beta$ ) found from the SNe and that from the attenuation of the whole galaxy. We interpret this as a consequence of the fact that the size and properties of the dust grains themselves, which are expected to drive  $R_V$  ([Salim et al. 2018](#)), are similar across the galaxy and any single location or line of sight to the SN is a representative probe.
- SN host galaxies’  $\delta$  and  $R_V$  span a wide range of values: while lower mass galaxies tend to have  $-0.5 < \delta < 0.1$  or  $2 < R_V < 4$  (68% CI of the distribution), a significant (of

Dust Prediction from SN Ia	Reference	Global Galaxy Dust Finding	Reference
$R_V$ ranges from $\sim 1$ -6	BS21	$R_V$ ranges from $\sim 1$ -6	Fig. 3 (this work) and Salim & Narayanan (2020)
High mass galaxies mean $R_V$ is $\sim 0.5$ lower than low mass galaxies	BS21	High mass galaxies mean $R_V$ is $\sim 0.7$ lower than low mass galaxies; high mass sample has two $R_V$ populations	Fig. 3 (this work) and Salim & Narayanan (2020)
The mass step is driven by dust differences	BS21	Correcting for galaxy dust properties reduces mass step significance to $2\sigma$	Fig. 5 and 8 (this work)
Host color - SN Ia Hubble Residual step is as large or larger than mass step	Roman et al. (2018)	Host color is strongly correlated with host $R_V$	Fig. 4 (this work)
The SN Ia color-luminosity relation ( $\beta$ ) is largely driven by $R_V$	BS21	Red-sequence galaxies have very low $R_V$ , and measurement of $\beta$ for same sample is very low	Fig. 5 (this work) and Chen et al. (2022)
SNe Ia in red-sequence galaxies have less HR scatter than SNe Ia in other galaxies	Chen et al. (2022)	Red sequence galaxies have smaller range of $R_V$ than other galaxies	Fig. 5 (this work)

**Table 5.** Comparison of the findings from galaxies studies on dust, to relevant SNIa predictions. HR stands for ‘Hubble Residual’.

the order of 10%) fraction of galaxies with mass  $> 10^{10} M_\odot$  reach below  $\delta \sim -0.9$  or  $R_V < 2$ . Our findings suggest that assuming a single value or a very narrow distribution for  $R_V$ , as often done in SN cosmology, may not be accurate.

A wider area, wider wavelength coverage (including more infrared bands than VISTA), and deeper photometric data set than what used here is under preparation for the final DES deep-field catalogs that will be used for the final DES cosmology analyses. This catalog will cover 3 times as many SN Ia hosts as those used in this work, and it will enable more precise dust parameter estimation for each galaxy. As such, this catalog will be a valuable dataset to replicate this analysis and confirm with more precision and confidence the results of this work. We will also be able to extend the analyses to use the Hubble residuals for the photometric sample of Möller et al. (2022), including more than 1000 SNe, as opposed to the 81 spectroscopically classified SNe used in the last part of this work. Precision spectroscopy of host galaxies in the optical to infrared, and redder bands (at  $\gtrsim 8 \mu\text{m}$  in rest frame), will also be valuable to derive more precise dust properties constraints in SN hosts and therefore strengthen the argument that the observed mass step is a consequence of dust properties in SN host galaxies. Measurements of this kind will also be useful to improve measurements of cosmological parameters from the SN Ia Hubble diagram, since we have shown that it is possible to use dust parameters information to improve the standardization of SN Ia. We conclude that a better understanding of the origin of the mass step and its correlation with dust may be crucial to obtain more precise and accurate cosmological constraints from SN cosmology.

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The Collaborating Institutions are Argonne National Lab-



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## APPENDIX A: COMPARISON OF HOSTS PHOTOMETRY

In this section we compare the host galaxy properties computed with two different sets of photometric data, as shown in Figure A1: that of the DES deep fields of [Hartley & Choi et al., \(2020\)](#) (on the  $x$  axis of each panel), and that from the catalog of [Wiseman et al. \(2020\)](#) (on the  $y$  axis of each panel). We remove objects that are brighter in the deep fields catalog (i.e. likely contaminated by the SN light) by more than twice the standard deviation of the residual magnitude distribution from the comparison of the two photometric datasets in any band. Switching between photometric data does not introduce a large amount of variation: only 1, 2, and 5% of galaxies had their  $A_V$ ,  $R_V$  or stellar mass respectively changed by more than  $2\sigma$  (magenta and orange points in Figure A1). We choose the SN Stack data in our fiducial dataset throughout the paper as it removes contamination due to the super-

nova that contaminates light from the host galaxy. While it is not obvious which threshold on the outlier fraction should be used in order to consider the change problematic for this work, we have made the entire analysis in this work using the photometry from the DES deep fields, and did not find any significant change in our conclusions. Moreover, the  $\chi^2$  cut we apply on the BAGPIPES outputs results in a reduction in the number of galaxies in the deepfields case compared to the W20 catalog because of the contaminated photometry, so that the W20 also provides better fits to the galaxies SEDs.

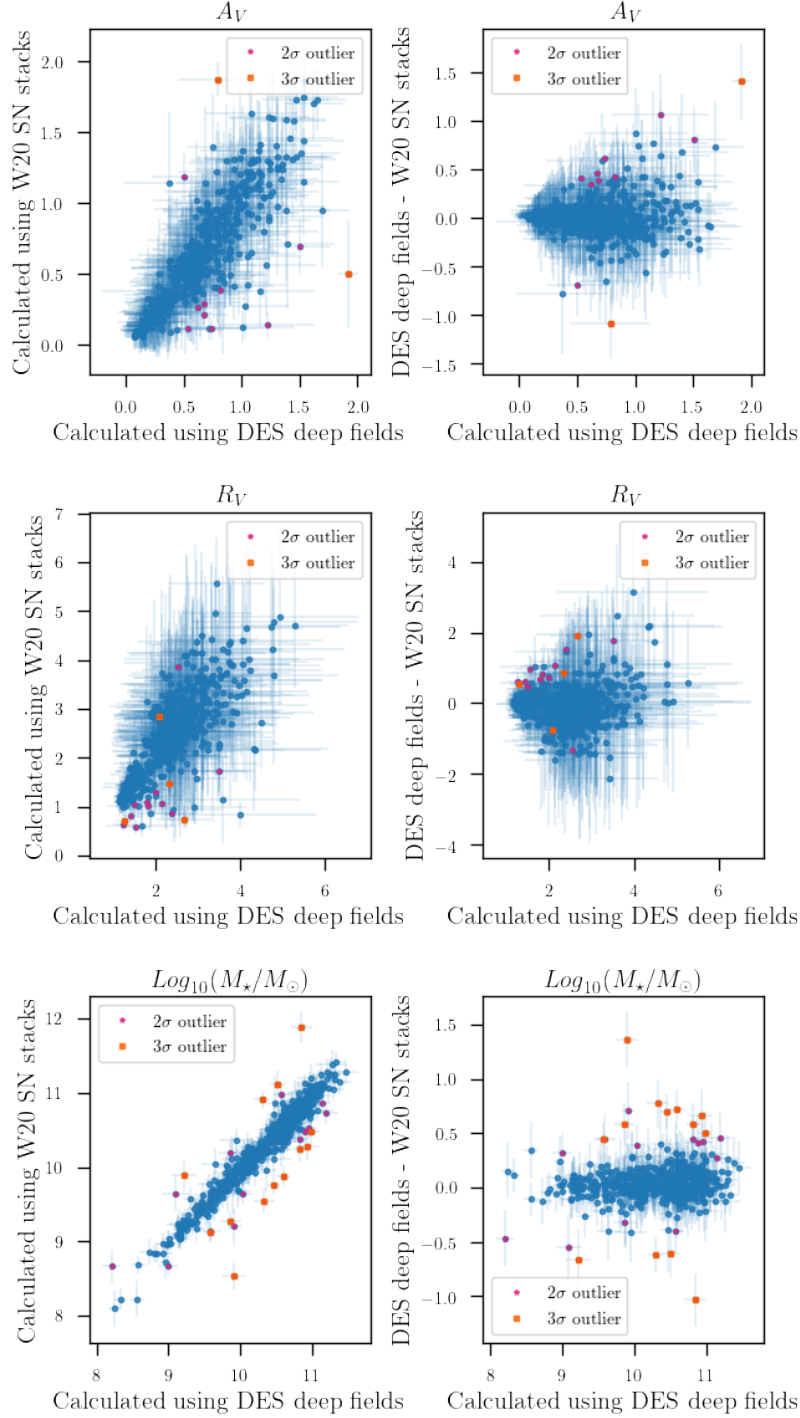
## APPENDIX B: STELLAR MASS - DUST PARAMETER RELATIONS FOR AN EXPONENTIAL STAR FORMATION HISTORY

While in Figure 3 we assume a log-normal SFH, Figure B1 contains similar plots for an exponential SFH with a prior of  $\tau < 2$  Gyr. In this figure, it is clear that a different assumption in the SFH makes the correlation between the dust parameters and the stellar mass less obvious, at least when splitting the sample at  $10^{10} M_\odot$ . We interpret this result as a consequence of the fact that there exists a degeneracy between the dust, the SFH, and the stellar mass of galaxies when their SEDs are fit using the available photometric bands. We note that, however, as soon as the prior on  $\tau$  is relaxed to a more realistic one (e.g.  $\tau < 10$  Gyr), the  $A_V/\delta$  versus mass plots resemble more closely our findings in Figure 3. We conclude that the main findings of this work are not largely affected by the exact choice of the SFH parameterisation, as long as the prior choices are realistic. We show here the effect of a more extreme and unrealistic SFH choice (although used in the literature, e.g. S20) on the recovered parameters, when only more peaked SFHs are considered (because of the low  $\tau$  values allowed).

## APPENDIX C: HUBBLE RESIDUALS AND DUST PARAMETERS WITH DIFFERENT HOST GALAXY SED FITTING PRIORS

In this Appendix we study the correlations between Hubble residuals and various dust parameters assuming different SFHs and dust parameters priors. Similarly to Figure 6, Figure C1 shows the step functions for host galaxy parameters computed with different assumptions: an exponential SFH with  $0 < \tau < 2$  Gyr, with  $R_V$  both free and fixed, an exponential SFH with  $0 < \tau < 10$  Gyr, with  $R_V$  free and fixed, and a log-normal SFH with  $R_V$  left free. The form of the SFH has a noticeable effect on the significance of the steps in the Hubble Residuals. We find a significant step with respect to mass in every SFH, but for those where  $\tau$  is restricted to  $[0, 2]$  Gyr there is no significant step in the dust parameters. In every other case where there is no such restriction on  $\tau$  there is a significant step in all dust parameters. This is important to note for the DES cosmology, as the  $0 < \tau < 2$  Gyr prior is the one used in S20, where the mass step is described in detail as resulting from the DES 3Yr data.

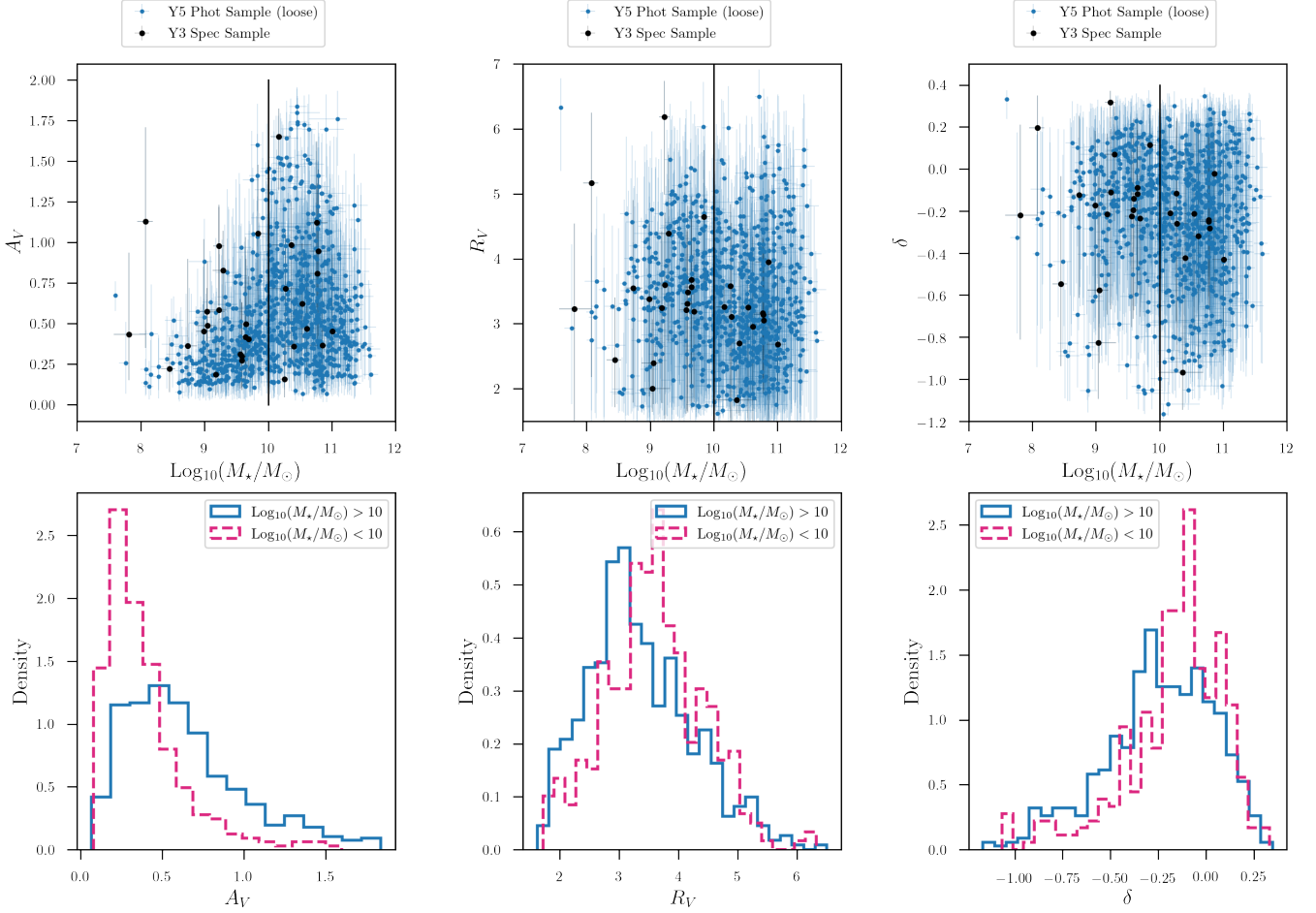
We also show the equivalent of the bottom panels in Figure 6 but for multiple SFH and parameter prior combinations in Figure C2. In this case, we show the step size, with its uncertainty, after splitting the sample into hosts of red (with color



**Figure A1.** Comparison of host galaxy properties computed with two different sets of photometric data: that of the DES deep fields of [Hartley & Choi et al., \(2020\)](#) (shown on the  $x$  axis of each panel), and that from the catalog of [Wiseman et al. \(2020\)](#) (shown on the  $y$  axis of each panel on the left hand side). The right-hand side shows, on the  $y$ -axis, the residuals given by the difference between the two estimations of the same parameter. We show results for the dust parameters  $A_V$  and  $R_V$ , and the stellar mass, as these are the relevant quantities for this work. The magenta and orange points represent the  $2\sigma$  and  $3\sigma$  outliers, respectively.

parameter  $c > -0.025$ , presumably in a dustier environment) and blue (with color parameter  $c < -0.025$ , presumably in a less dusty environment) SNe. There are a few important conclusions to be drawn from these plots. Firstly, all the plots that were considered significant in Figure C1 have sharp up-

ward trends in step size as we move to supposedly “dustier” SN environments. This is further evidence that the mass step has some dependence on dust parameters. Secondly, the cases that were not deemed significant do not experience this trend. We define as significant a step where the mean values on



**Figure B1.** Same as Figure 3, but with parameters derived assuming an exponential SFH. Top: Dust parameters  $A_V$  and  $R_V$  as a function of stellar mass, derived using an exponential SFH. Bottom: Distribution of  $A_V$  and  $R_V$  values above and below the  $10^{10} M_\odot$  split, the location of the observed mass step.

each point are at least  $3\sigma$  away from each other. The two non-significant cases in  $A_V$  (the leftmost column) lack the increase in step when going from a less dusty to a dustier environment on average, while the two non-significant cases in the two rightmost columns (they are actually the same, since they are  $\delta$  and  $R_V$  for the same case) increase very slightly but stay non-significant in all three samples. Since there is no theoretical or empirical reason to restrict  $\tau$  to  $< 2$  Gyr, we consider the results assuming the log-normal SFH (or similarly, the exponential SFH with  $0 < \tau < 10$  Gyr), as our fiducial results, also considering how the change in SFH parametrisation does not appear to have a noticeable effect on the step size or significance.

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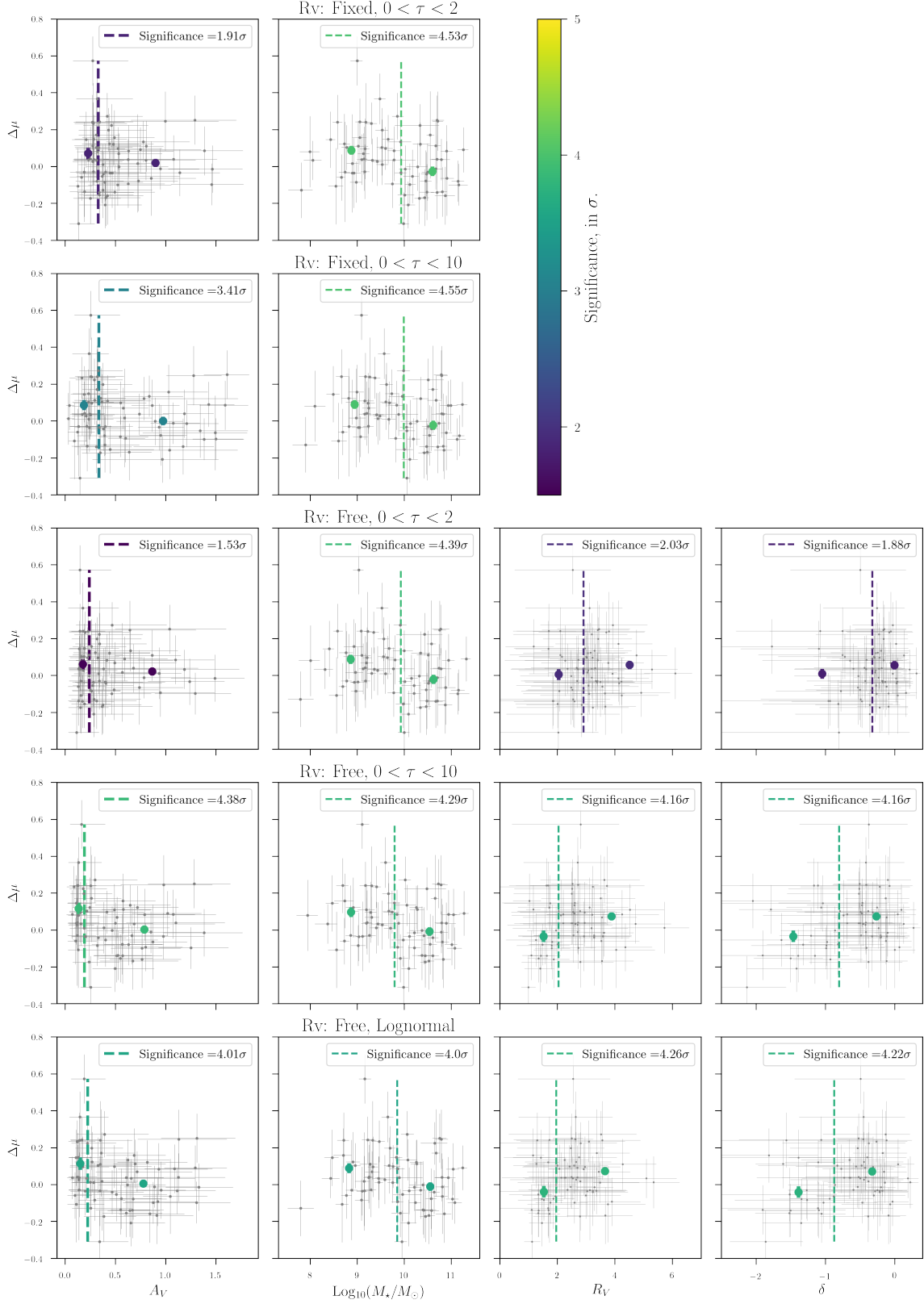
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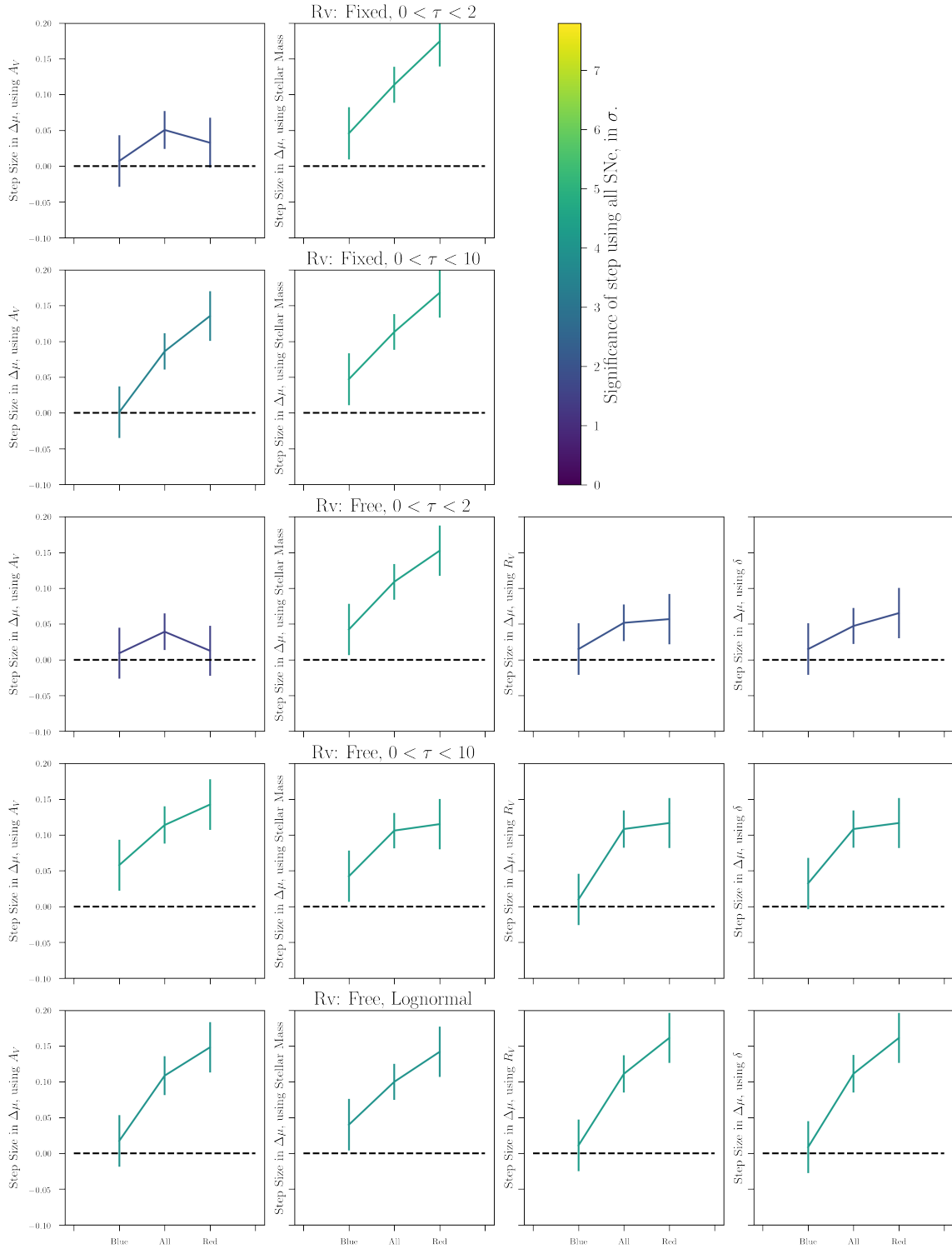
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**Figure C1.** Step-like behaviour in Hubble residuals as a function of  $A_V$ , stellar mass,  $R_V$  and  $\delta$ . These are calculated using five SFHs, four of which are exponential with  $0 < \tau < 2$  or  $0 < \tau < 10$  Gyr with  $R_V$  free or fixed, the final being log-normal with  $R_V$  left free. We then scan over each parameter to search for a significant step in the Hubble Residuals, i.e. a certain point where the mean of the Hubble residuals on either side of said point is significantly different. The most significant step found is displayed in each panel, along with its  $3\sigma$  errorbars, and is coloured according to its significance.





**Figure C2.** Step size in the Hubble Residuals as we move from what we expect to be a less dusty to a more dusty sample. The “red” and “blue” samples refer to red and blue SNe, having a colour above or below -0.025. The “All” sample refers to what was shown in Figure 6. As in Figure 6, the steps are coloured according to their significance where significance was determined by the step calculated from the SNe and host properties in each subsample.

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